

Feedback: Now With Physics!*

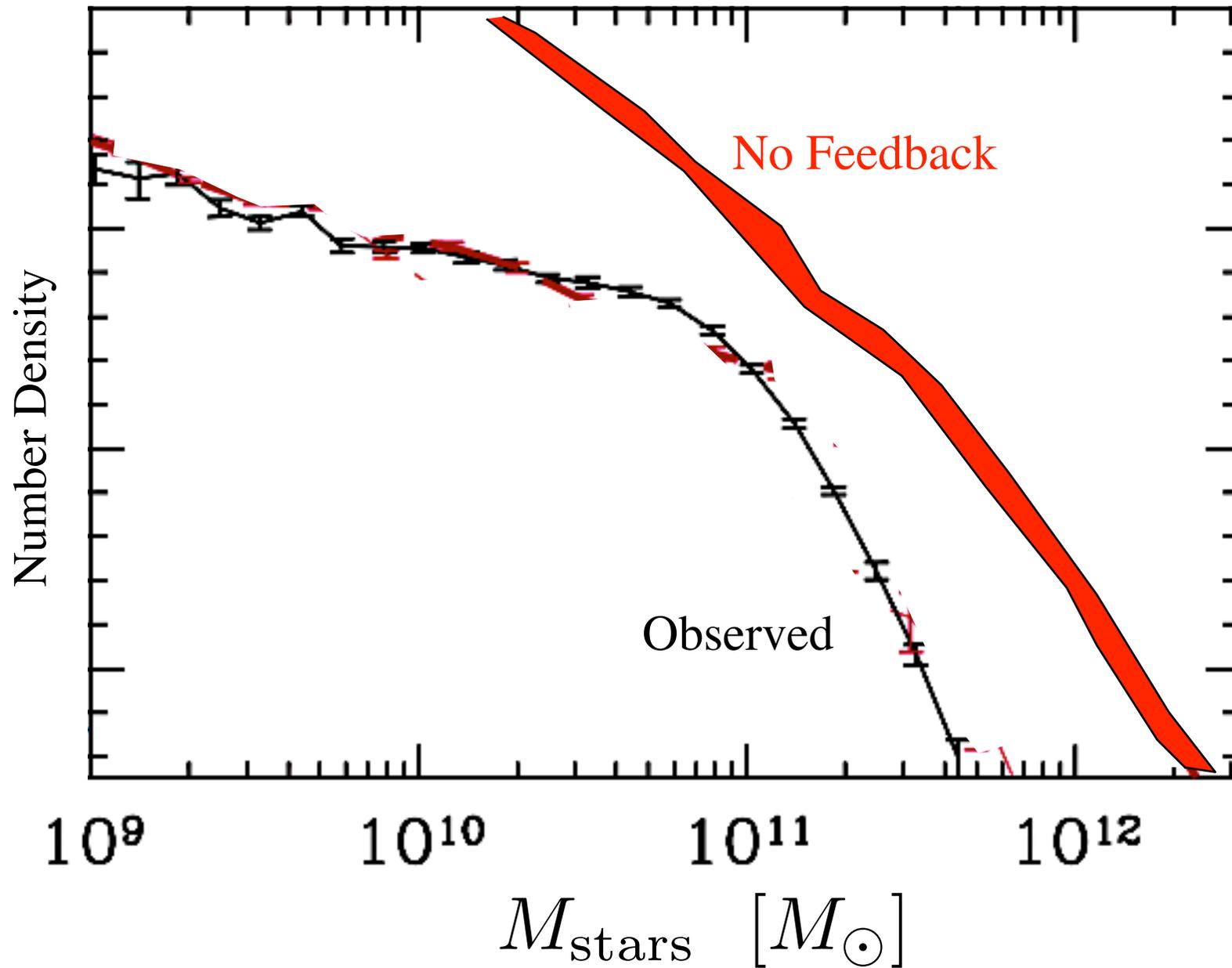
Dusan Keres

Philip Hopkins, Eliot Quataert, Norm Murray, Jose Onorbe

* *Real* physics not necessarily included

Motivation

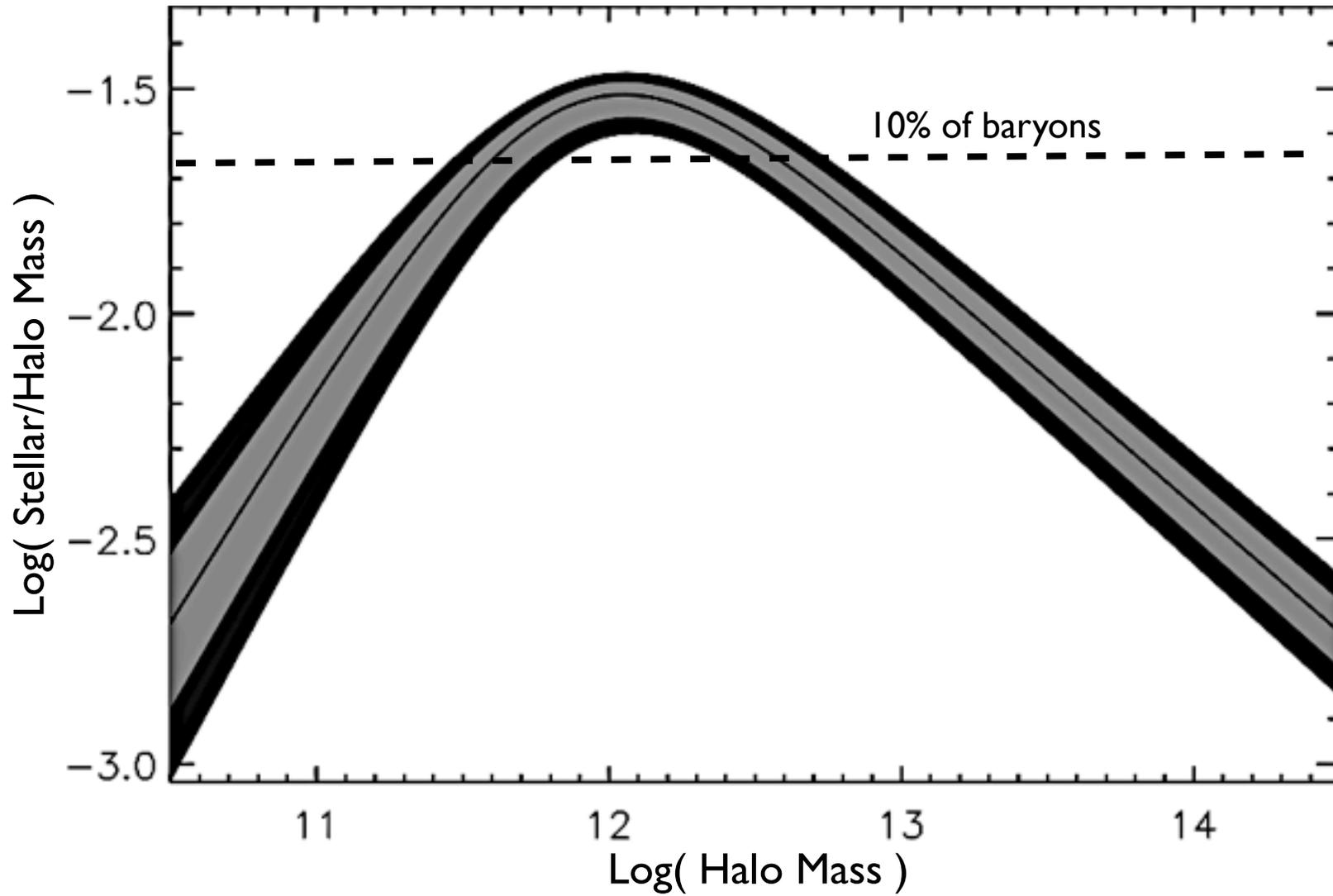
Q: WHY IS STAR FORMATION SO INEFFICIENT?



Motivation

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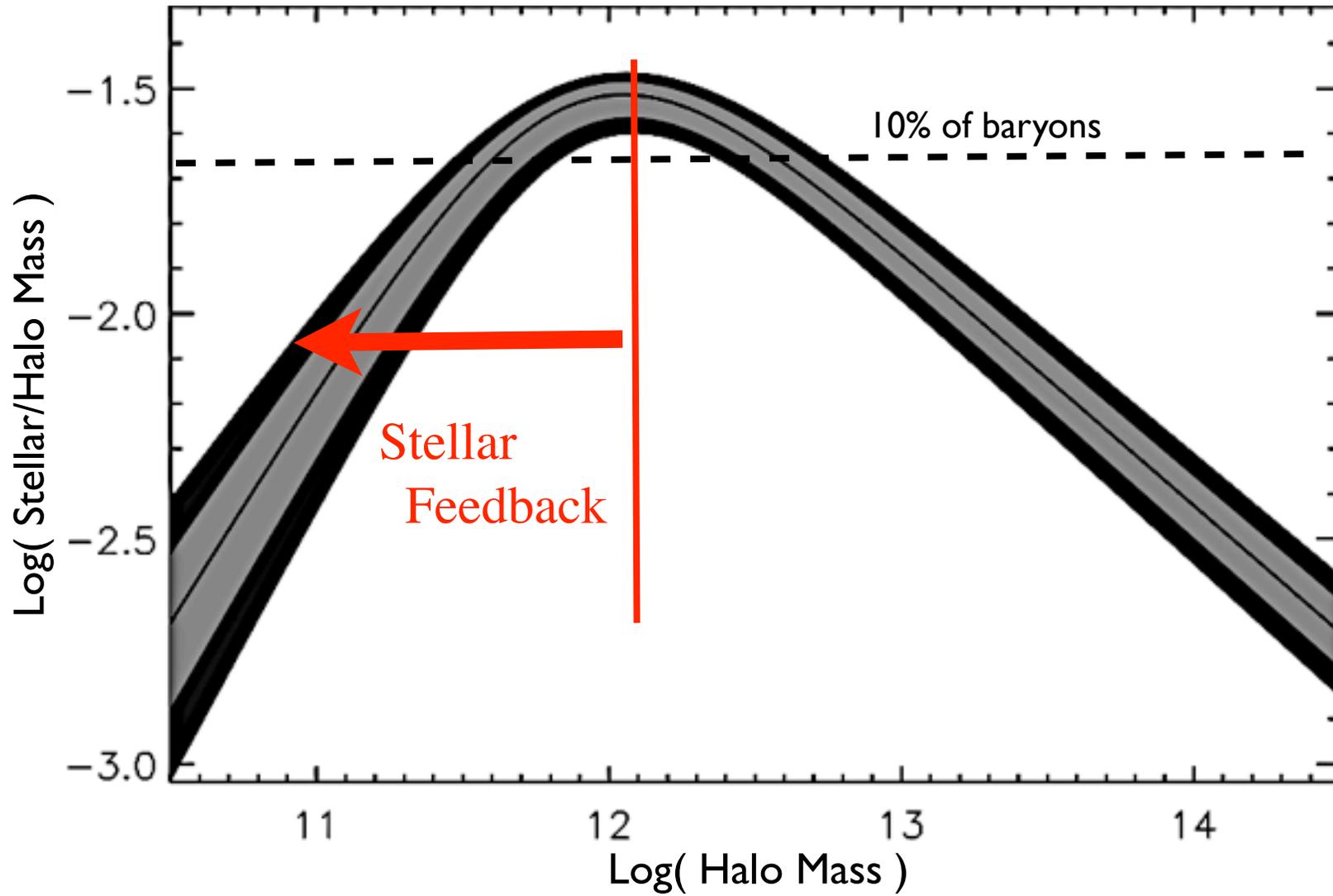
Moster 2009



Motivation

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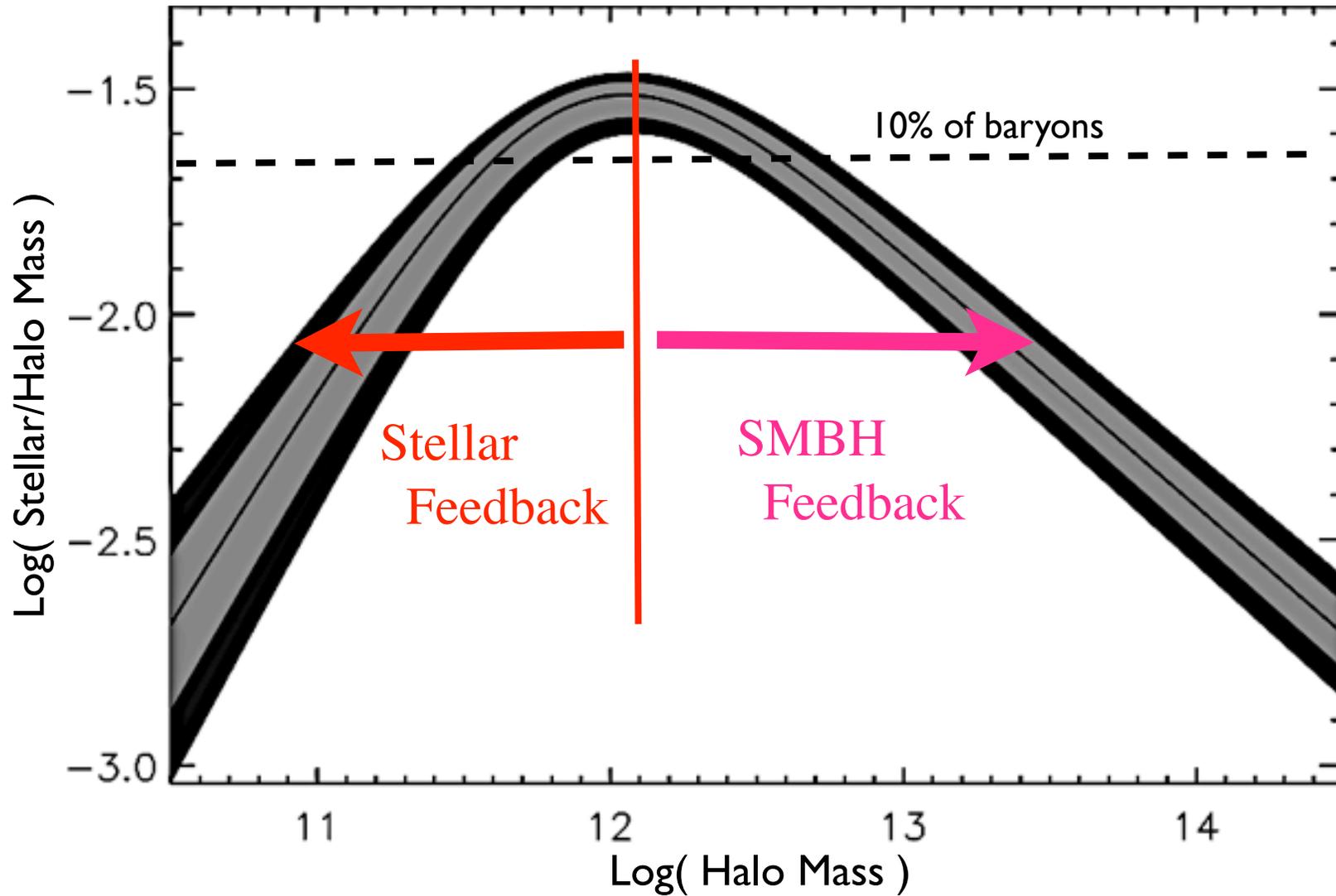
Moster 2009



Motivation

Q: WHY IS STAR FORMATION SO INEFFICIENT?

Moster 2009



Stellar Feedback is the Key!

SO WHAT'S THE PROBLEM?

- Standard (in Galaxy Formation):
Couple SNe ($\sim 1e51$ erg/SN)
as “heating”/thermal energy

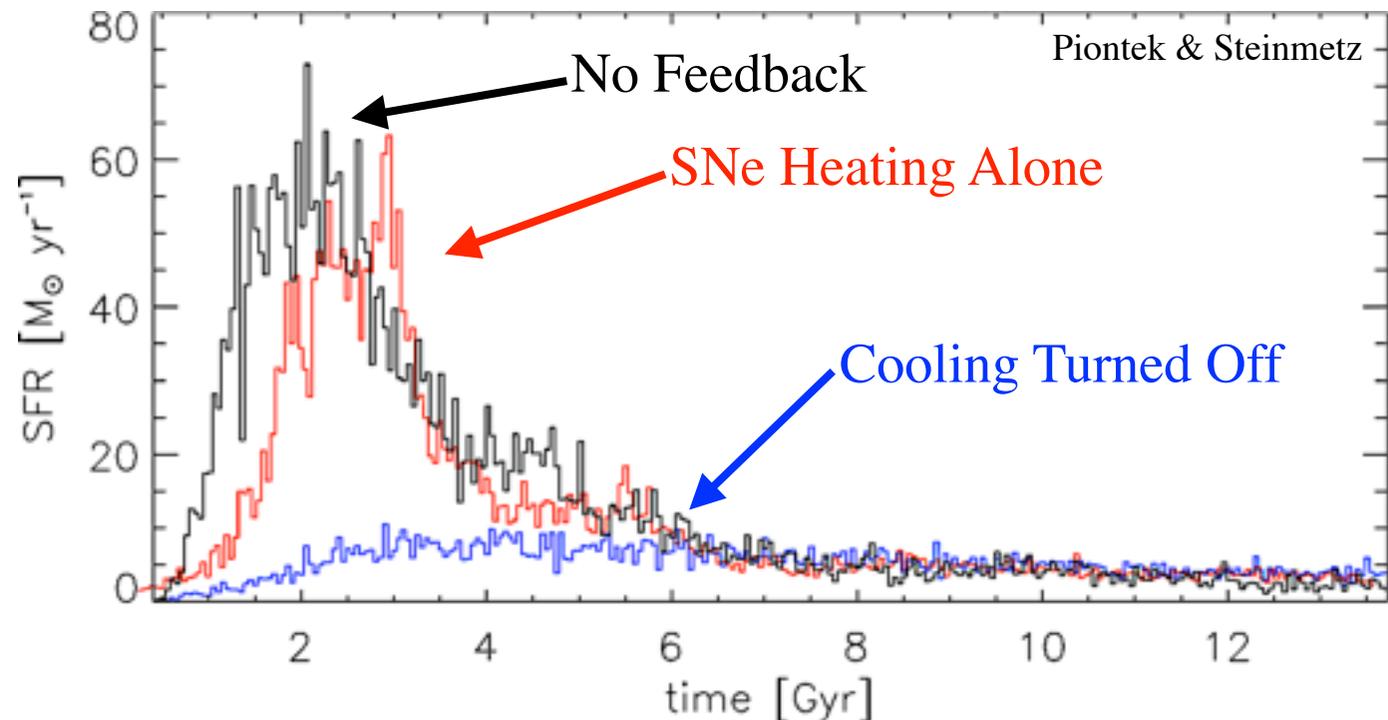
- FAILS:

$$t_{\text{cool}} \sim 4000 \text{ yr} \left(\frac{n}{\text{cm}^{-3}} \right)^{-1}$$

$$t_{\text{dyn}} \sim 10^8 \text{ yr} \left(\frac{n}{\text{cm}^{-3}} \right)^{-1/2}$$

- “Cheat”:

- Turn off cooling
- Force wind by hand
(‘kick’ out of galaxy)

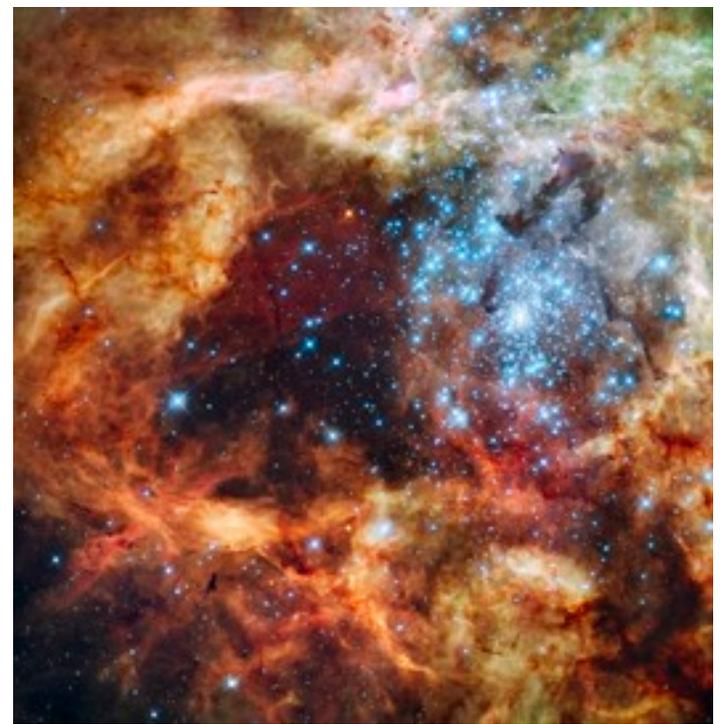


Stellar Feedback: Understanding the key Physics



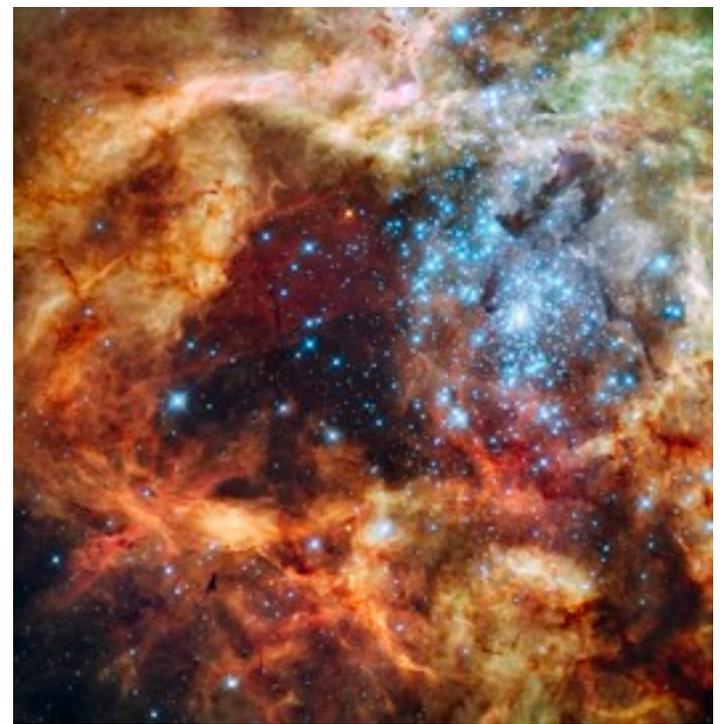
Stellar Feedback: Understanding the key Physics

- High-resolution ($\sim 1\text{pc}$), molecular cooling ($<100\text{ K}$), SF only at highest densities ($n_{\text{H}} > 1000\text{ cm}^{-3}$)



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- “Energy Injection”:
 - SNe (II & Ia)
 - Stellar Winds
 - Photoionization (HII) + Photoelectric



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- *Explicit* Momentum Flux:

- Radiation Pressure

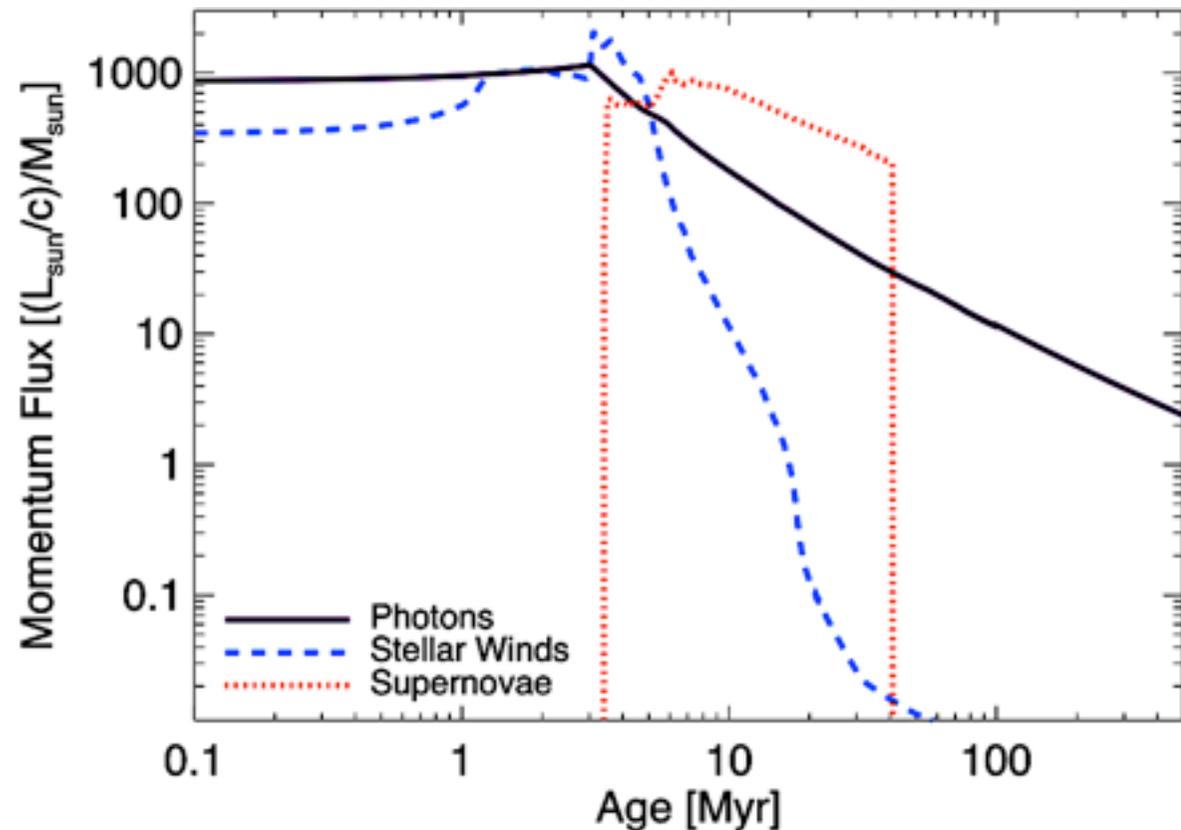
$$\dot{P}_{\text{rad}} \sim \frac{L}{c} (1 + \tau_{\text{IR}})$$

- SNe

$$\dot{P}_{\text{SNe}} \sim \dot{E}_{\text{SNe}} v_{\text{ejecta}}^{-1}$$

- Stellar Winds

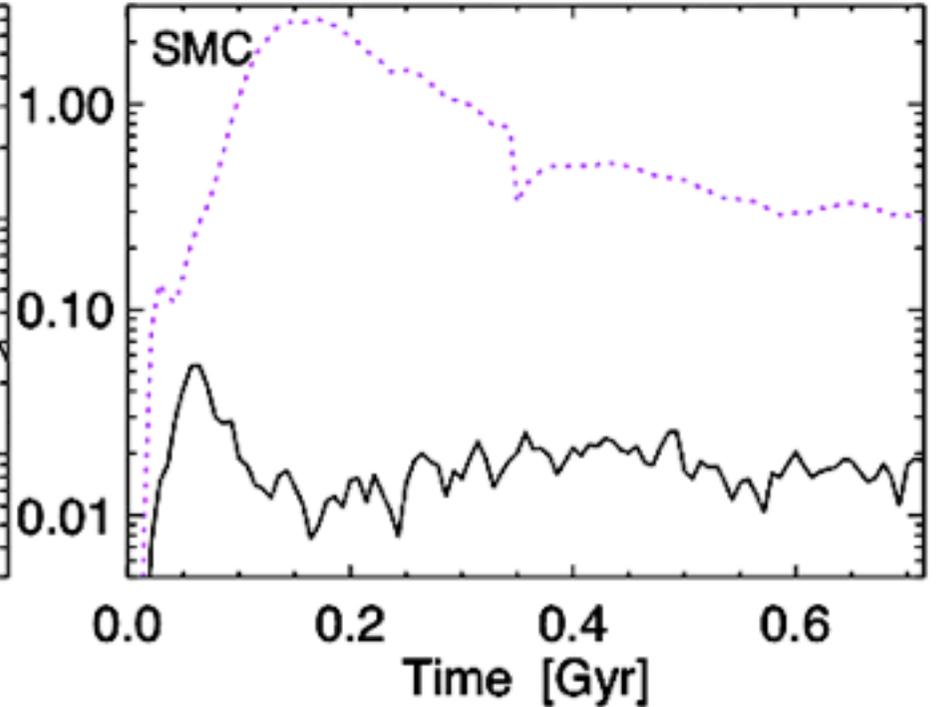
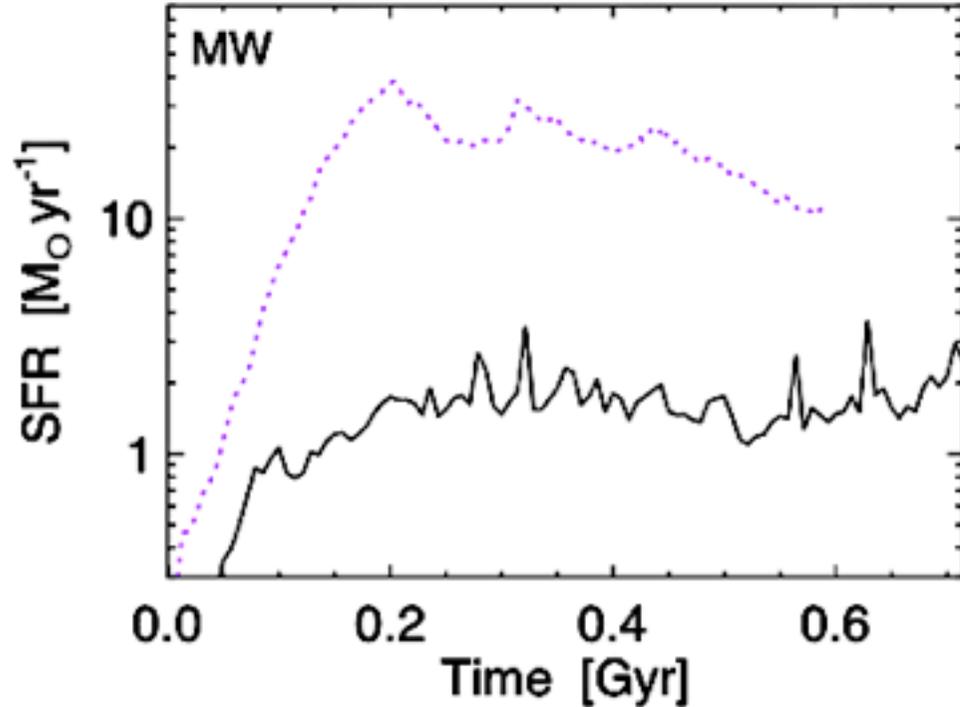
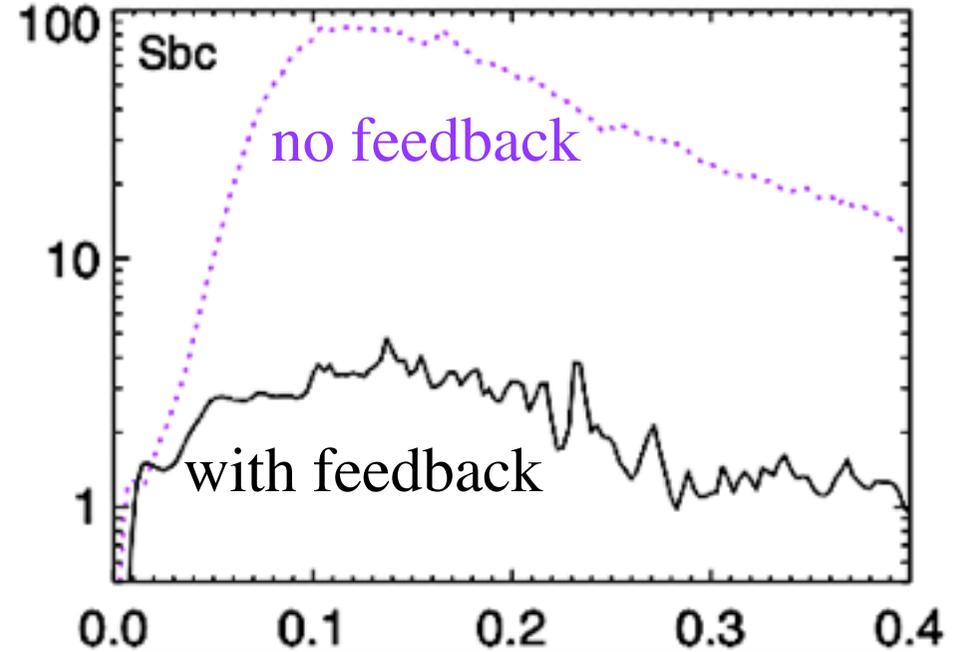
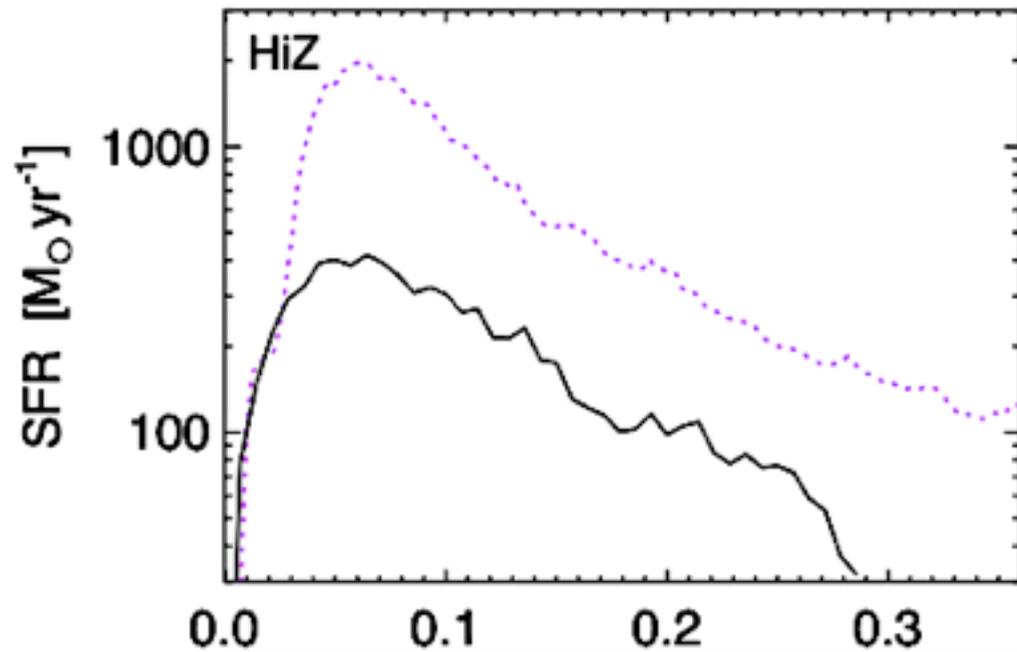
$$\dot{P}_{\text{W}} \sim \dot{M} v_{\text{wind}}$$



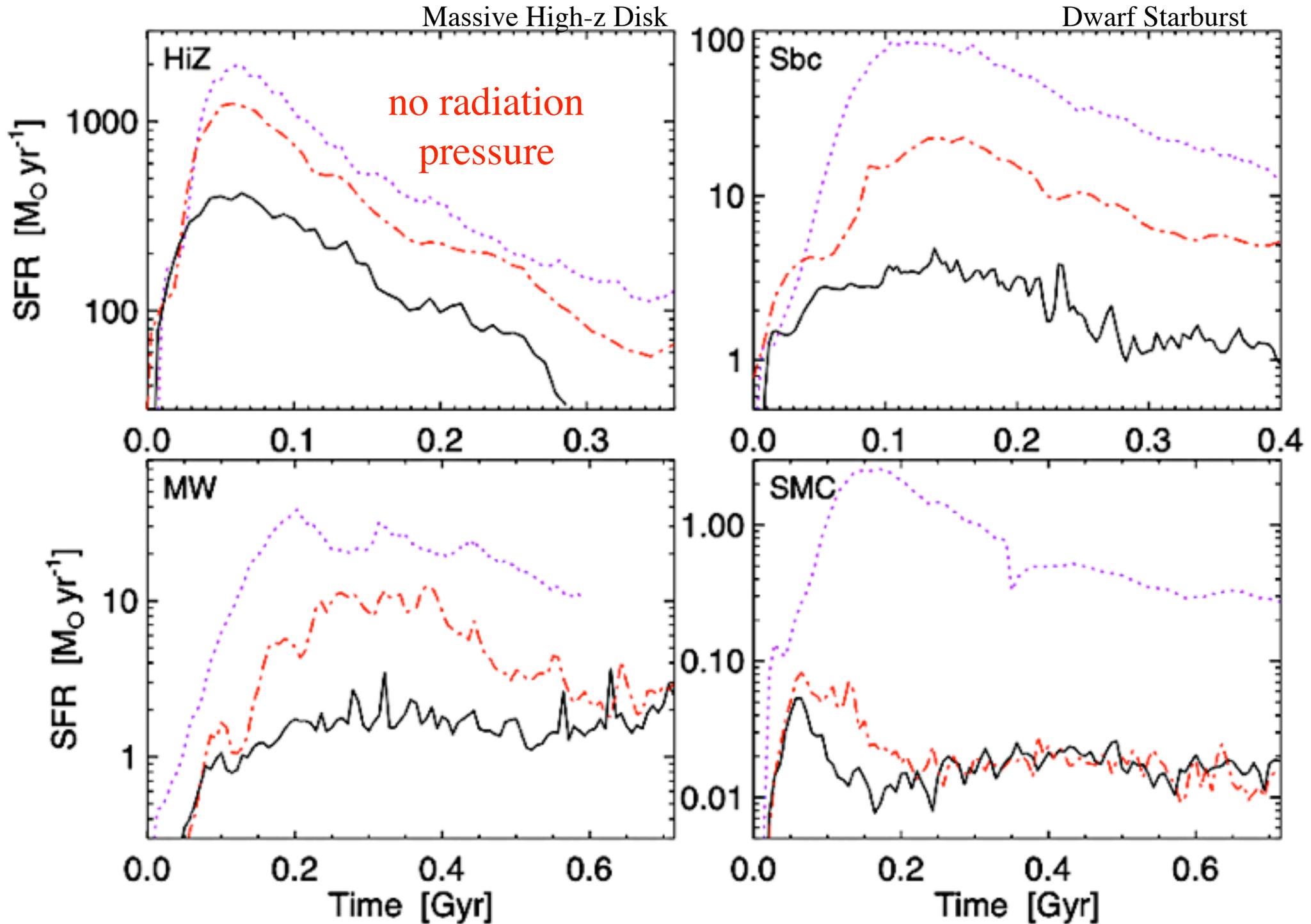
Stellar Feedback gives Self-Regulated Star Formation

Massive High-z Disk

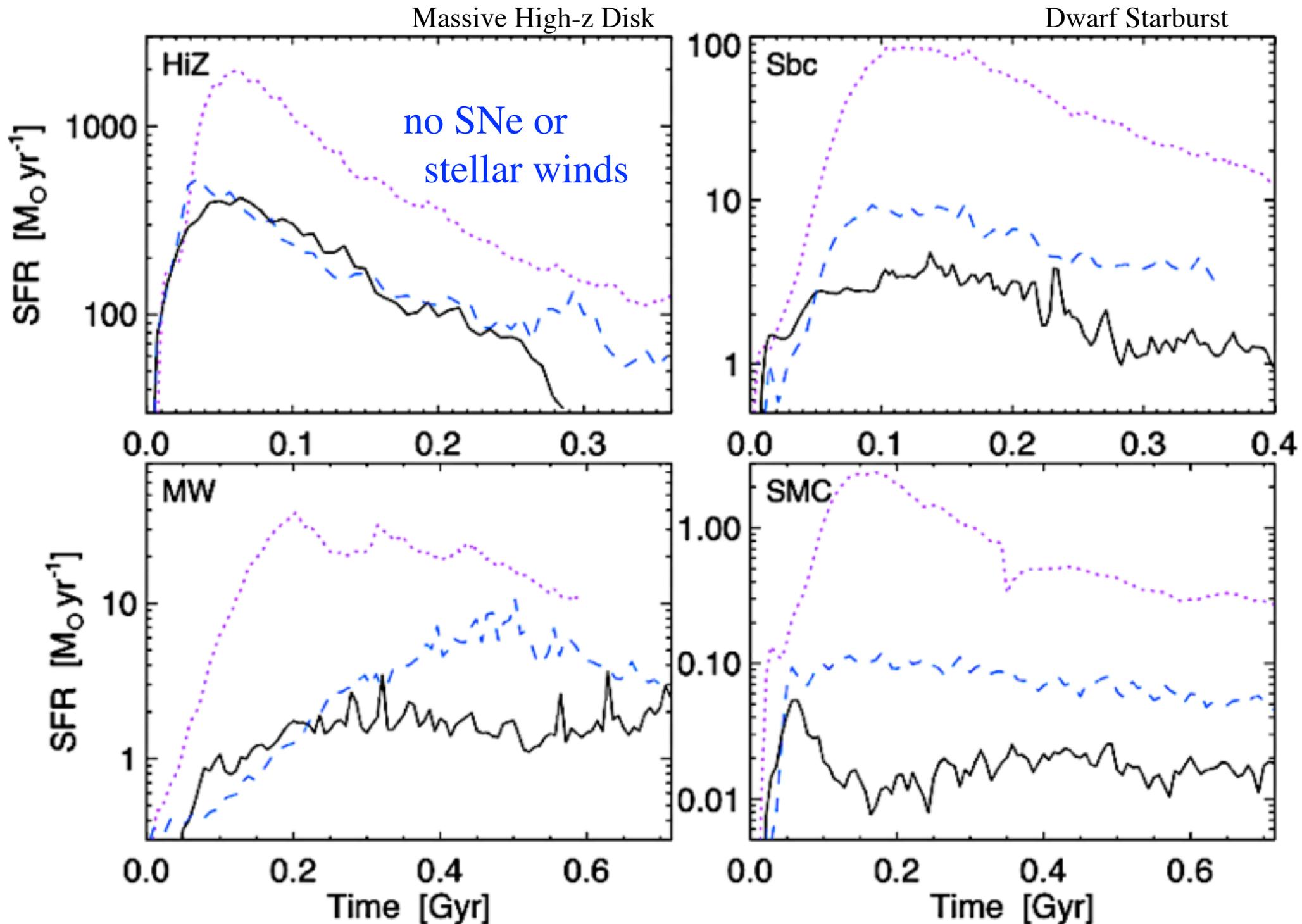
Dwarf Starburst



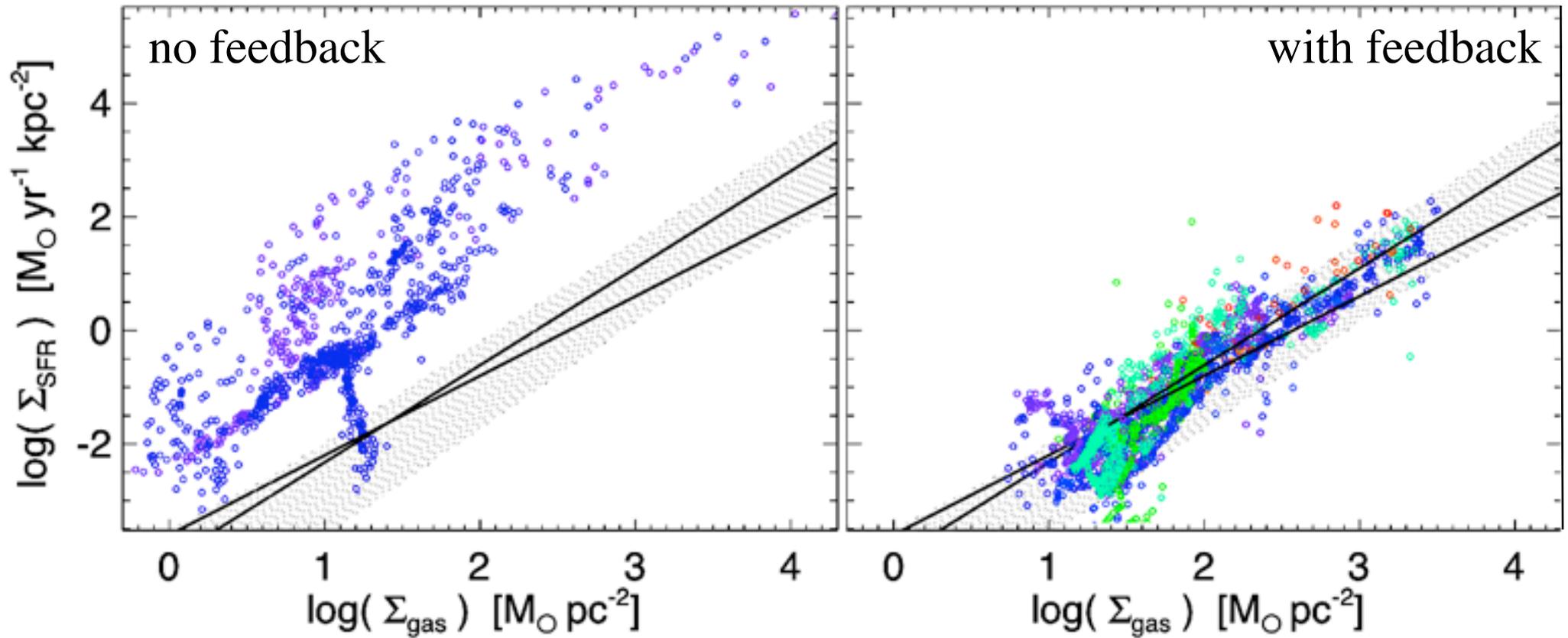
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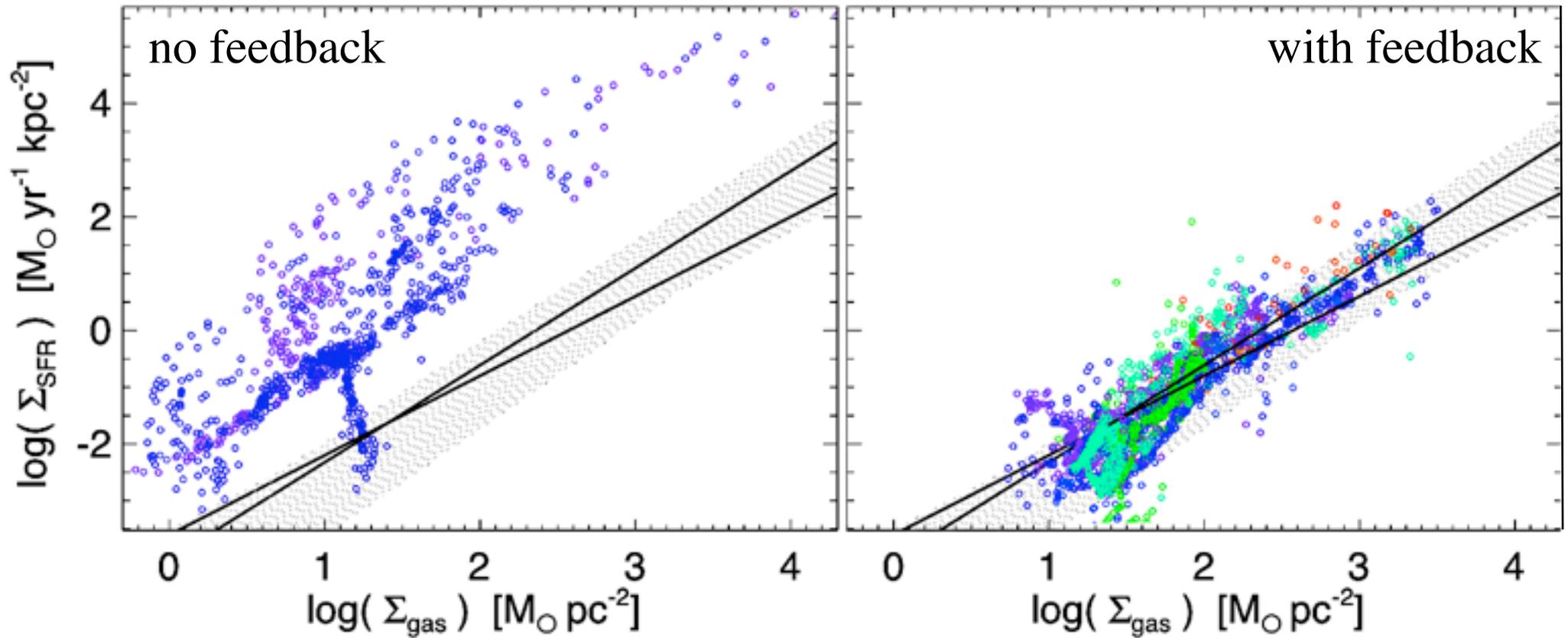
Kennicutt-Schmidt relation emerges naturally



PFH, Quataert, & Murray, 2011a

Kennicutt-Schmidt relation emerges naturally

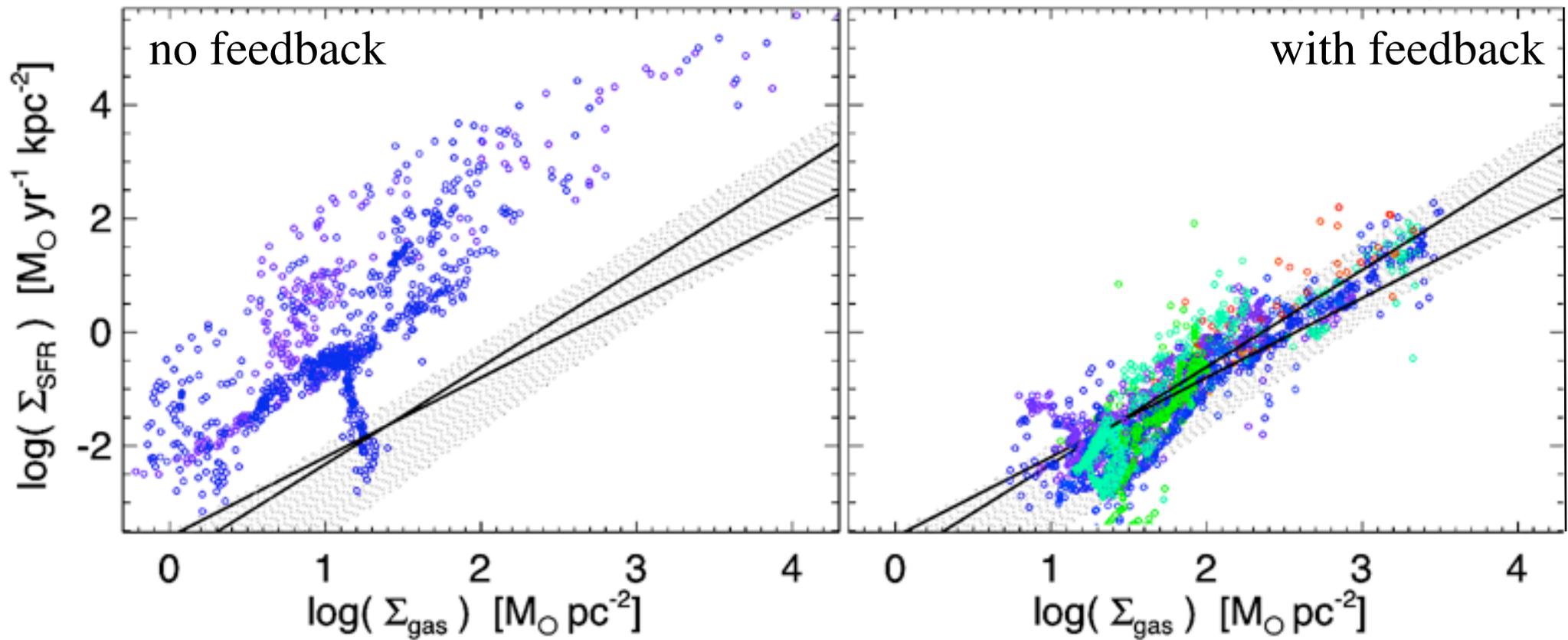
$$\dot{\Sigma}_* \sim \Sigma_{\text{gas}} / \tau_{\text{dyn}}$$



Kennicutt-Schmidt relation emerges naturally

$$\dot{\Sigma}_* \sim \Sigma_{\text{gas}} / \tau_{\text{dyn}}$$

$$\dot{\Sigma}_* \sim 0.02 \Sigma_{\text{gas}} / \tau_{\text{dyn}}$$

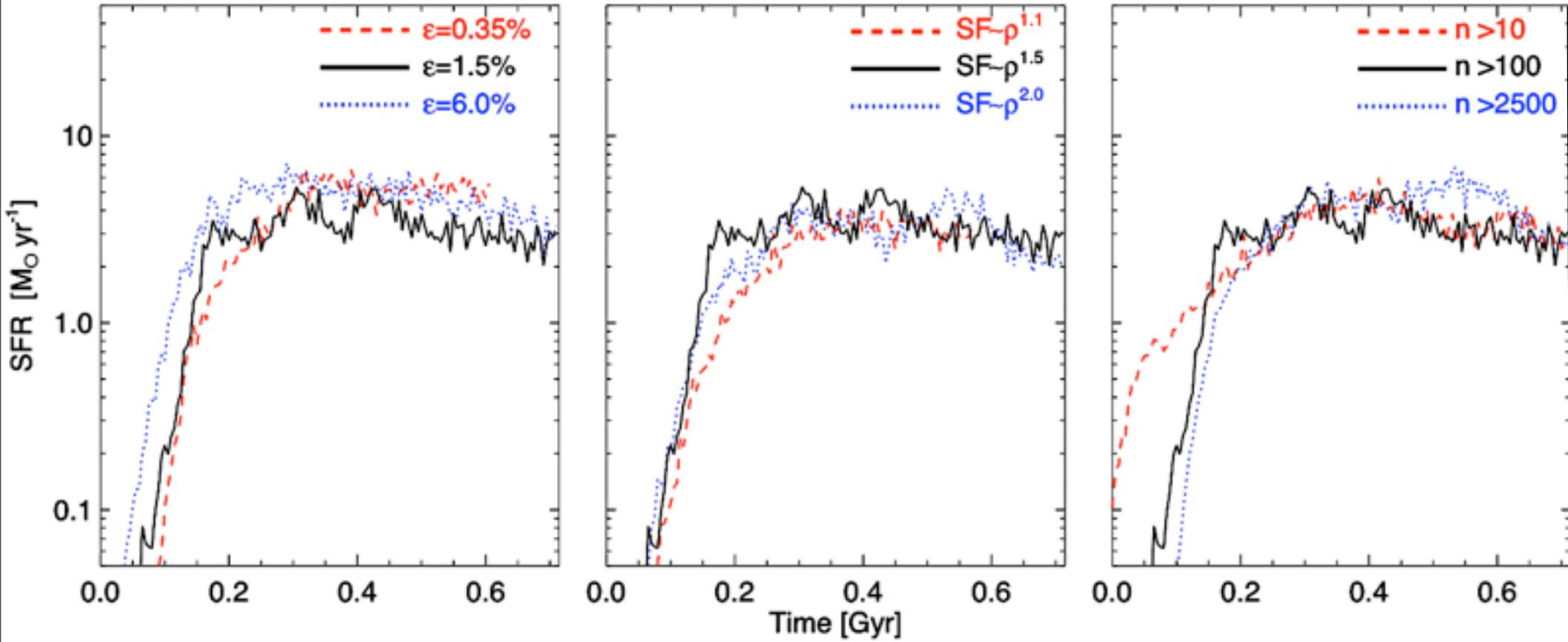


Global Star Formation Rates are *INDEPENDENT* of High-Density SF Law

Efficiency (SF per t_{dyn})

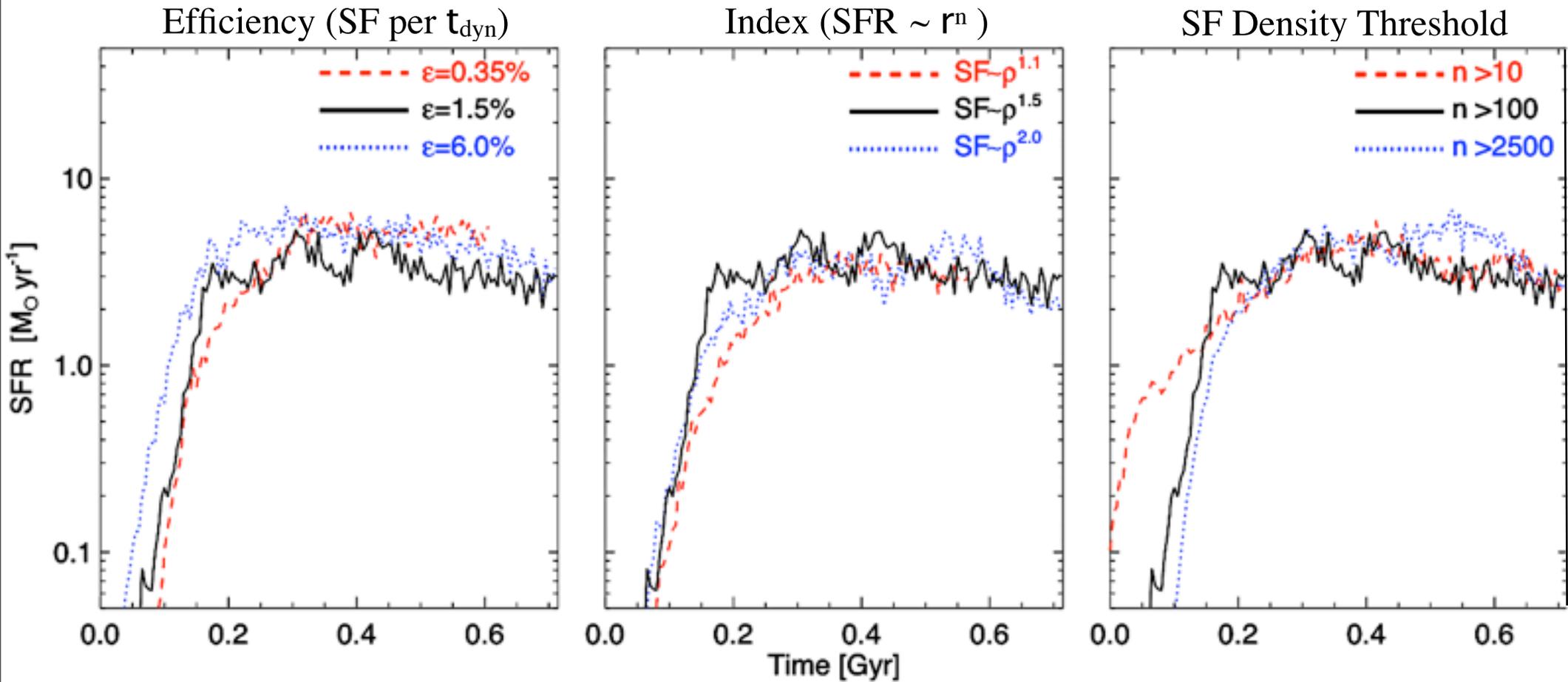
Index (SFR $\sim r^n$)

SF Density Threshold



Hopkins, Quataert, & Murray 2011
also Saitoh et al. 2008

Global Star Formation Rates are *INDEPENDENT* of High-Density SF Law



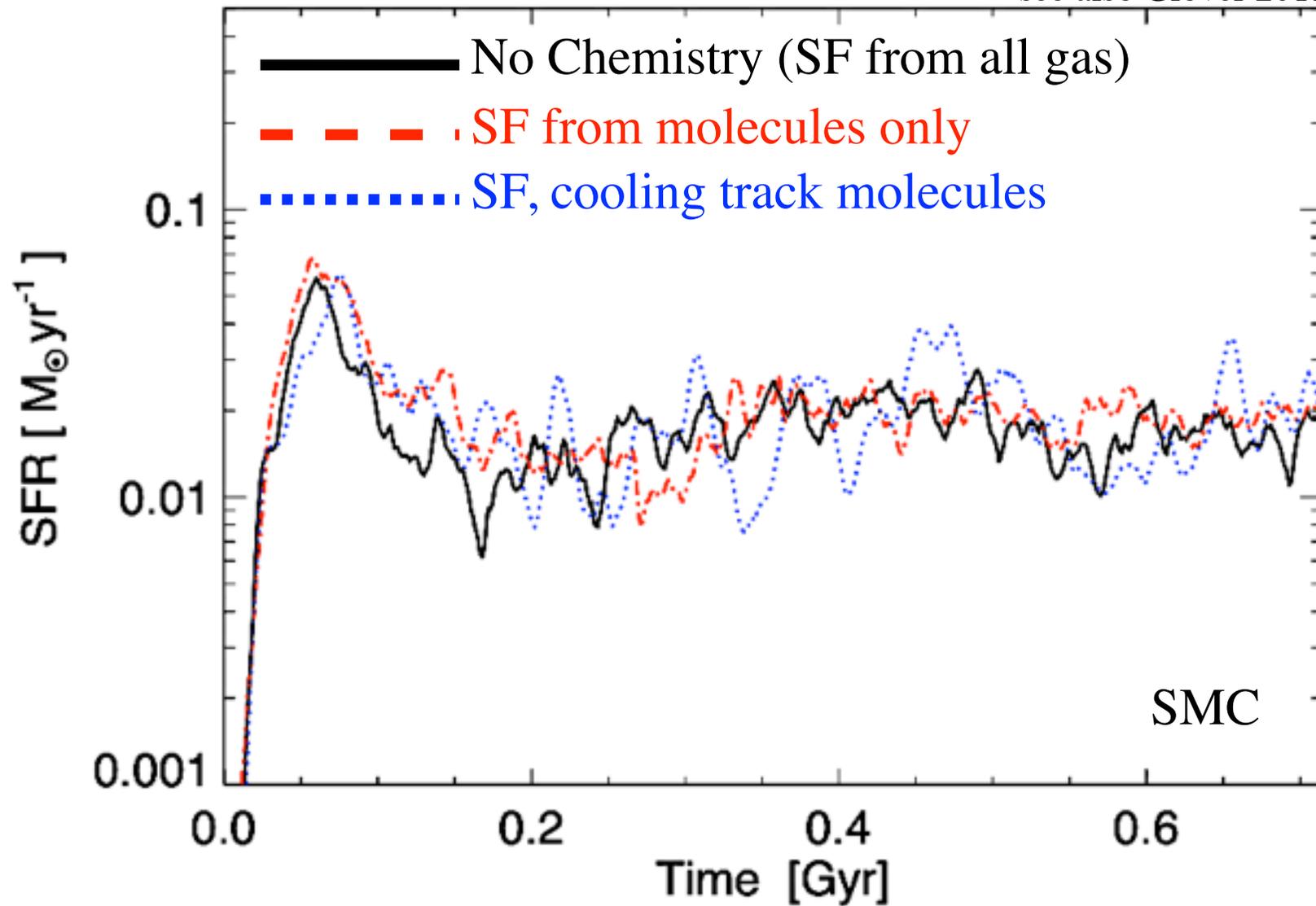
➤ Set by feedback (i.e. SFR) needed to maintain marginal stability

Hopkins, Quataert, & Murray 2011
also Saitoh et al. 2008

Molecular Chemistry doesn't change things above modest Metallicity

MOLECULES ARE A *TRACER*

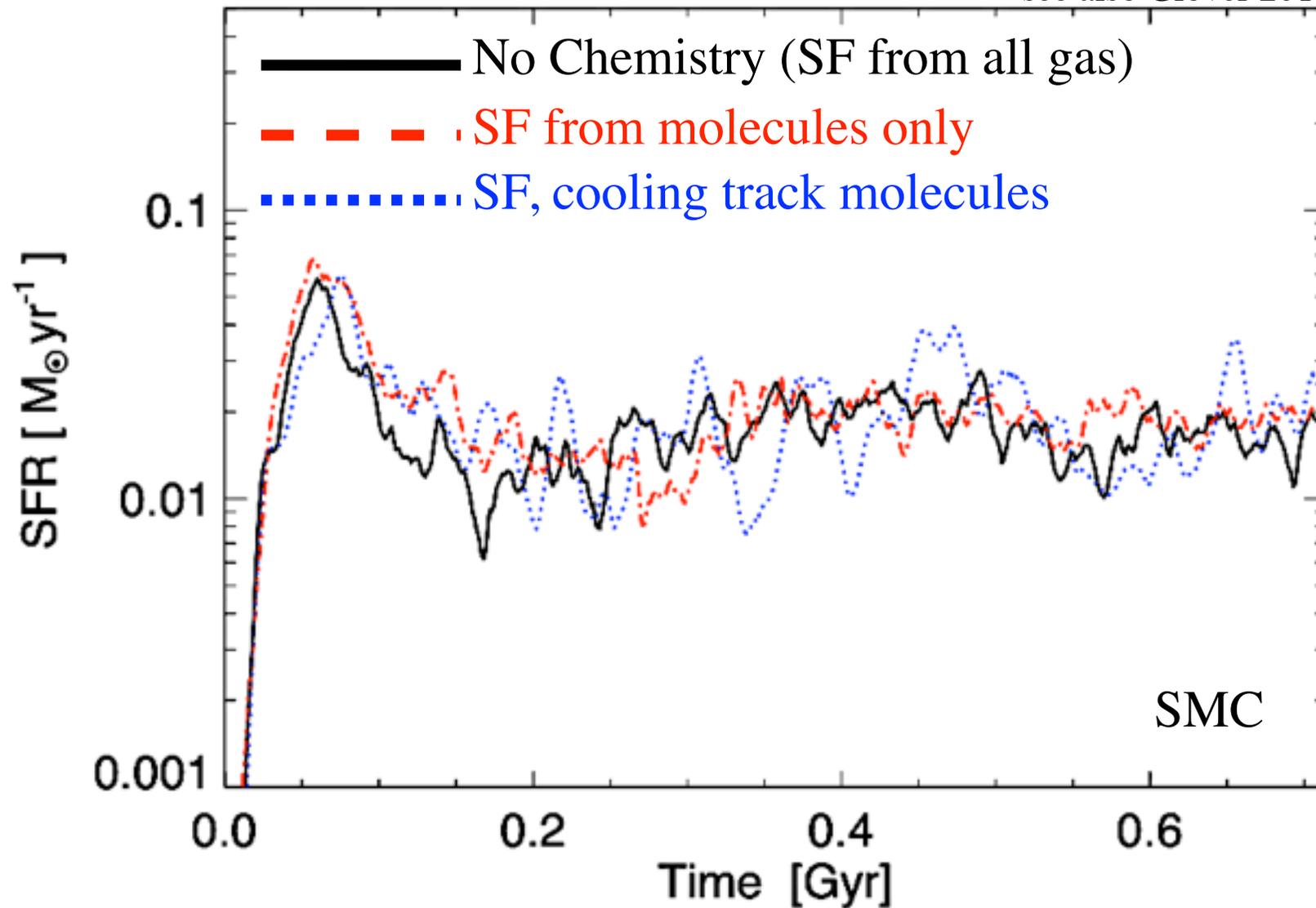
see also Glover 2011



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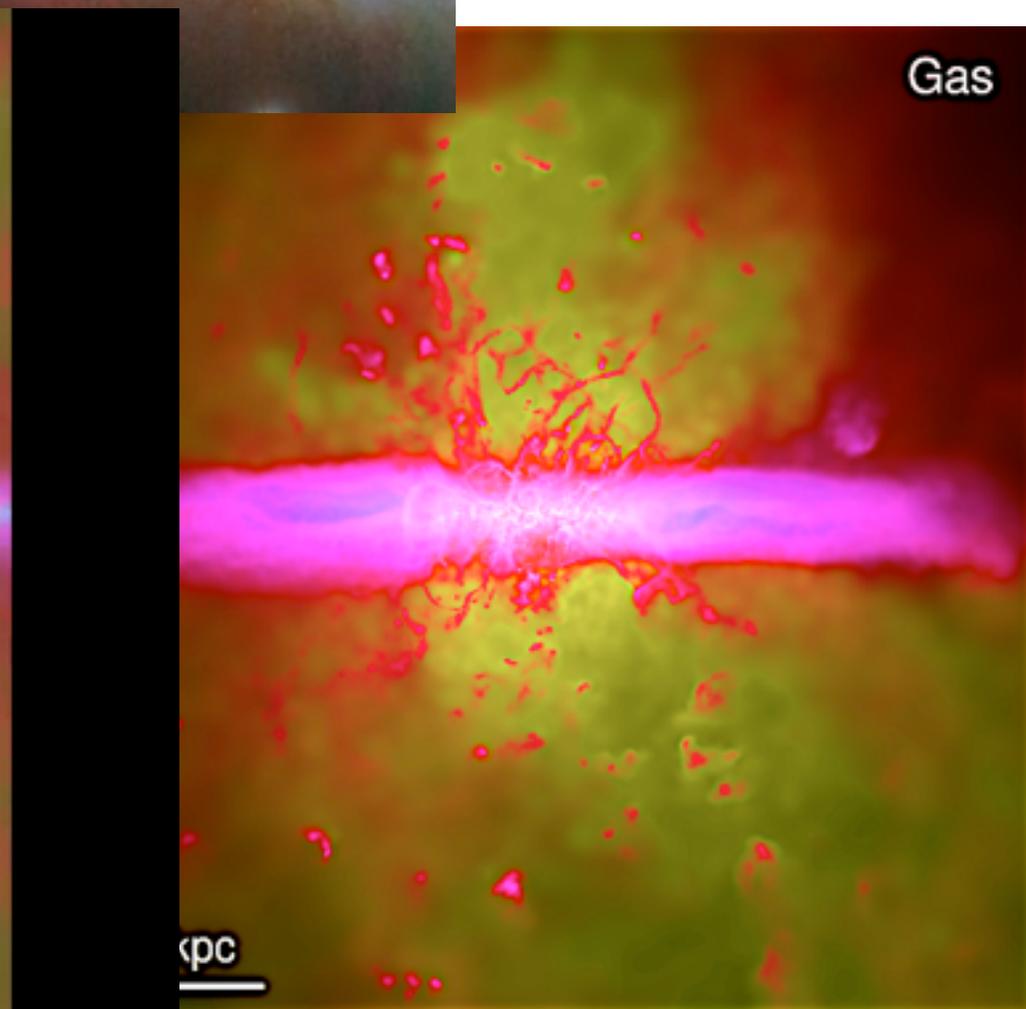
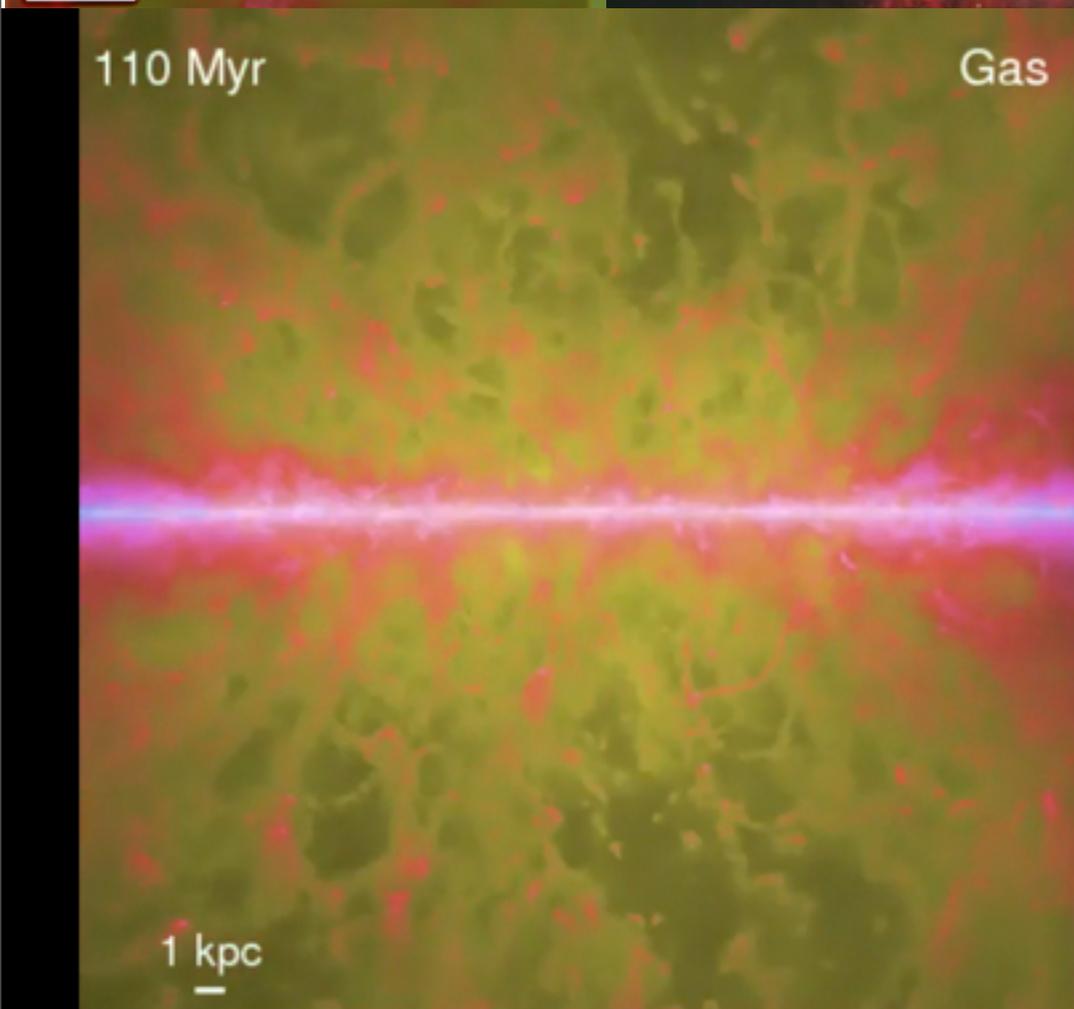
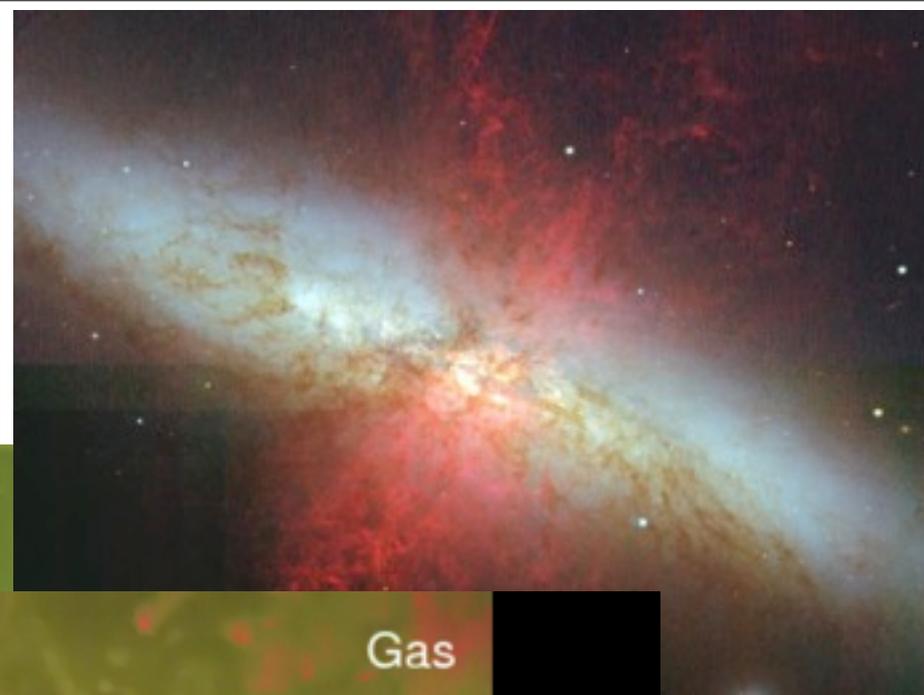
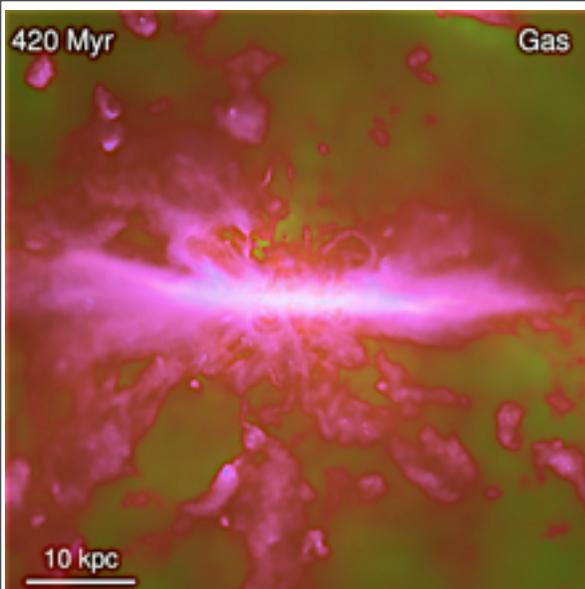
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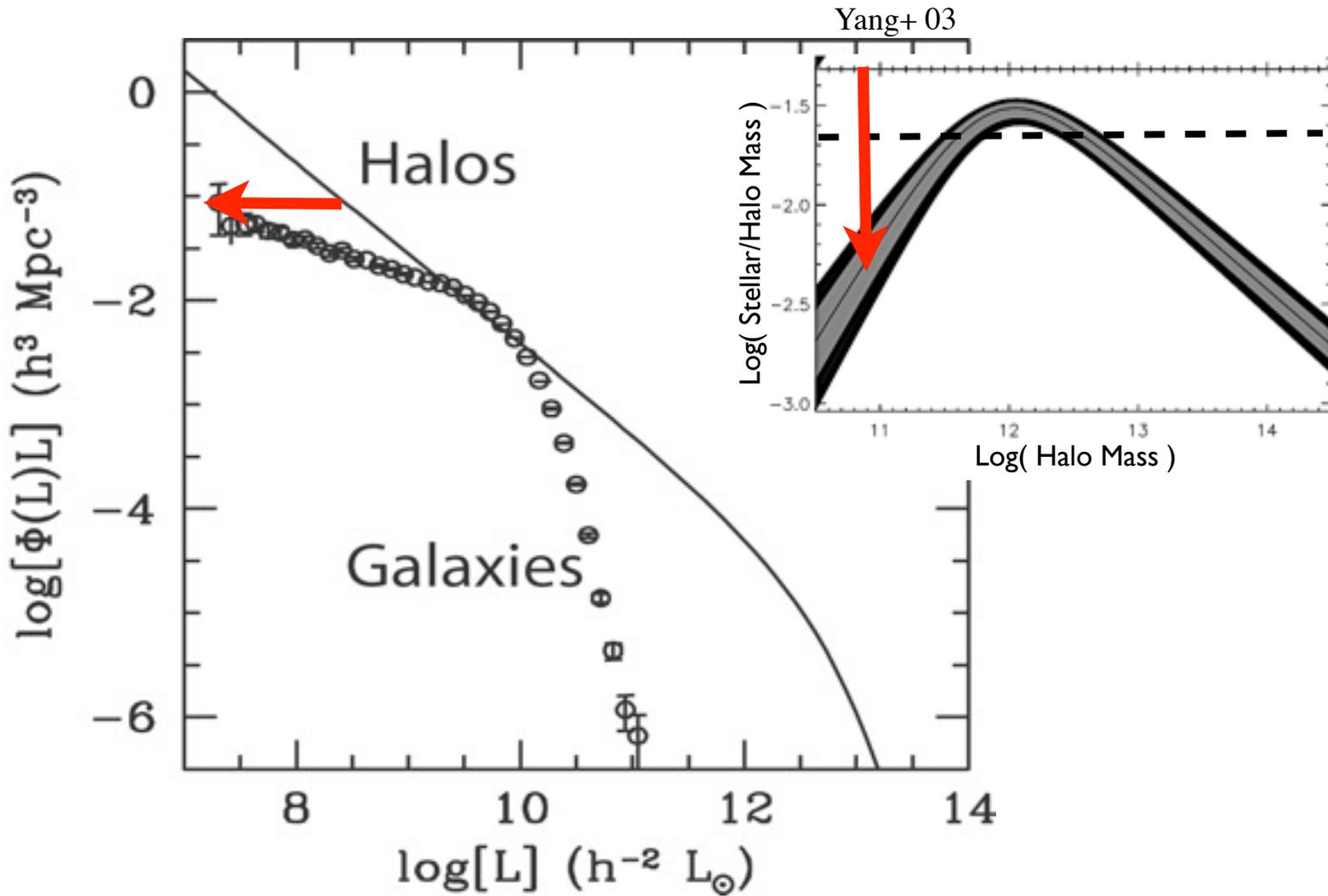


➤ Just need *some* cooling channel: changes at $M_{\text{gal}} < 10^6 M_{\text{sun}}$, $Z < 0.01 Z_{\text{sun}}$

Galactic Super-Winds



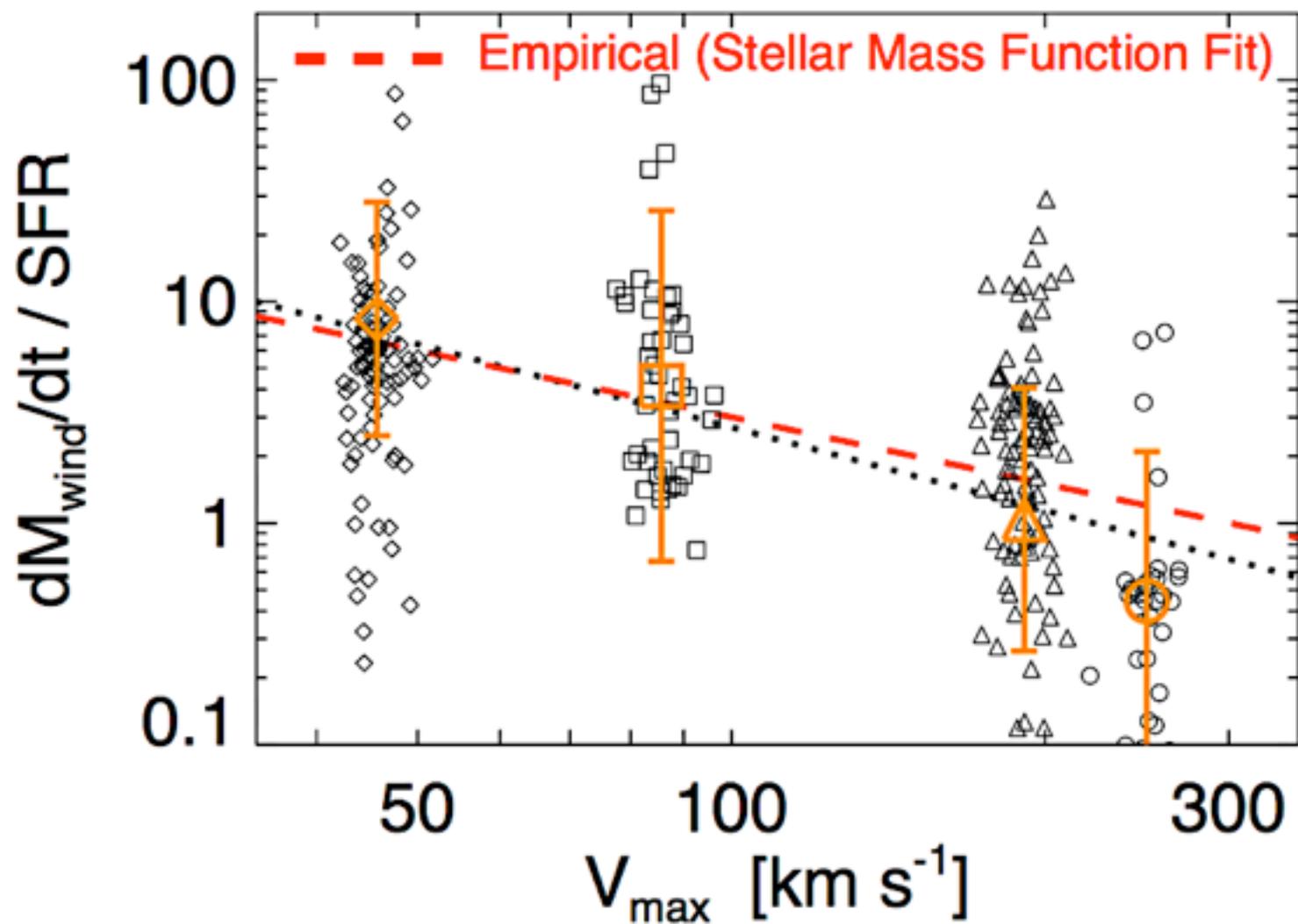
How Efficient Are Galactic Super-Winds?



➤ Large mass-loading:

$$\dot{M}_{\text{wind}} \approx 10 \dot{M}_{*} \left(\frac{V_c}{100 \text{ km s}^{-1}} \right)^{-1.1} \left(\frac{\Sigma_{\text{gas}}}{10 M_{\odot} \text{ pc}^{-2}} \right)^{-0.5}$$

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The Cosmological Inflow/Outflow Cycle

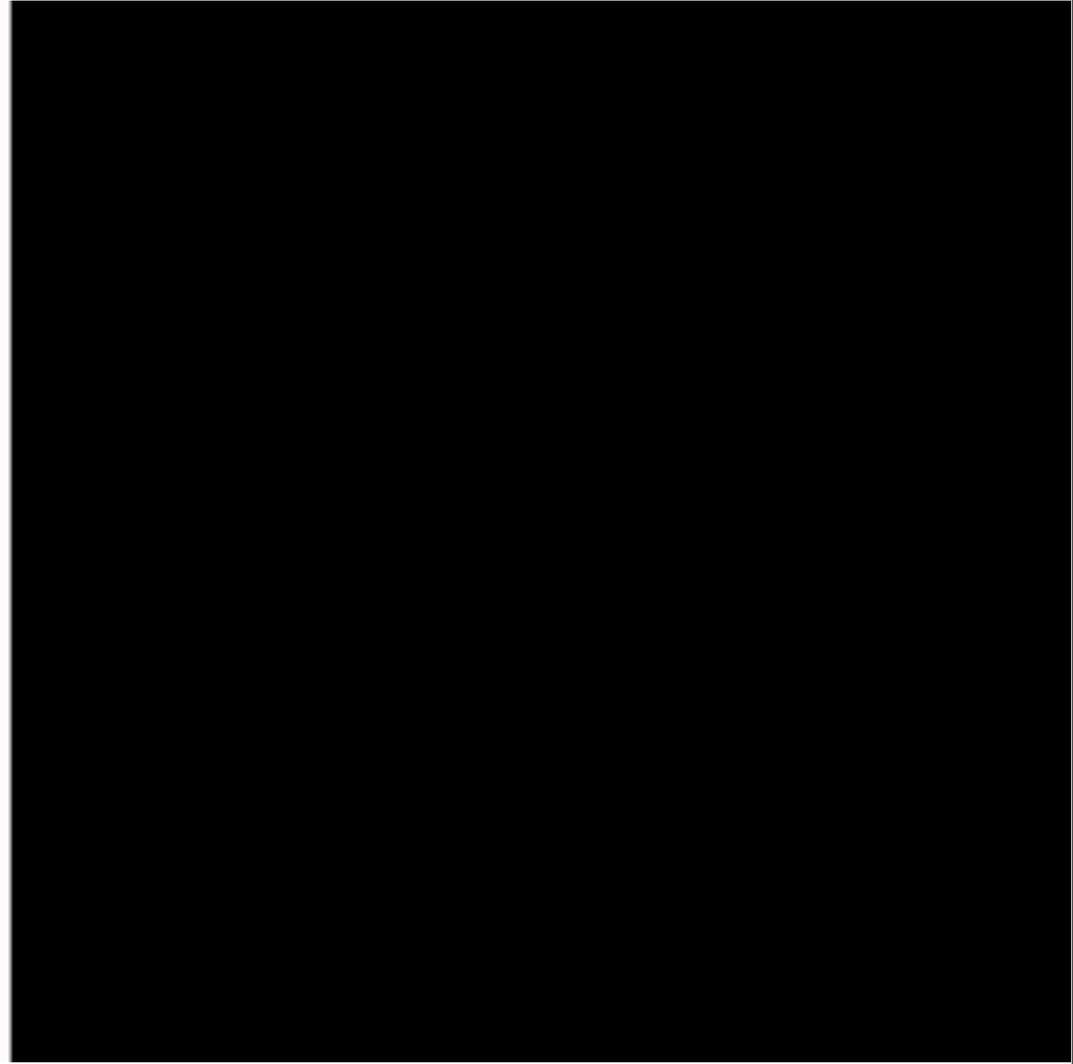
Cosmological Simulations

“ZOOM-IN” ON THE FORMATION OF A MASSIVE GALAXY

$z=29.99$ box=200/h kpc(phys)



IGM Density

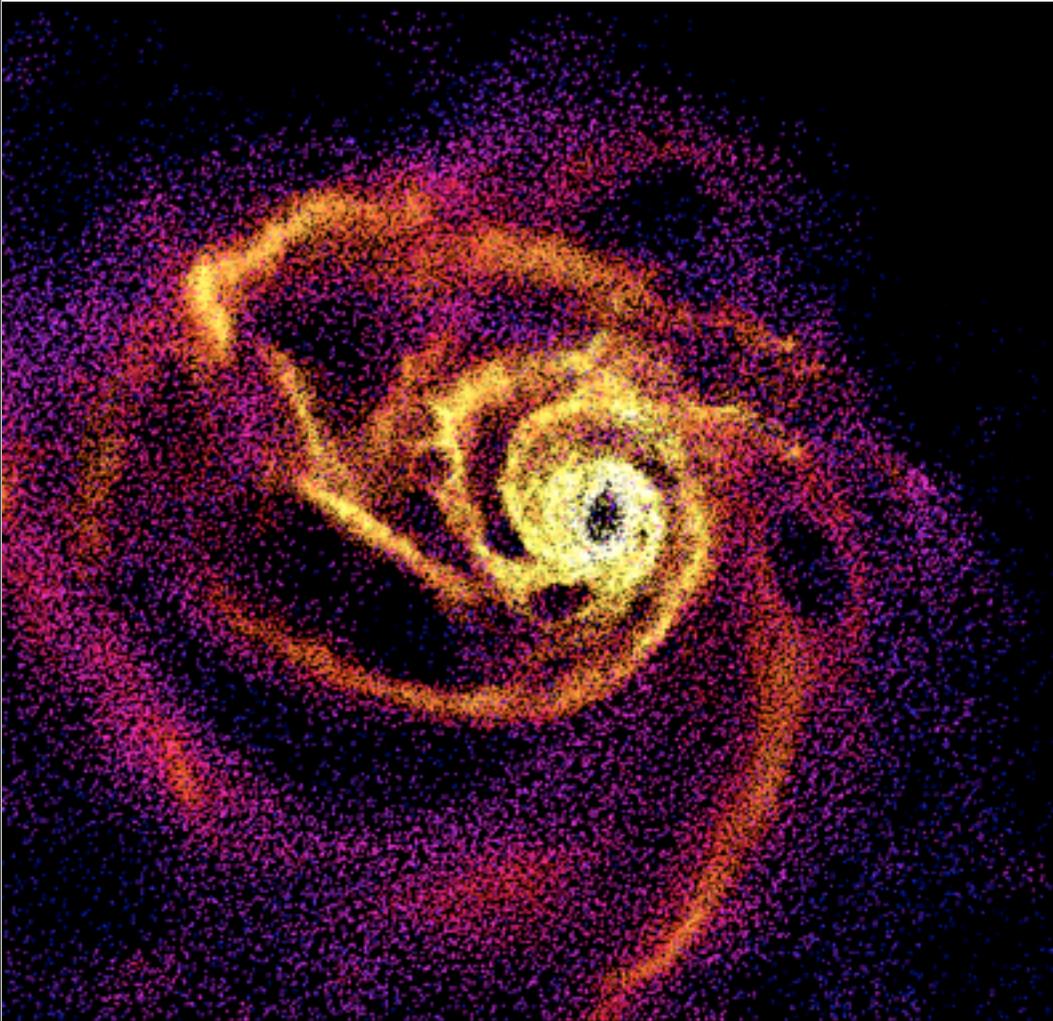


IGM Temperature

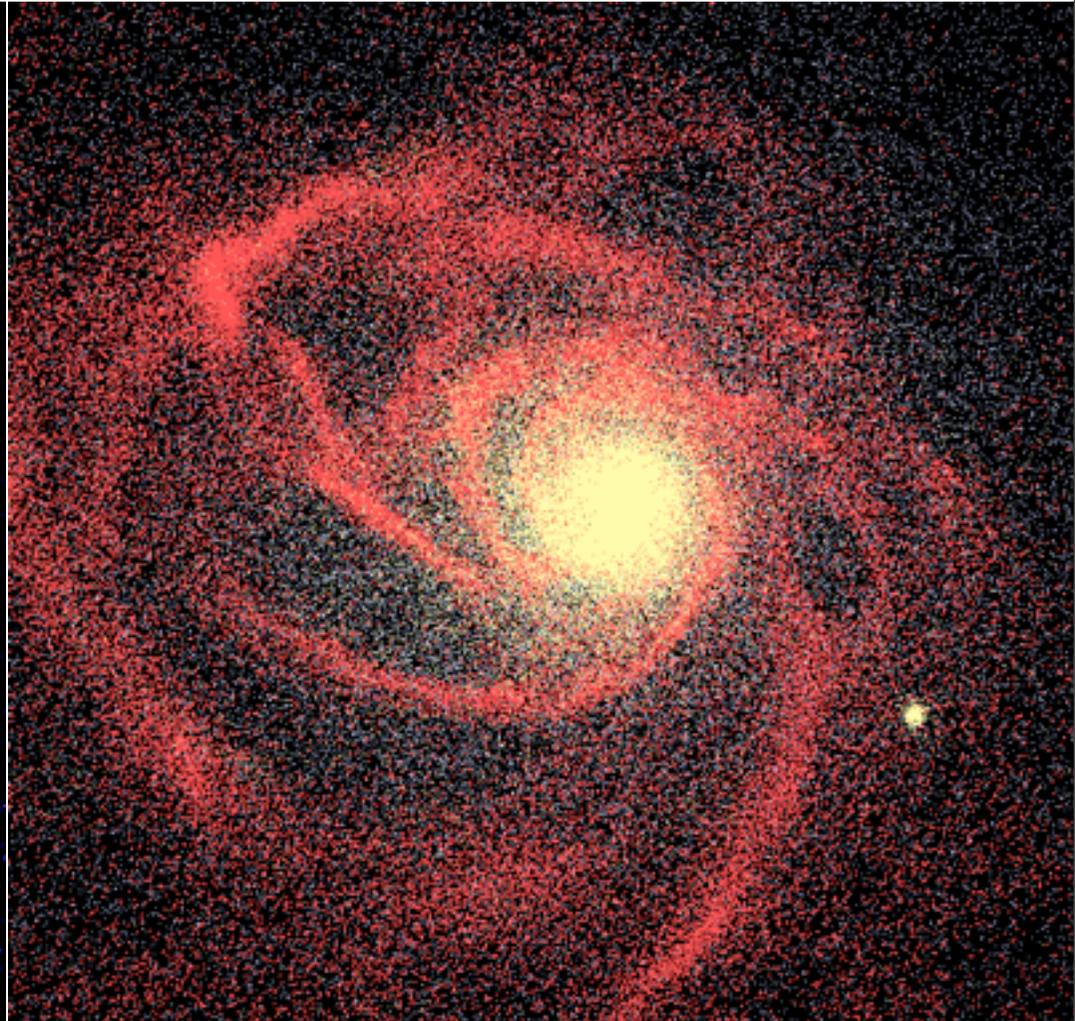
Keres & PFH et al

Cosmological Simulations

“ZOOM-IN” ON THE FORMATION OF A MASSIVE GALAXY



Gas Density

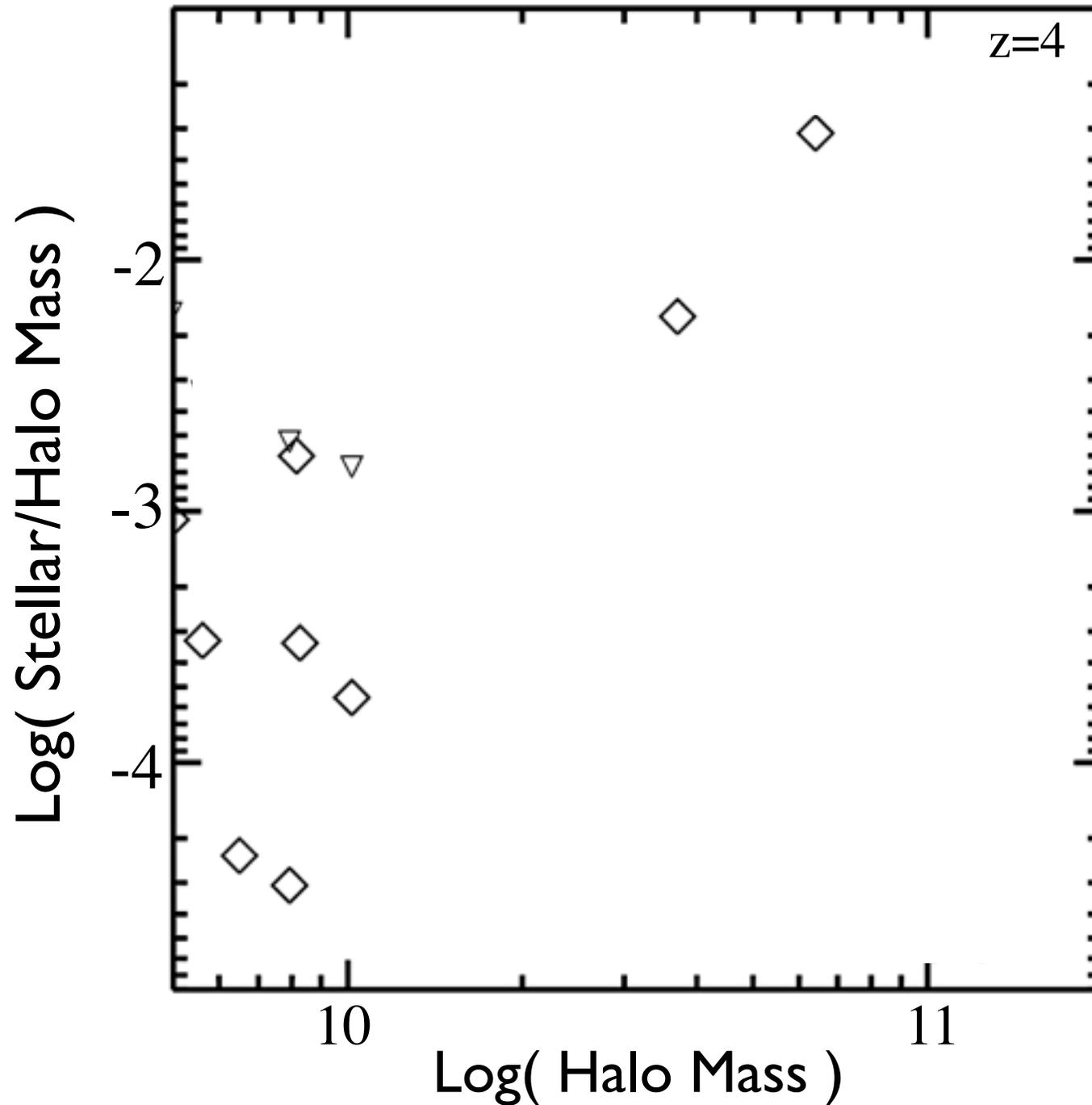


Stars

Keres & PFH et al

How Inefficient is Galaxy Formation?

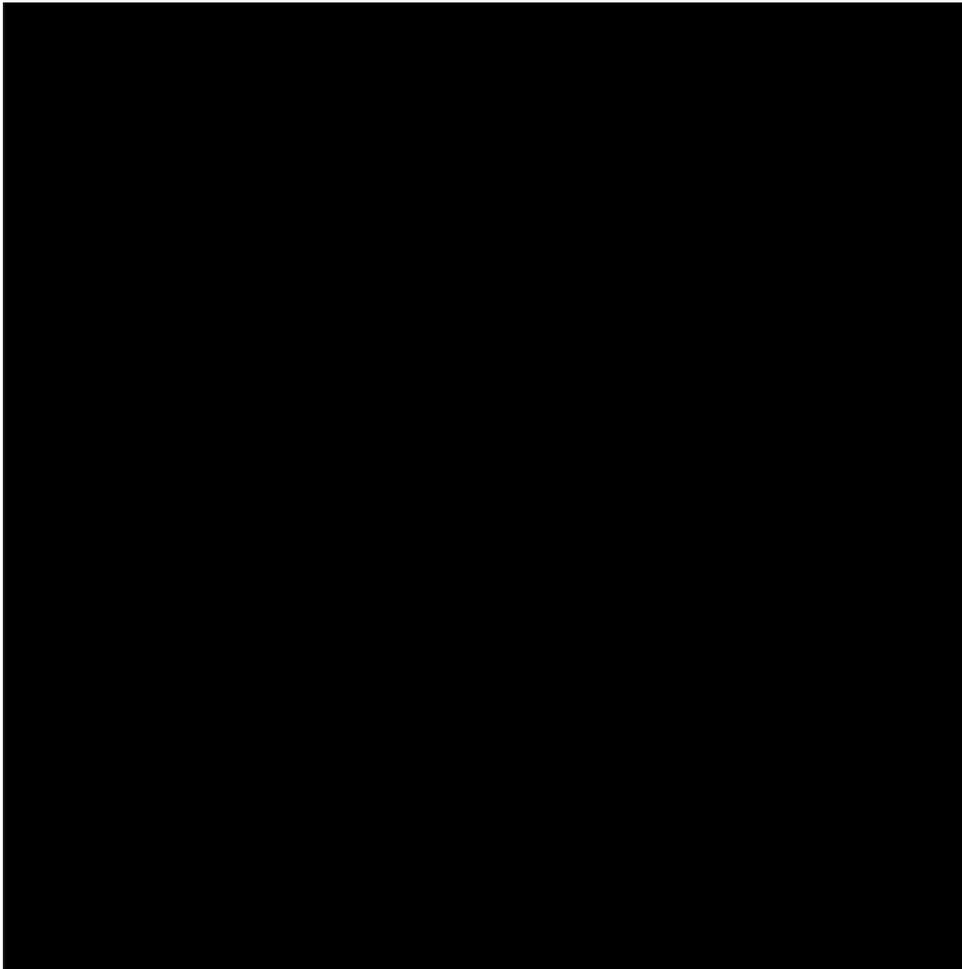
HELP WITH THE *FORMING-TOO-MANY-STARS-AT-HIGH-REDSHIFT-CATASTROPHE*?



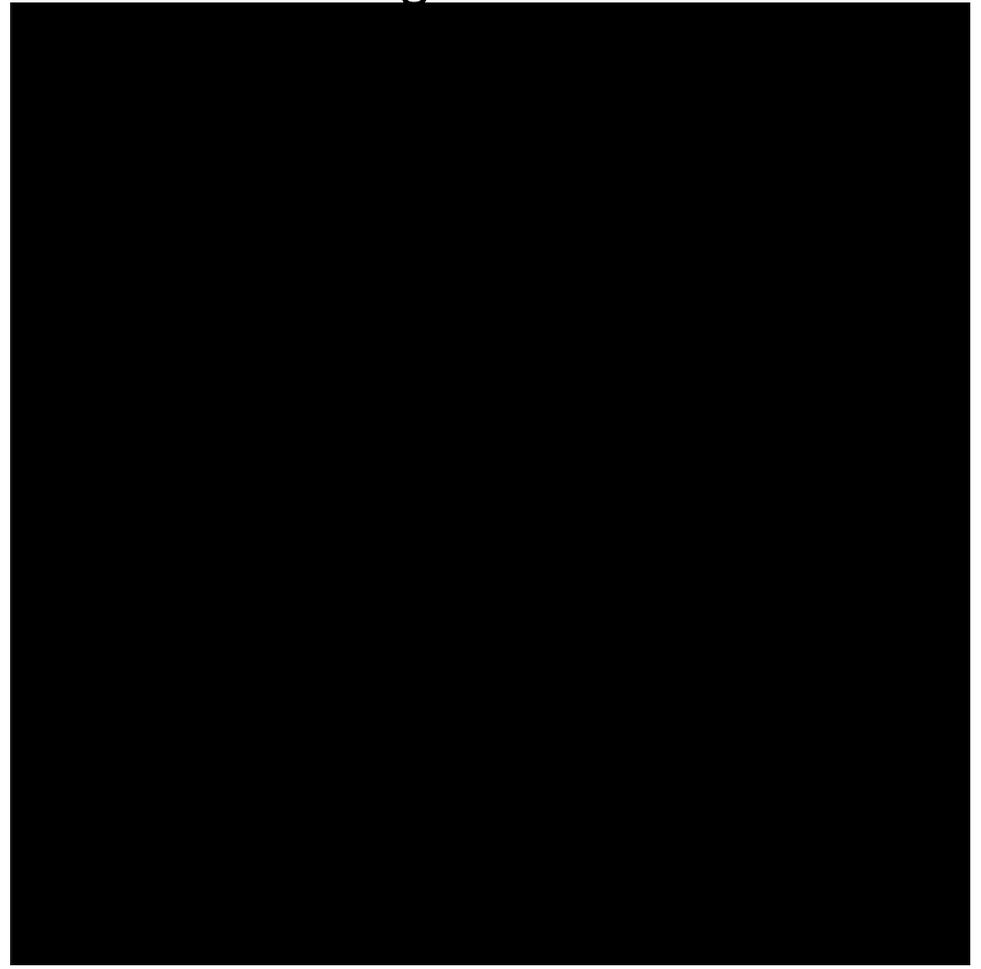
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Proto-MW: Gas Temperature:

No Feedback

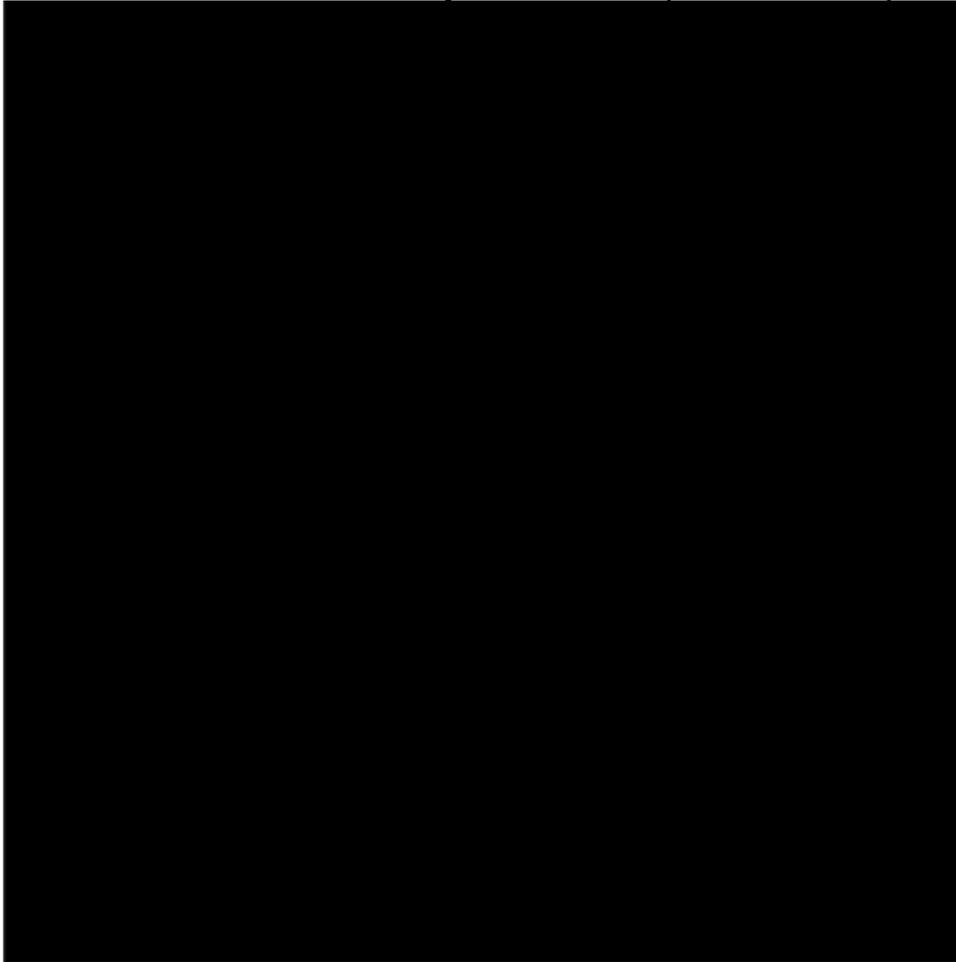


Following Full Feedback

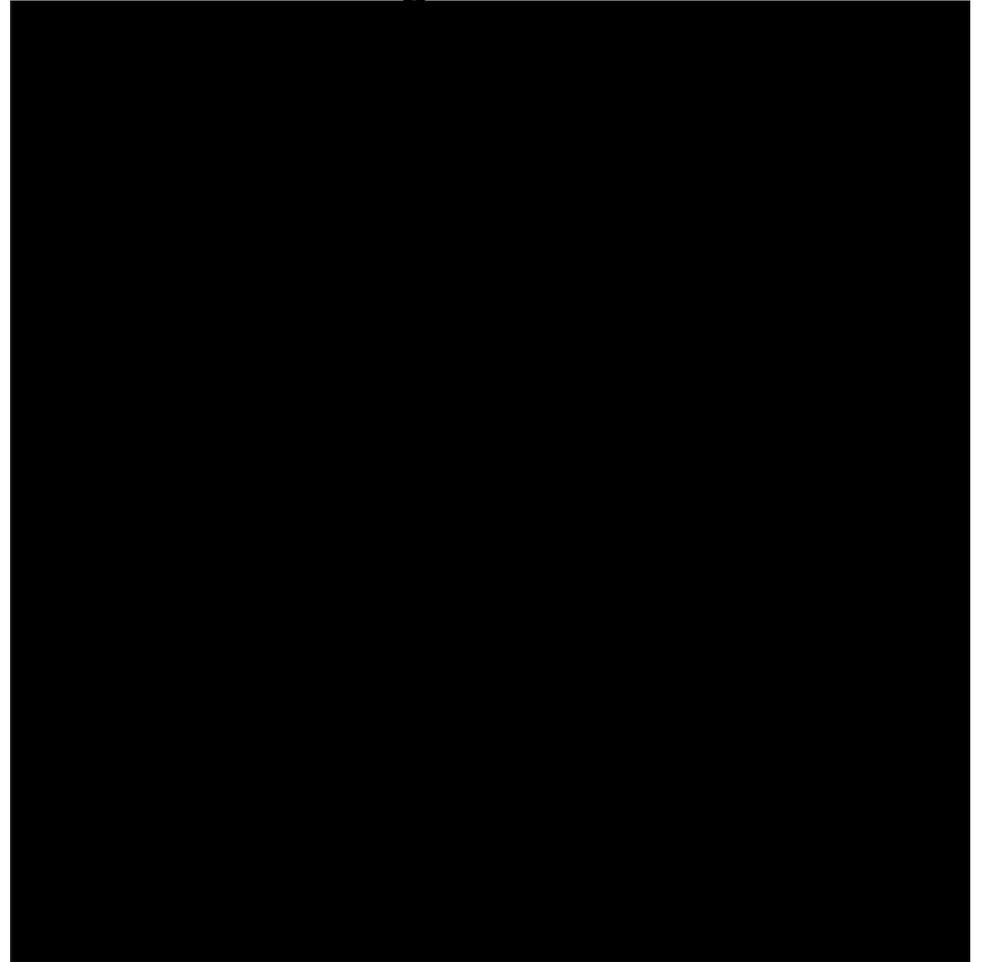


Proto-MW: Gas Temperature:

Insert Winds “By Hand” (Sub-Grid)

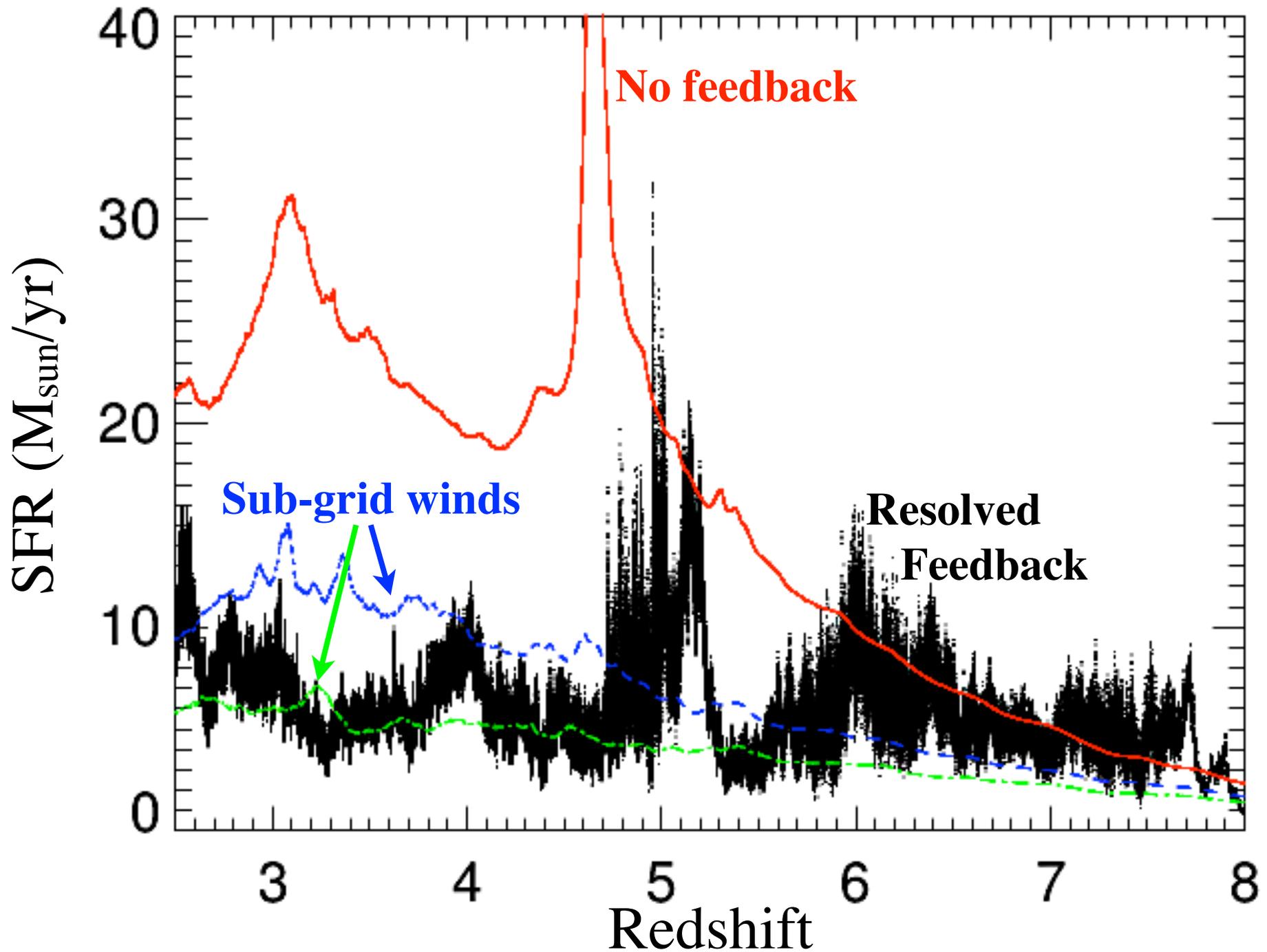


Following Full Feedback



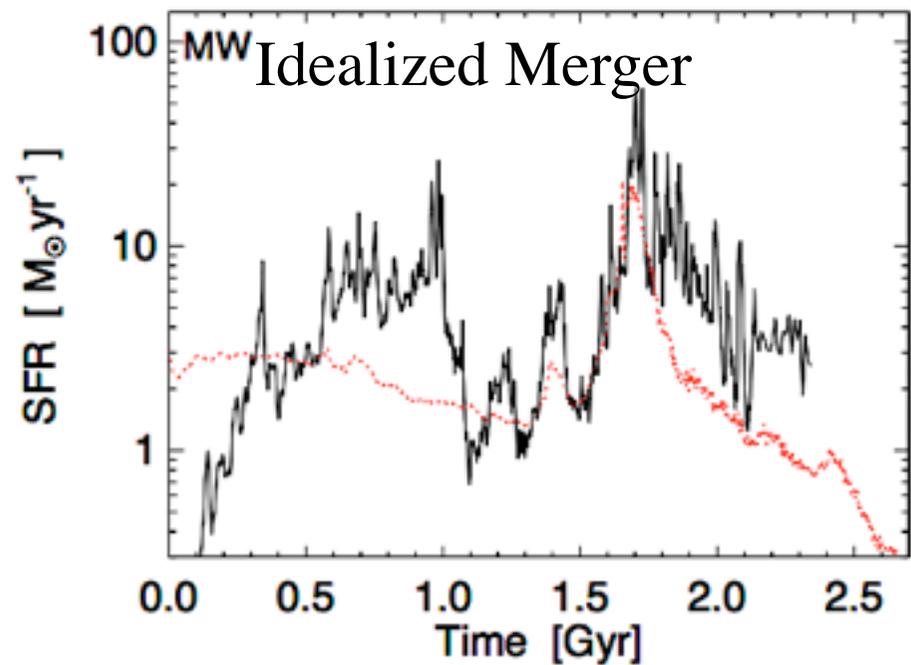
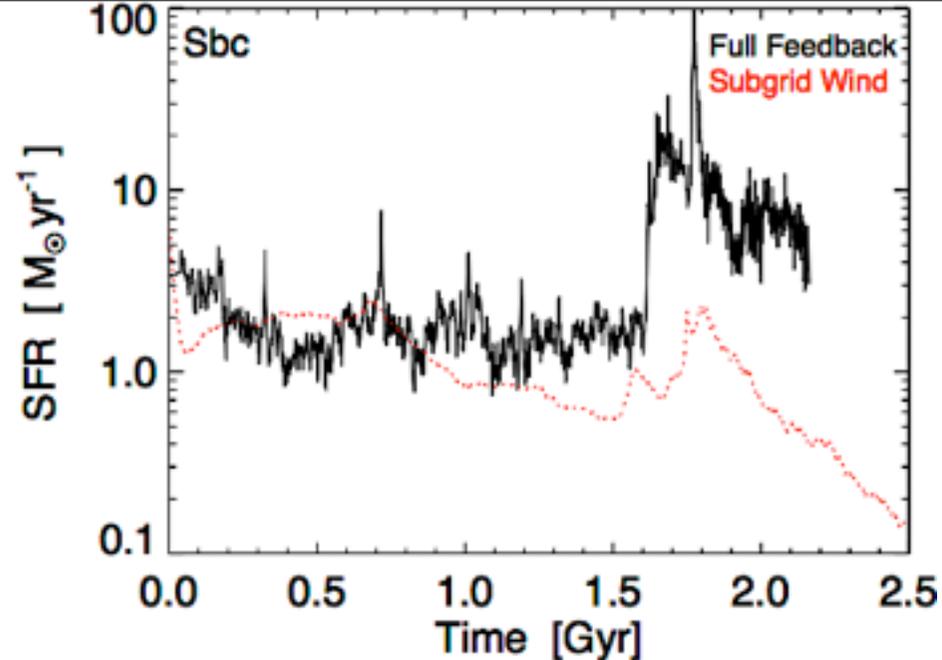
Starburst-Driven Winds

SUB-GRID vs. RESOLVED MATTERS!



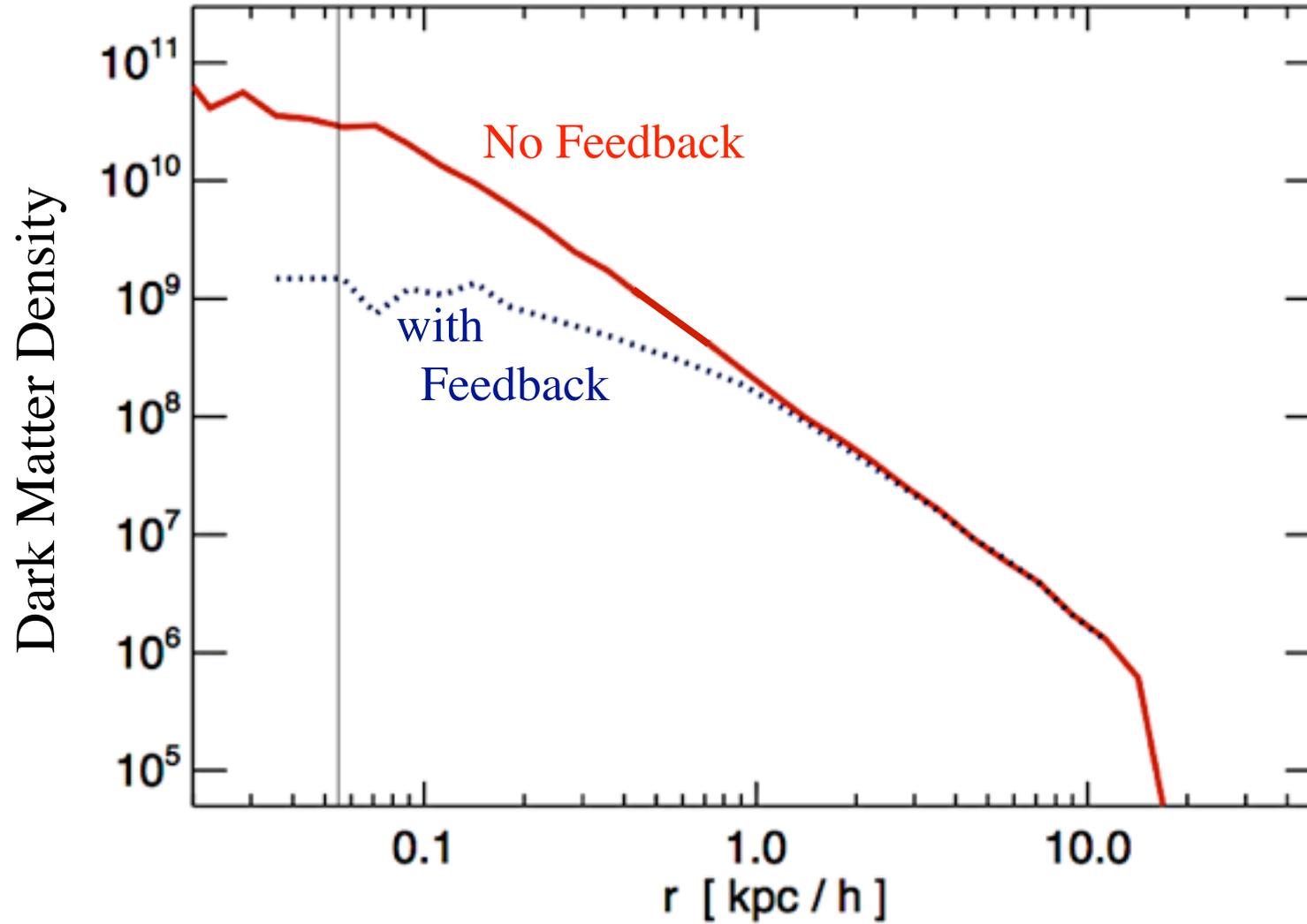
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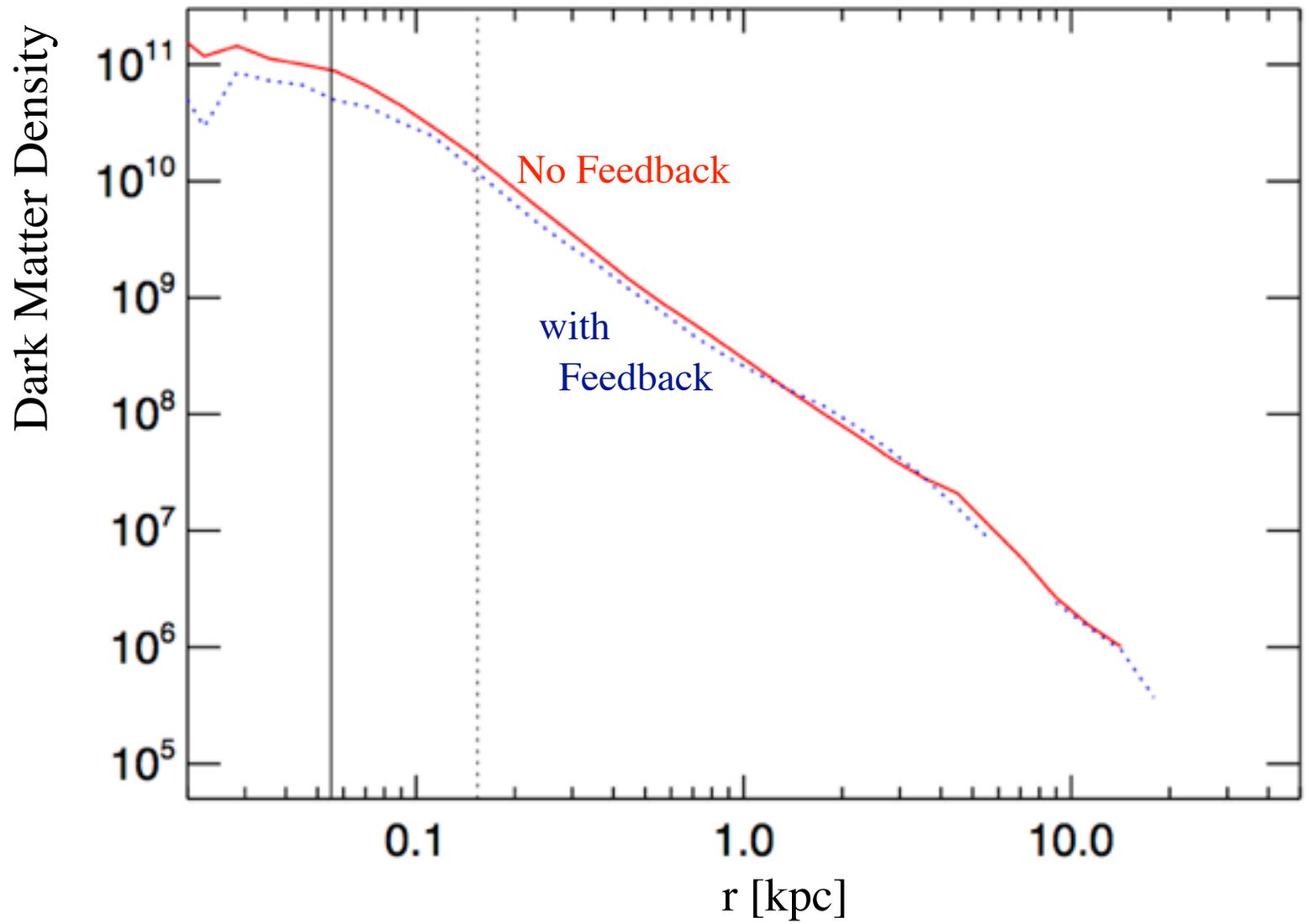
Dark Matter Profiles: Baryons or Cosmology?

DO RESOLVED WINDS ACTUALLY MAKE CORES?



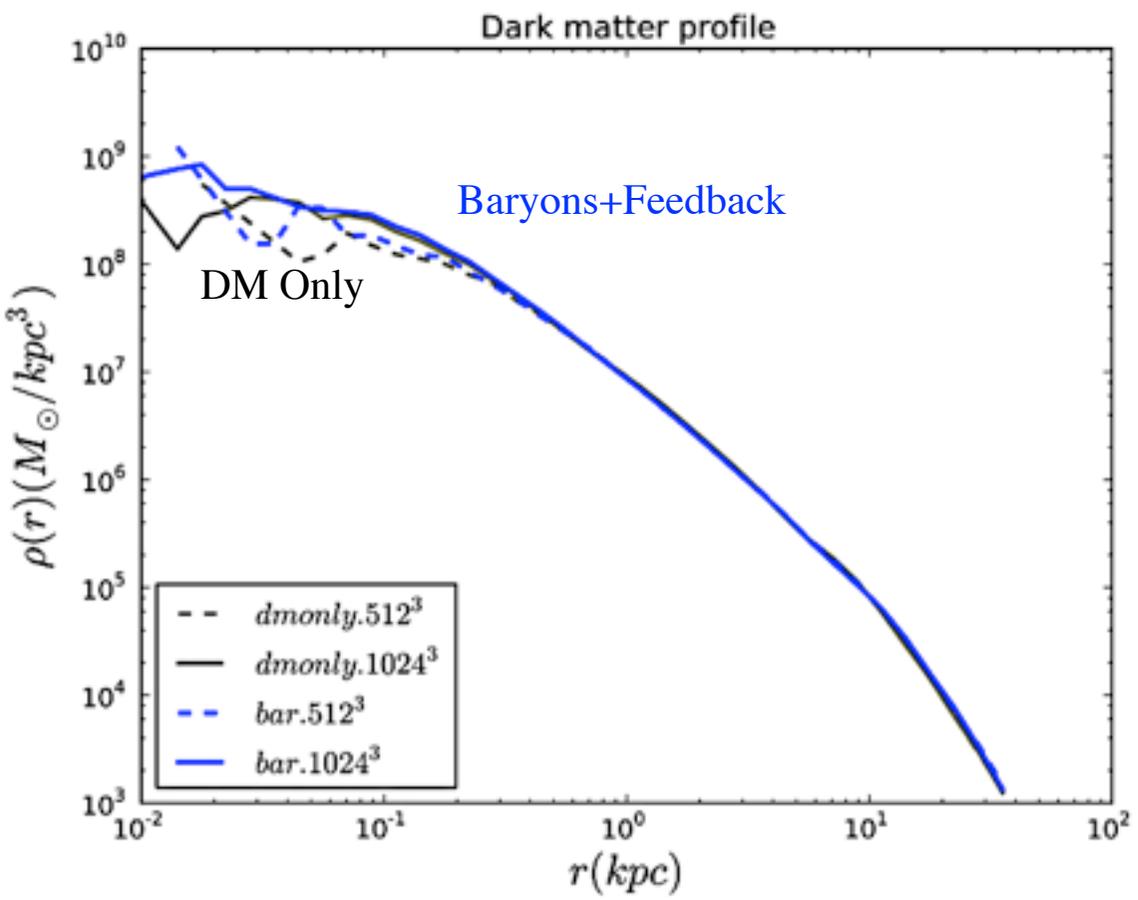
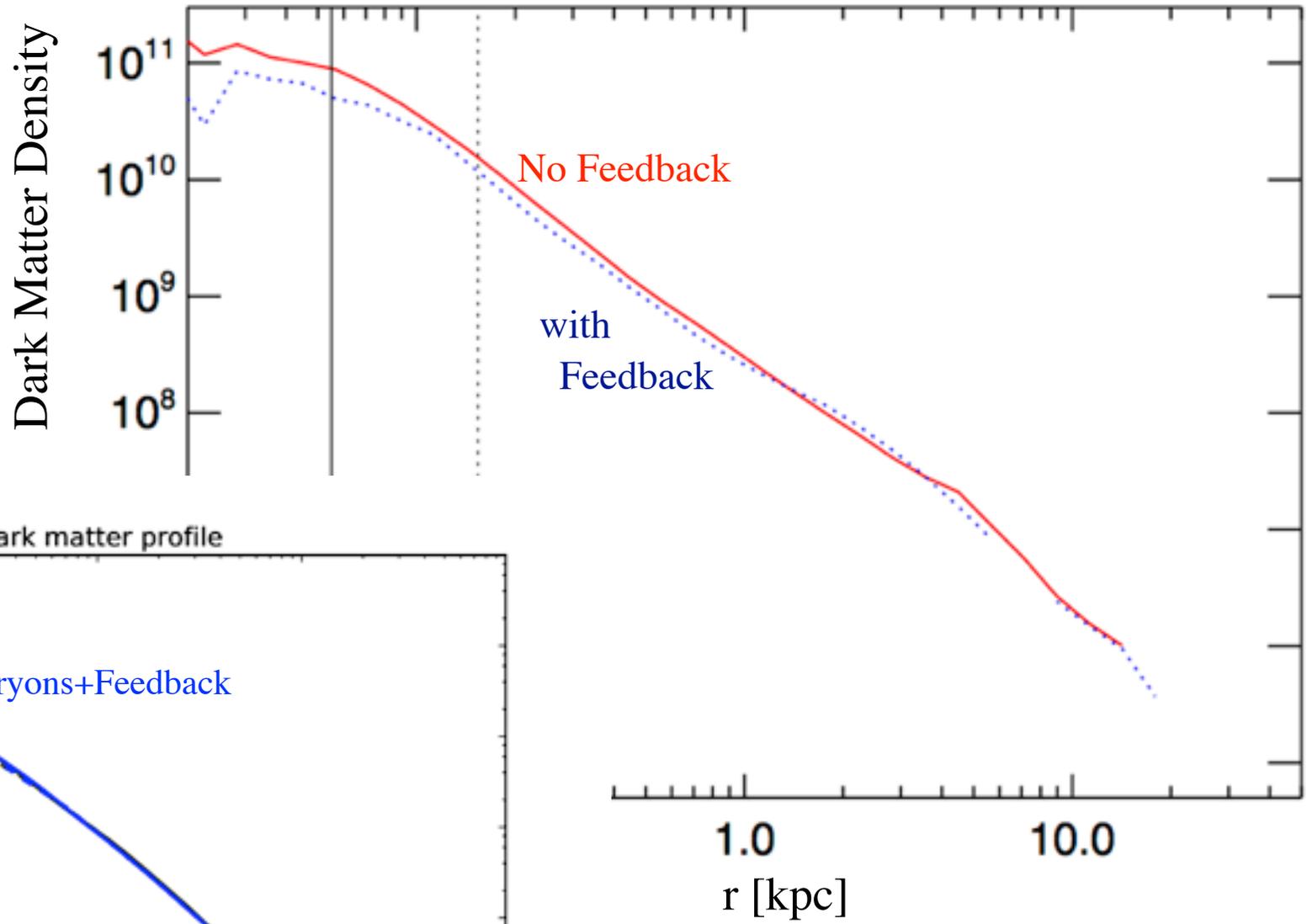
Keres & PFH et al
Bullock,
& Onorbe et al

BUT...



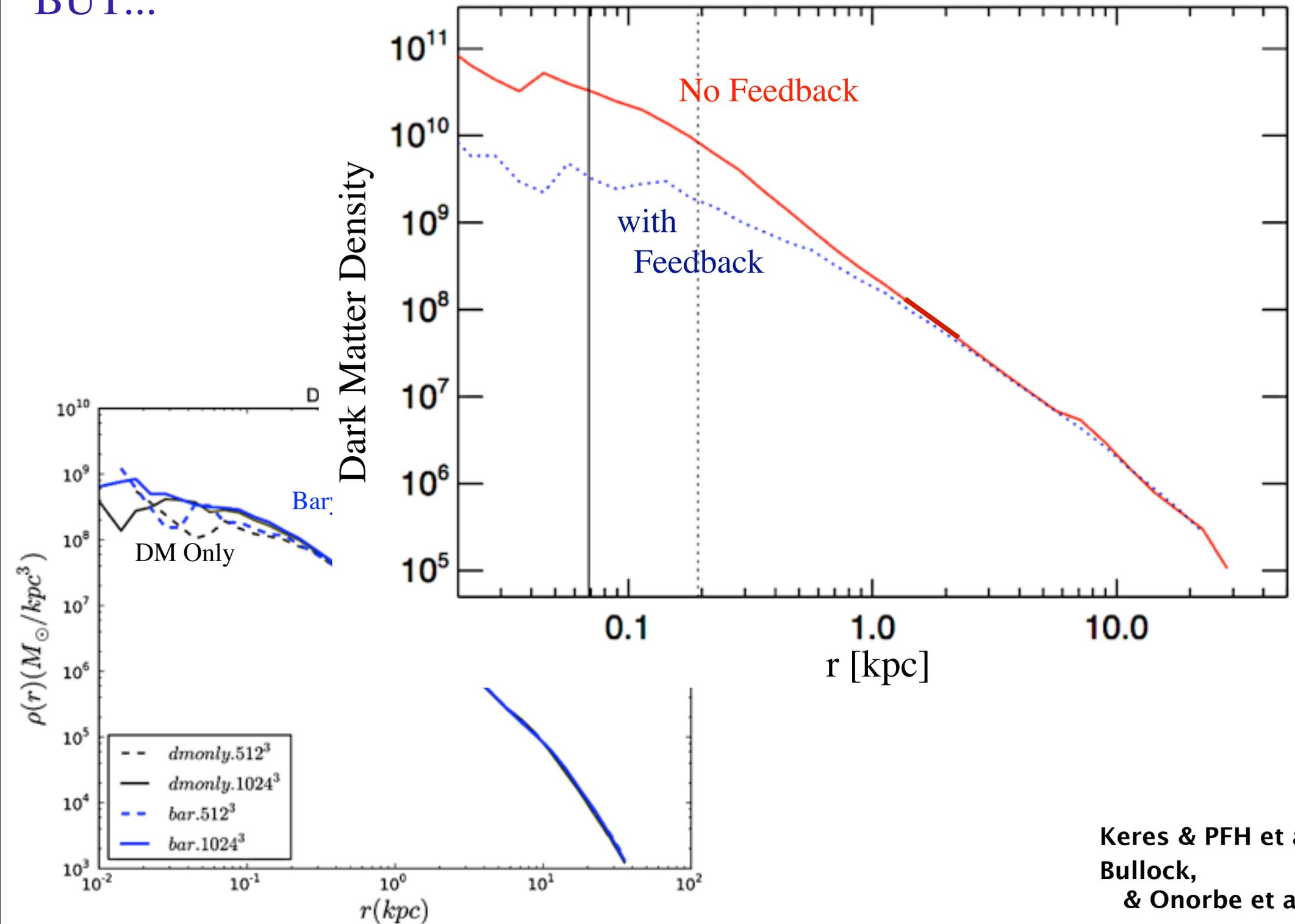
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BUT...



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BUT...



Summary:

- **Star formation is Feedback-Regulated:** *independent* of small-scale SF ‘law’
 - Need enough stars to offset dissipation (gravity)
 - Leads to Kennicutt relation & **super-winds**
- Different mechanisms dominate different regimes:
 - High- r : radiation pressure
 - Intermediate: HII heating, stellar wind momentum
 - Low- r : SNe & stellar wind shock-heating
 - **No *one* mechanism works**
- Cosmologically: *Not* just top-down inflows:
 - Winds determine **IGM enrichment, temperature, & subsequent inflow** structure
 - Cores? Be VERY careful!