Numerics,Physics,Resolution

Towards Predictive Galaxy Formation Simulations

Phil Hopkins
Caltech

Physics

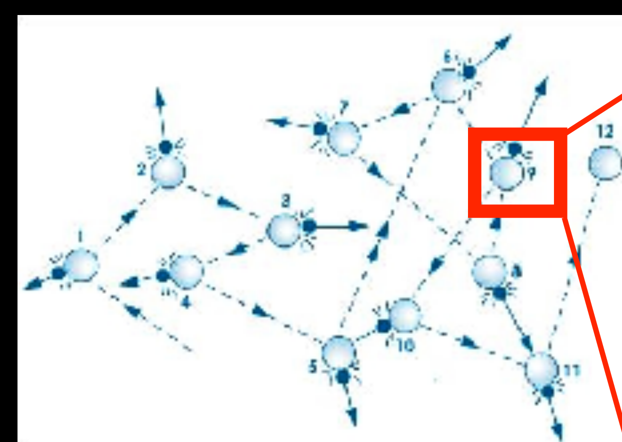
(a question of philosophy)

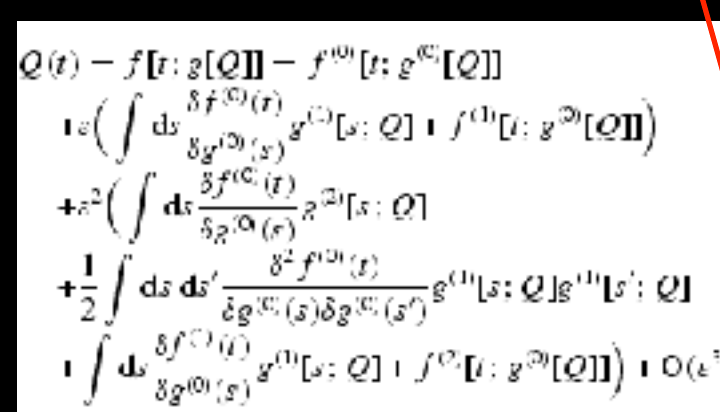
Everything is sub-grid

Hydrodynamics

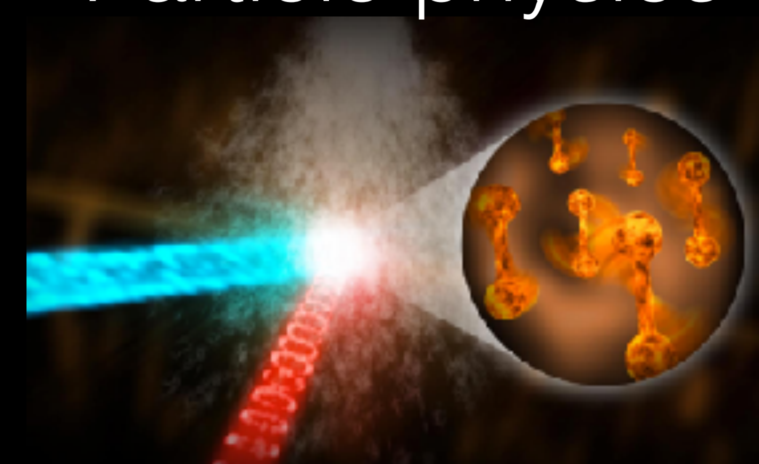
$$\frac{D\rho}{Dt} = -\rho \nabla \cdot \mathbf{v}$$

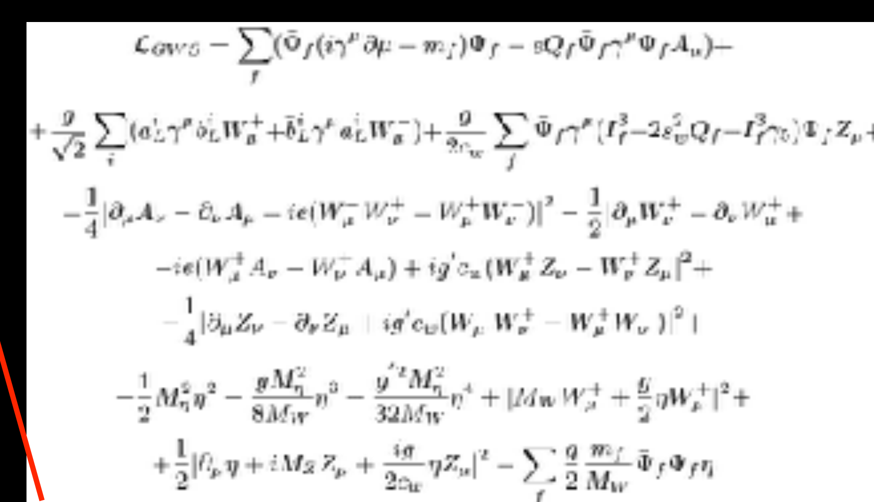
Statistical mechanics





Particle physics





2 philosophies of sub-grid:

- 1. Parameterize unknowns, marginalize over them (fit to observations)
    - bias in BAO/LSS cosmology
    - MCMC SAMs / Illustris/Eagle philosophy
- 2. Derive from theory/observations on small scales, after “smoothing”
    - (magneto) hydrodynamics
    - FIRE philosophy: $M_{\rm wind}$ = (whatever the input physics predicts)

Example: Supernovae

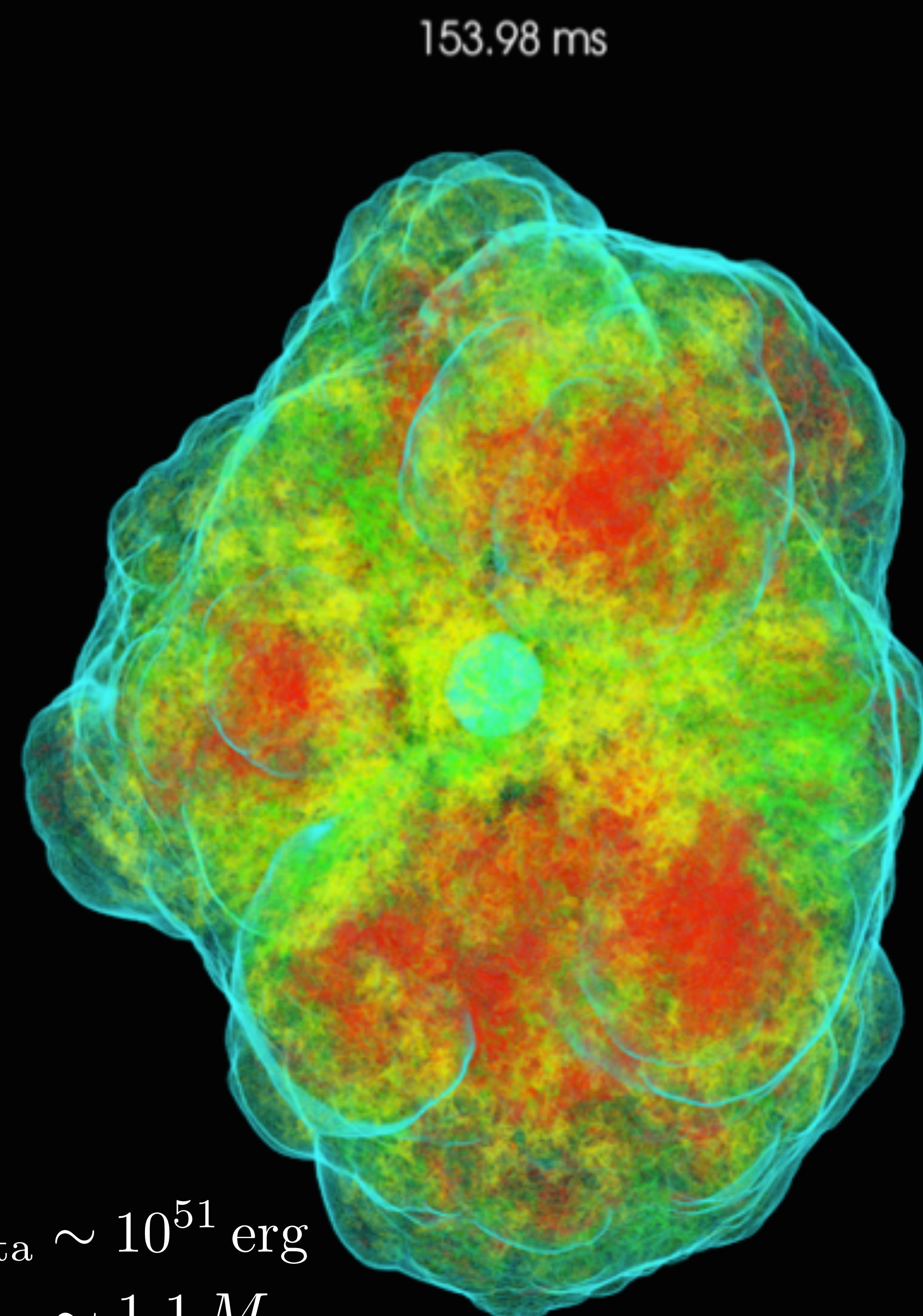
(building up a sub-grid model)

Example: SNe

Predict: Explosion

Resolution:

$$m_i < 10^{-6} \, M_\odot$$



$\mathrm{KE}_{\mathrm{ejecta}} \sim 10^{51}\,\mathrm{erg}$

$M_{\mathrm{Oxygen}} \sim 1.1\, M_\odot$

$\dots$

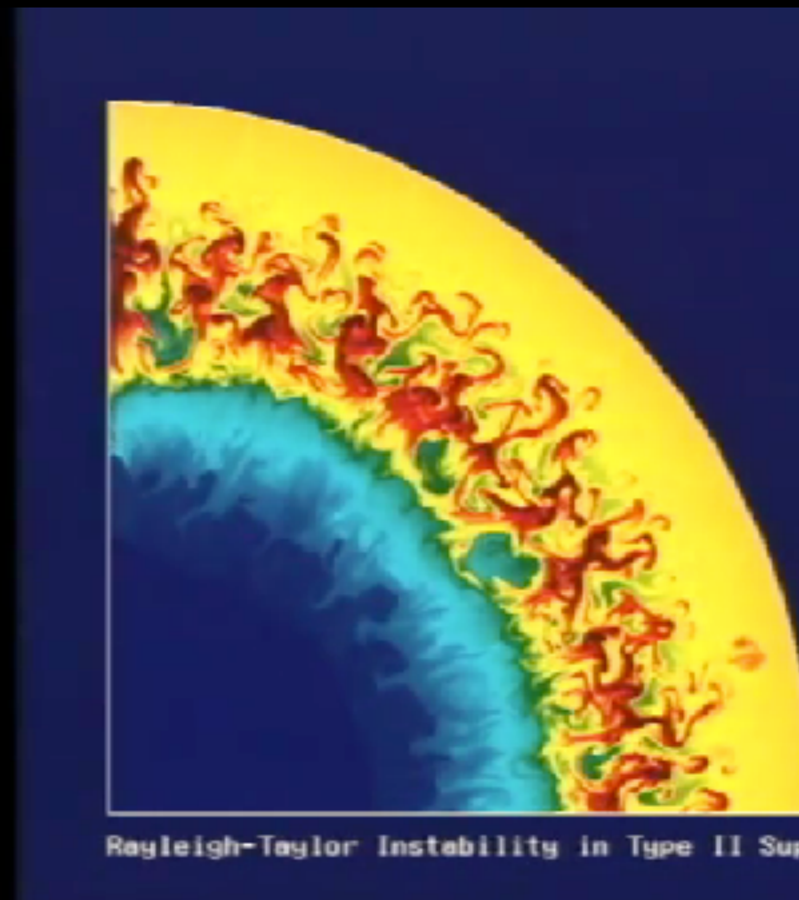
Moesta et al. '14

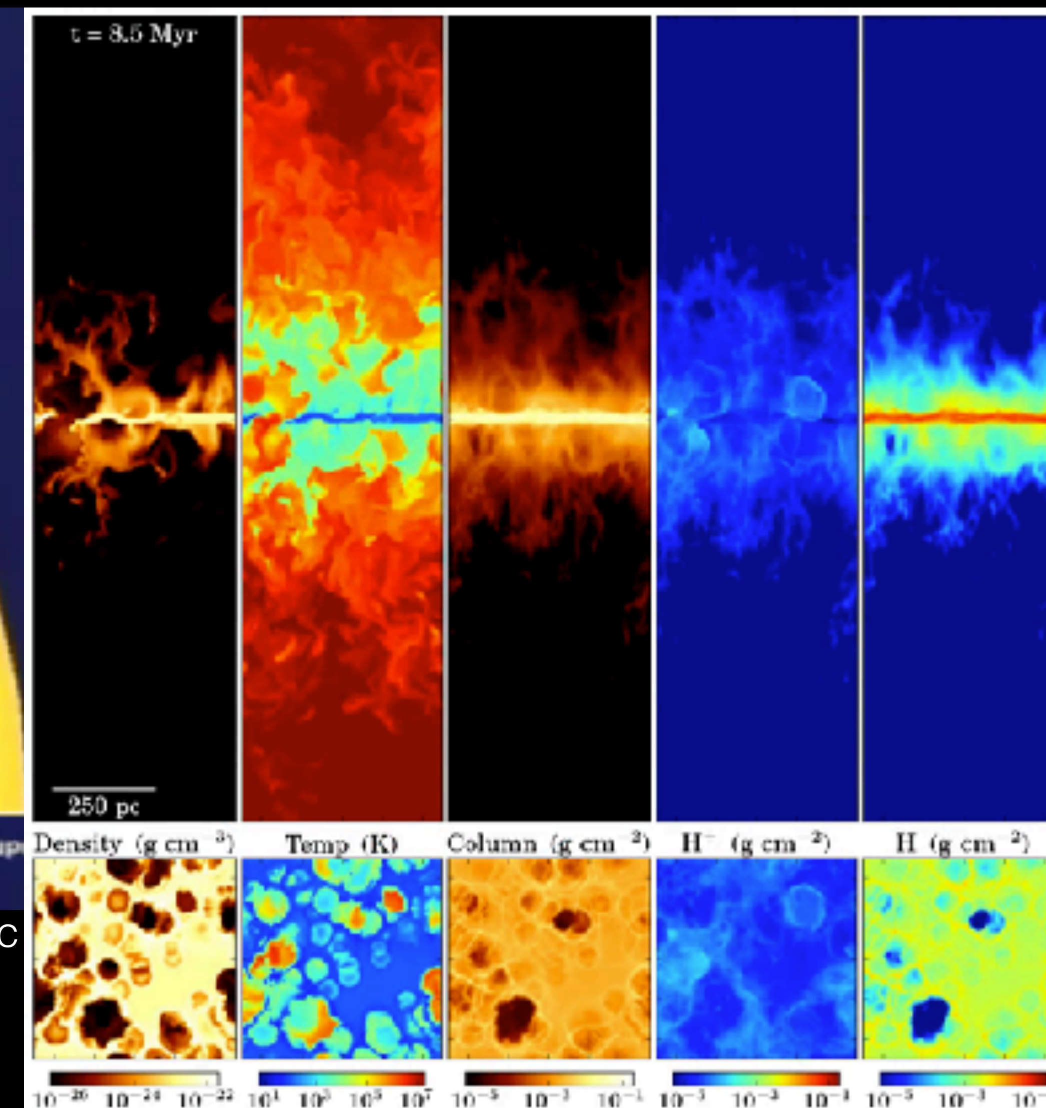
Sub-grid physics:

- (magneto) hydrodynamics
- nuclear Rx rates
- neutrino transfer

Example: SNe

Predict: Blastwave Evolution/ISM Interaction





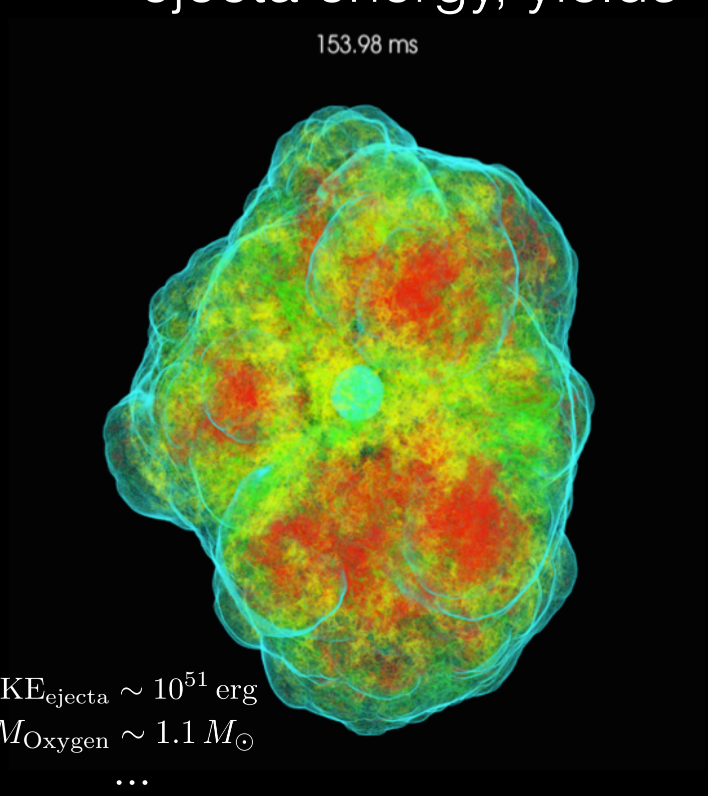
Walch, Martizzi, Barnes, Cioffi, etc

Resolution:

$$m_i \sim 1 - 100\, M_\odot$$

Sub-grid physics:

- SNe explosion
- ejecta energy, yields



End of energy-to-momentum (single SNe):

$$M_{\mathrm{snowplow,\,final}} \sim 3000\, M_\odot$$

Final momentum:

$$\langle M_s\, v_s \rangle_{\mathrm{final,\,SNr}} \sim 10^{5.5}\, M_\odot\, \frac{\mathrm{km}}{\mathrm{s}}$$

Example: SNe

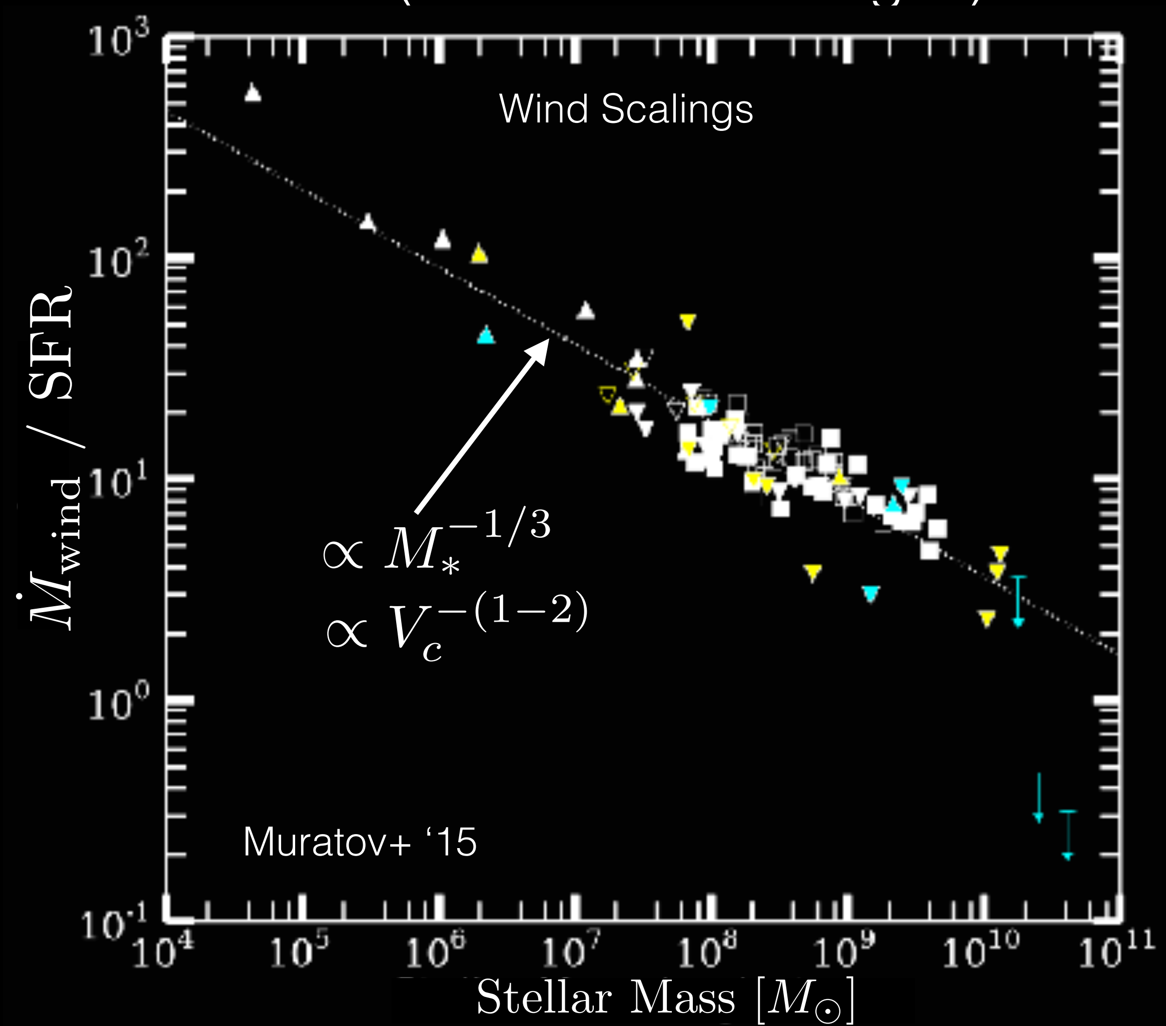
Predict: Overlap: super-bubbles & winds

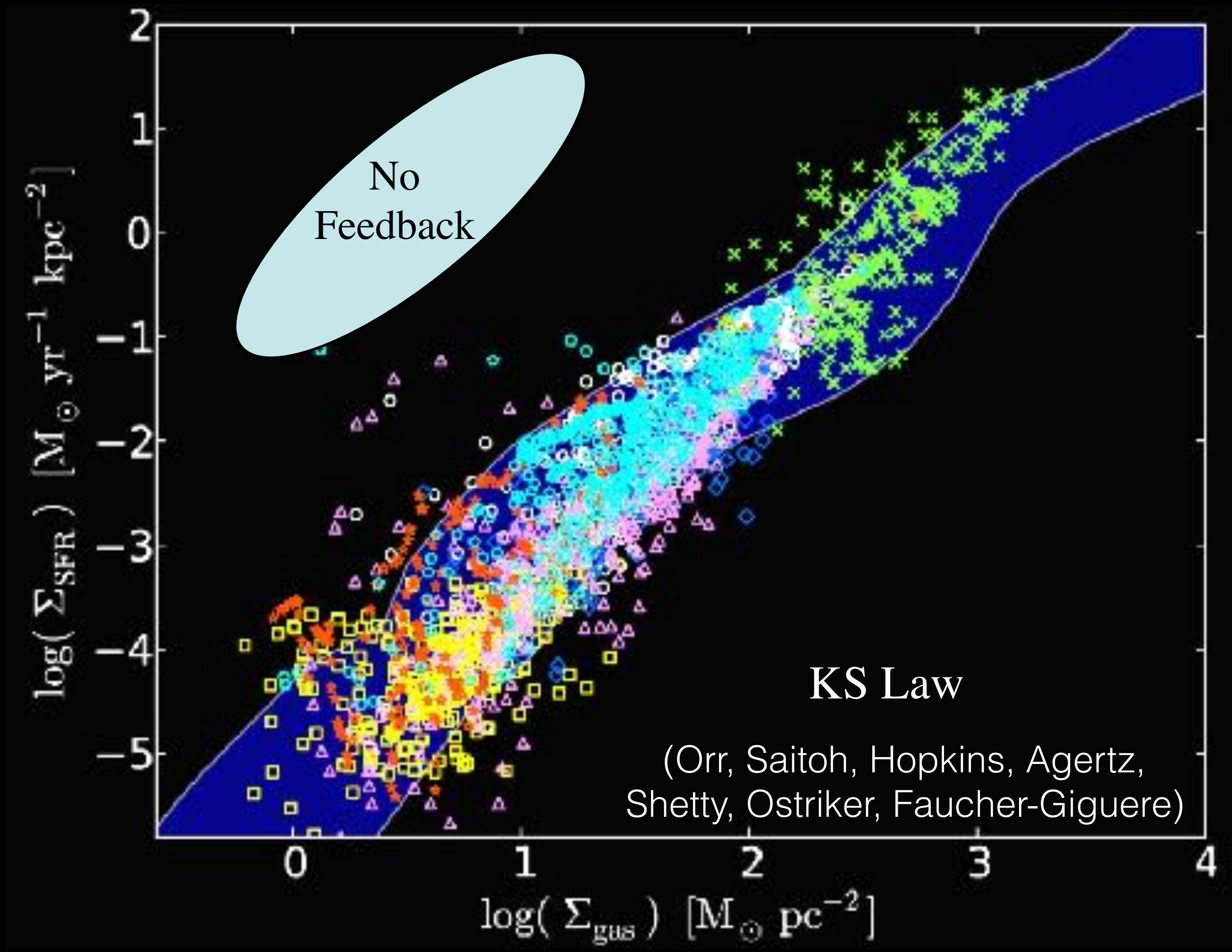
Resolution:

$$m_i \sim 10^{2-4}\, M_\odot$$

Sub-grid physics:

- single SNr evolution
- stellar evolution (rates)
- ~~SFR (dense molecular gas)~~

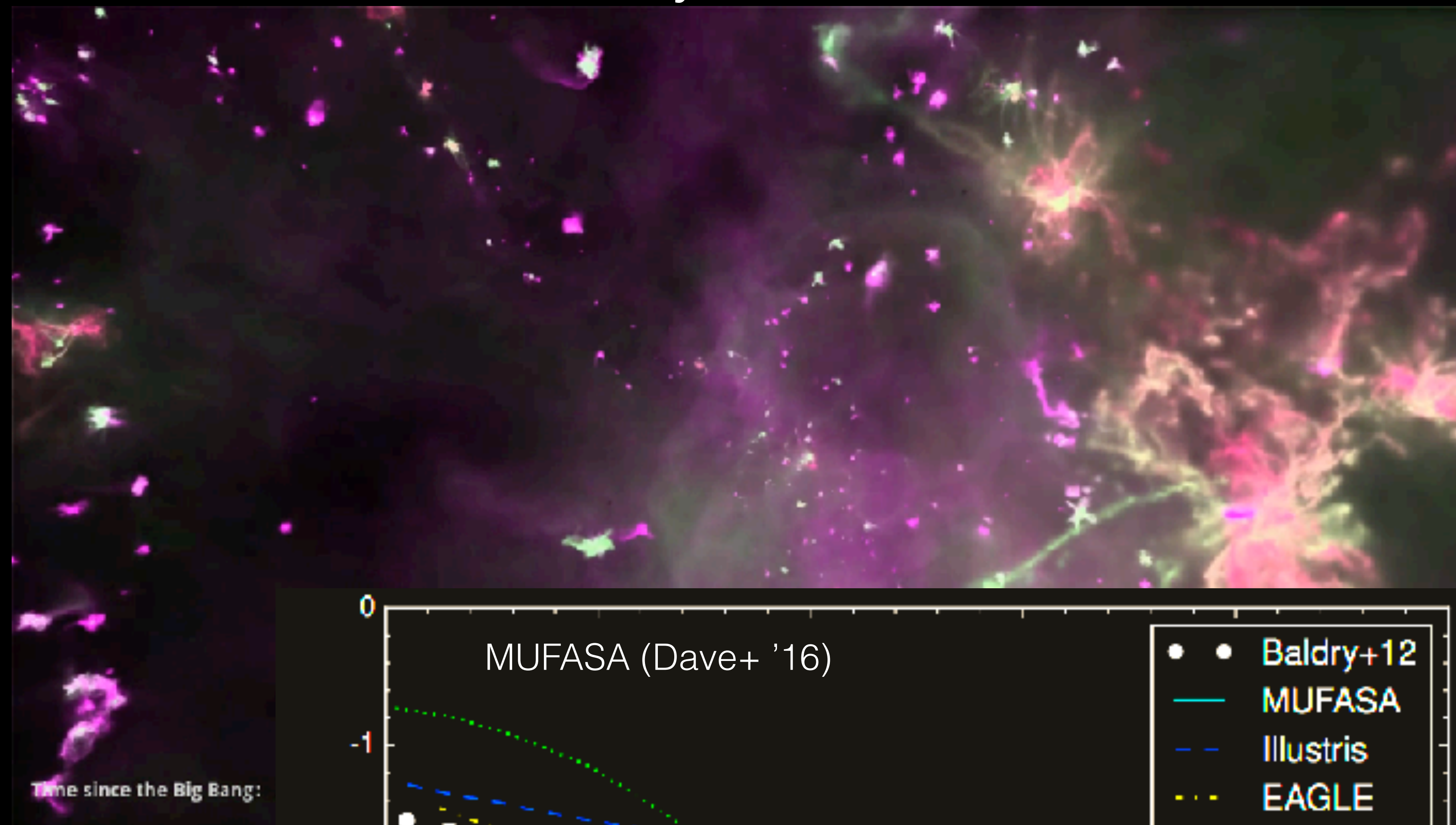




FIRE

Example: SNe

Predict: Galaxy SFHs, IGM enrichment

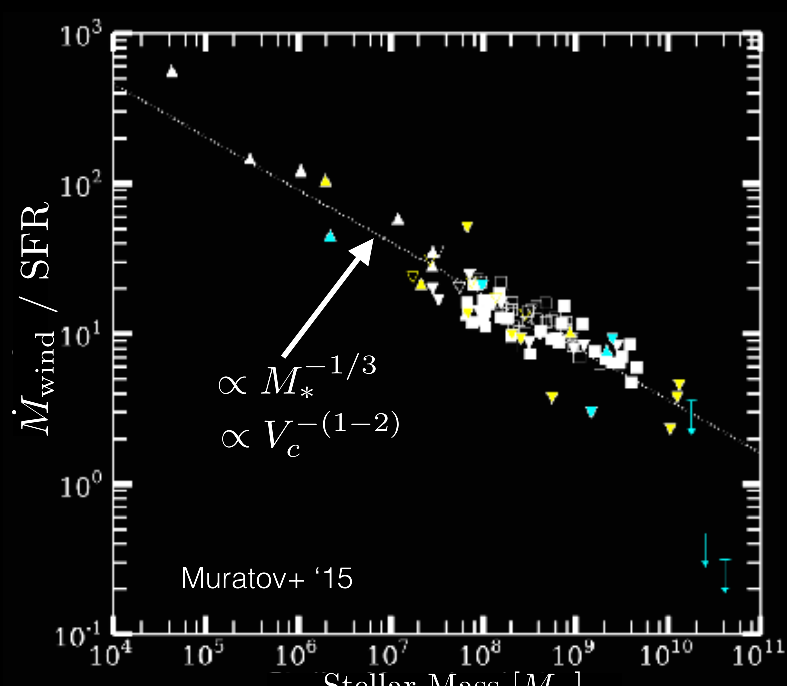


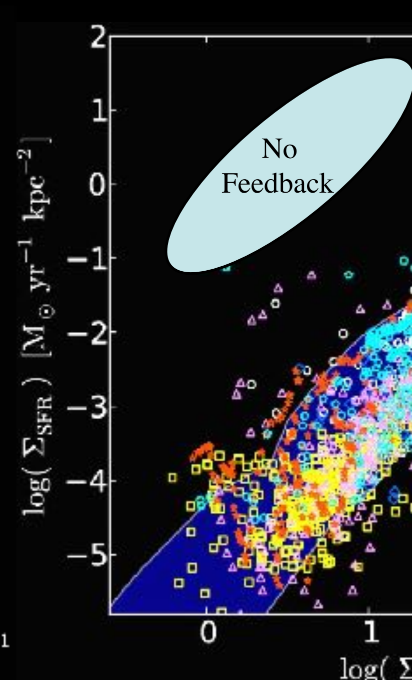
Resolution:

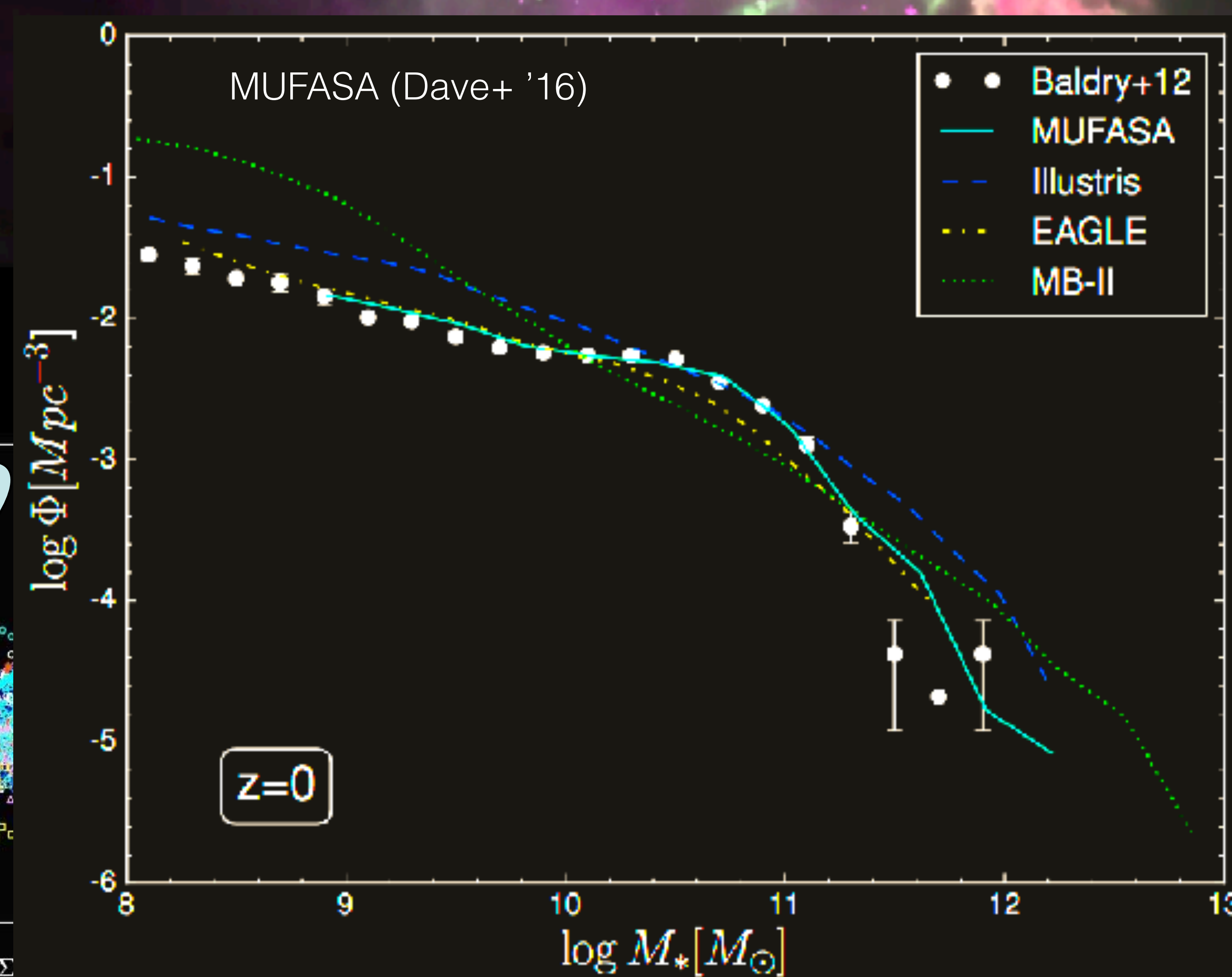
$$m_i \gtrsim 10^6\ M_\odot$$

Sub-grid physics:

- ~~SFR (kpc/low-density gas)~~
- wind scalings (galaxy-scale)



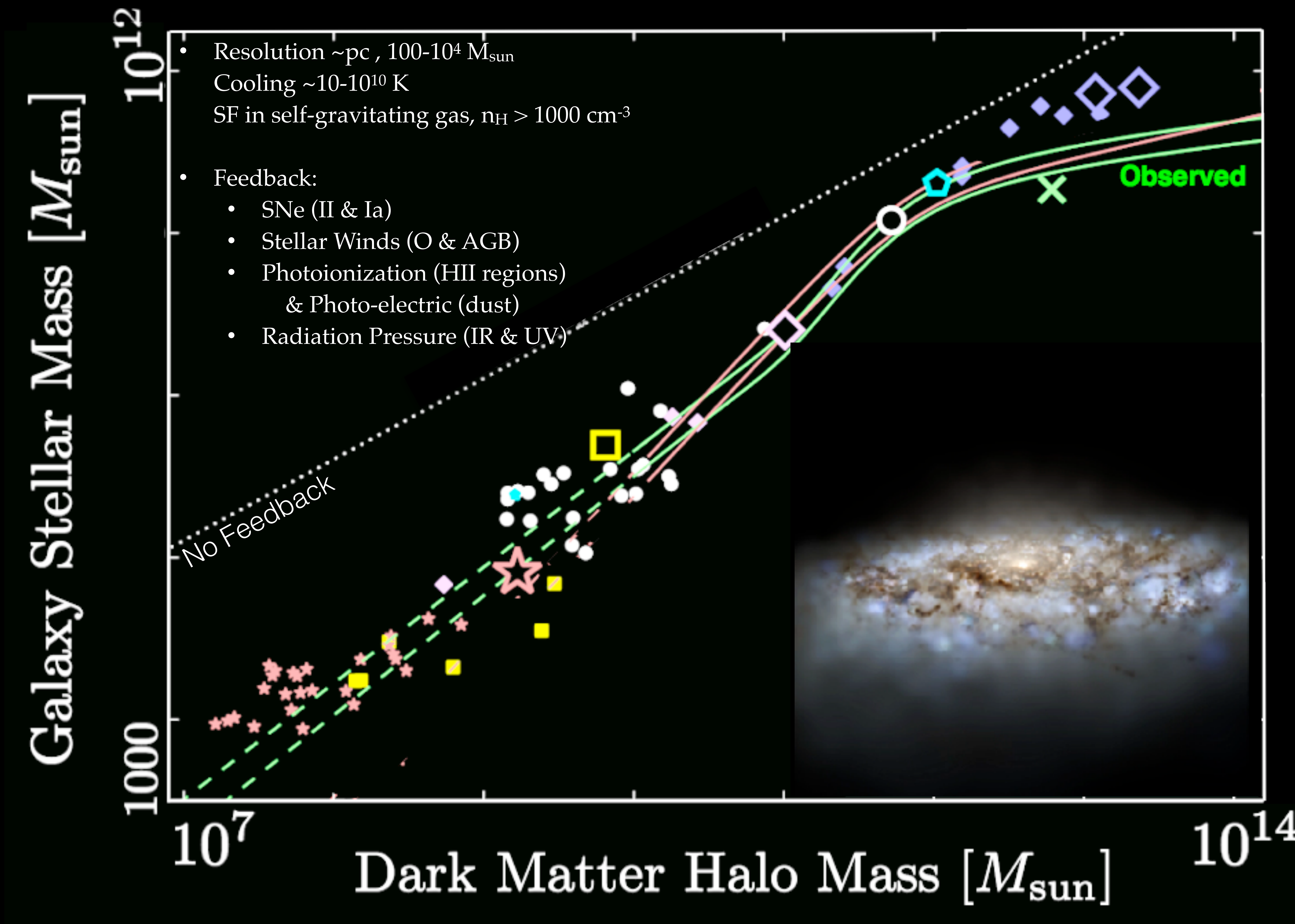




It Works!

THIS APPROACH IS PRODUCING REALISTIC GALAXIES

PFH et al.
(arXiv:1311.2073)



It Works!

- Resolution ~pc , 100-10^4 M_{sun}
  Cooling ~10-10^{10} K
  SF in self-gravitating gas, $n_H > 1000$

- Feedback:
  - SNe (II & Ia)
  - Stellar Winds (O & AGB)
  - Photoionization (HII regions) & Photo-electric (dust)
  - Radiation Pressure (IR & UV)

What Matters?

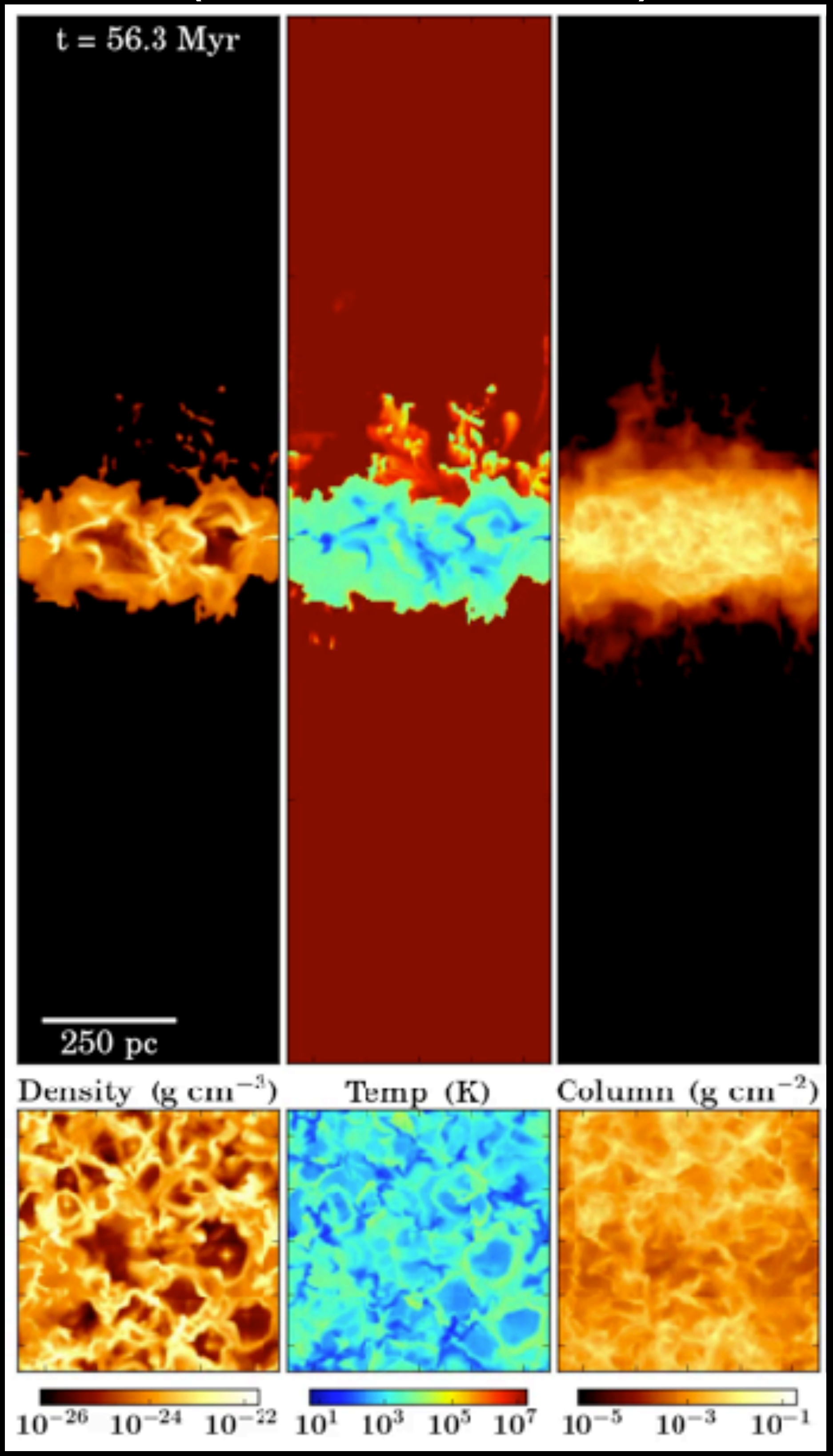
(depends 100% on *what you care about predicting*)

Doing the "sub-grid" right can matter

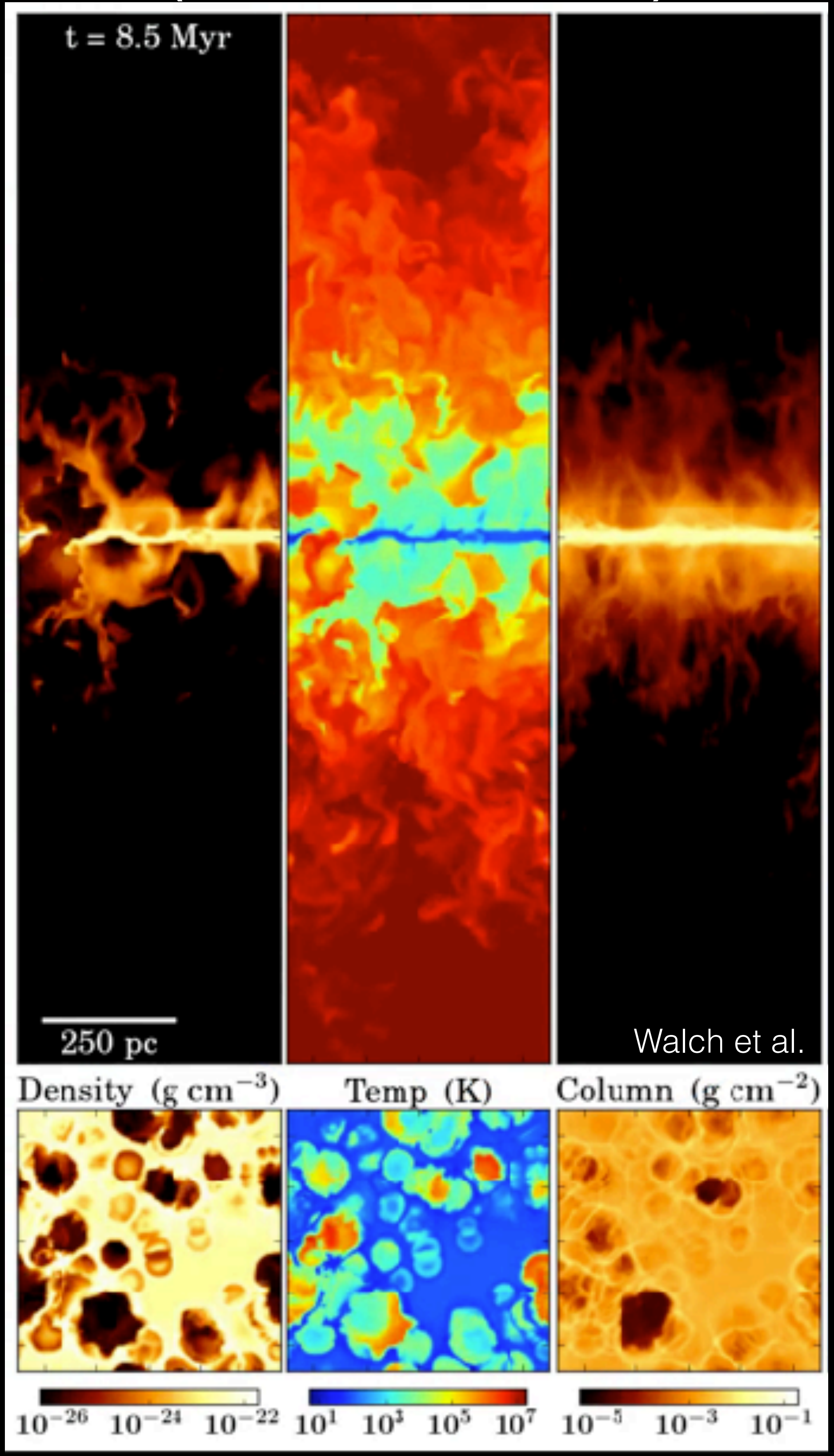
IF RESOLVE DENSE GAS, NEED PHYSICS FOR IT!

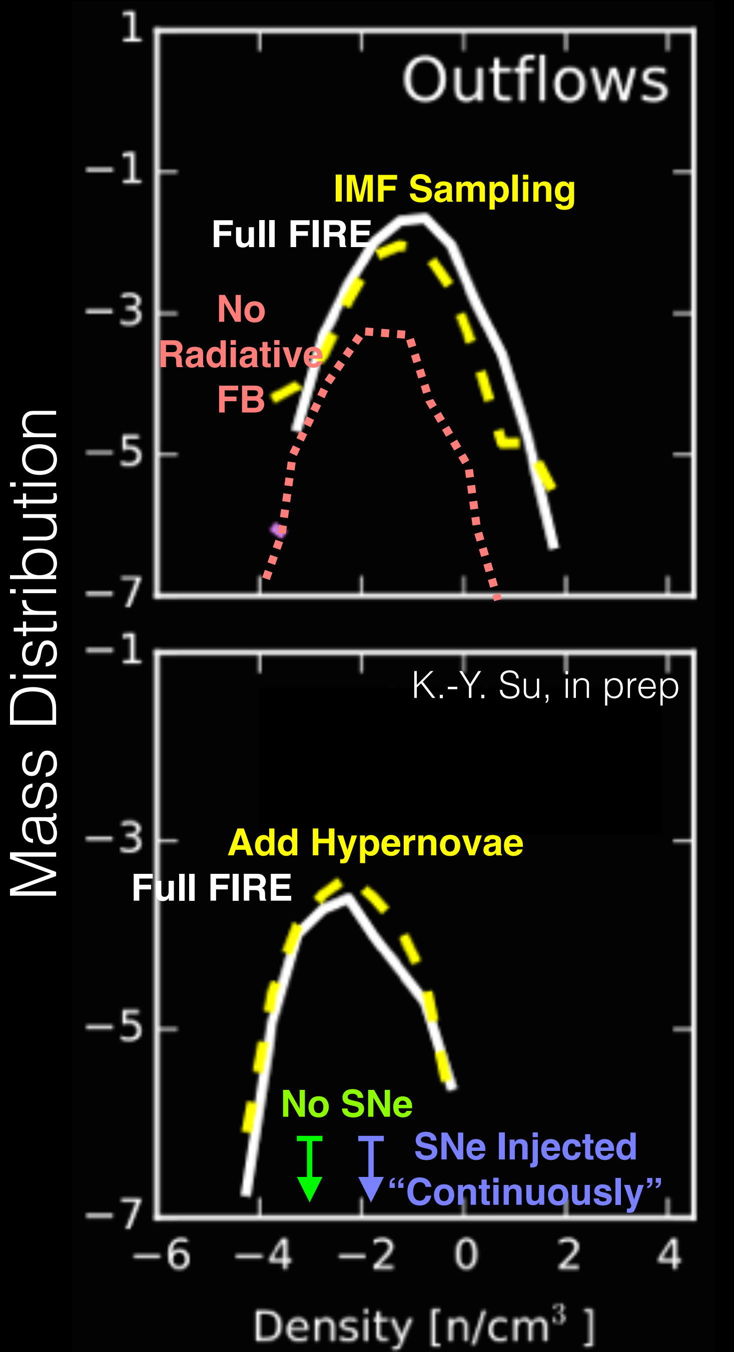
Murray+, Martizzi+, Walch+, Barnes+ Hopkins+, Hayward+, Shetty+, Hennebelle+

SNe Explode in Density Peaks (no radiative feedback)



SNe Clustered & Off-Peak (with radiative feedback)

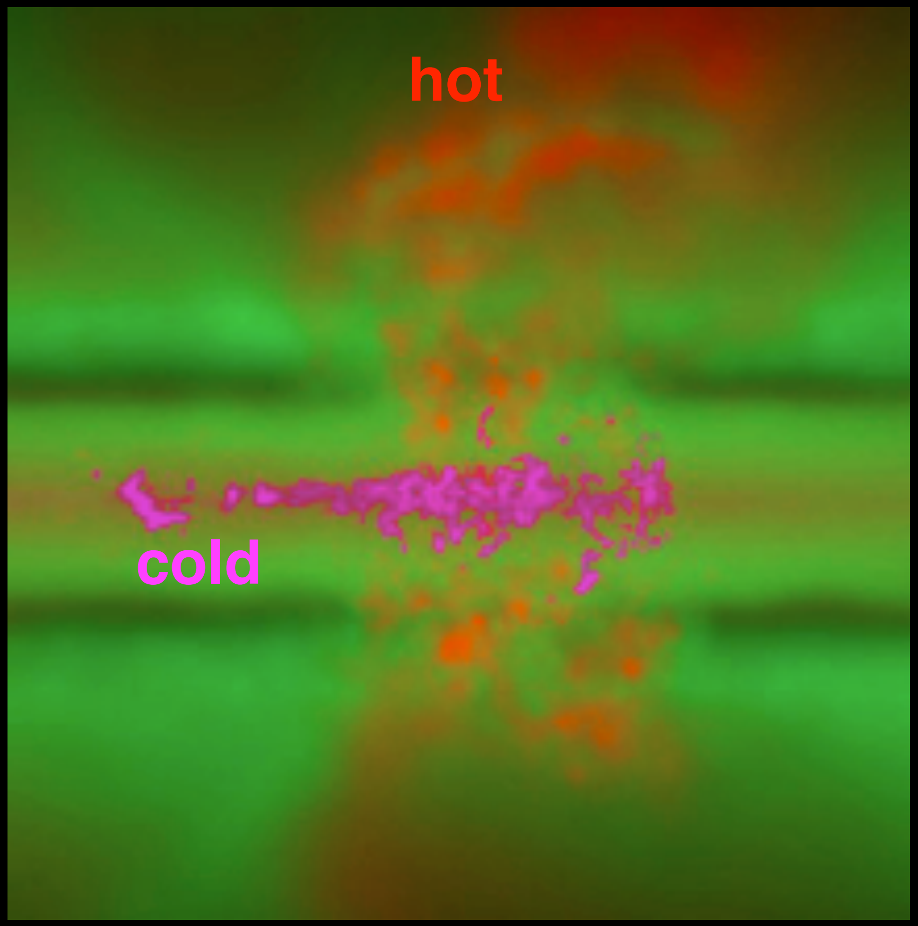




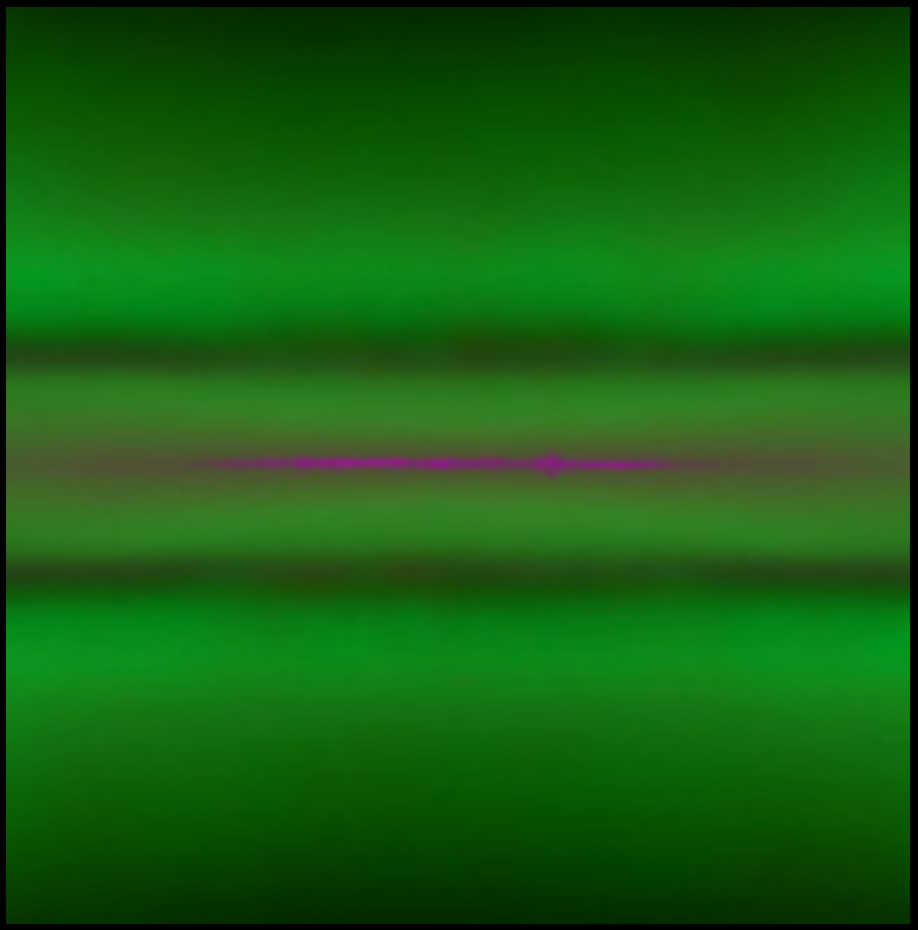
Doing the "sub-grid" right can matter
IF RESOLVE BUBBLES, NEED PHYSICS FOR IT!

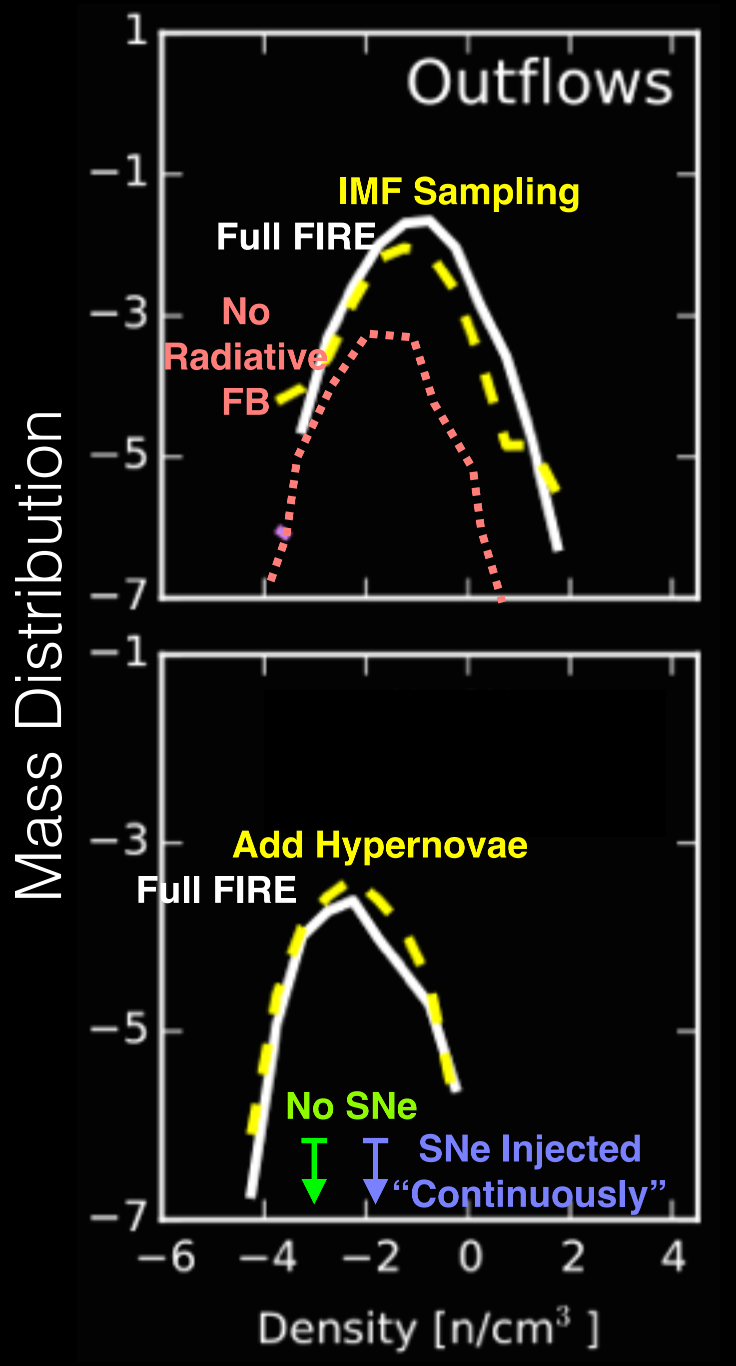
Klessen+, Ostriker+
Hopkins+
K.-Y. Su, in prep

Treat each SNe explicitly following resolved explosion



Continuously dump thermal energy ~SFR





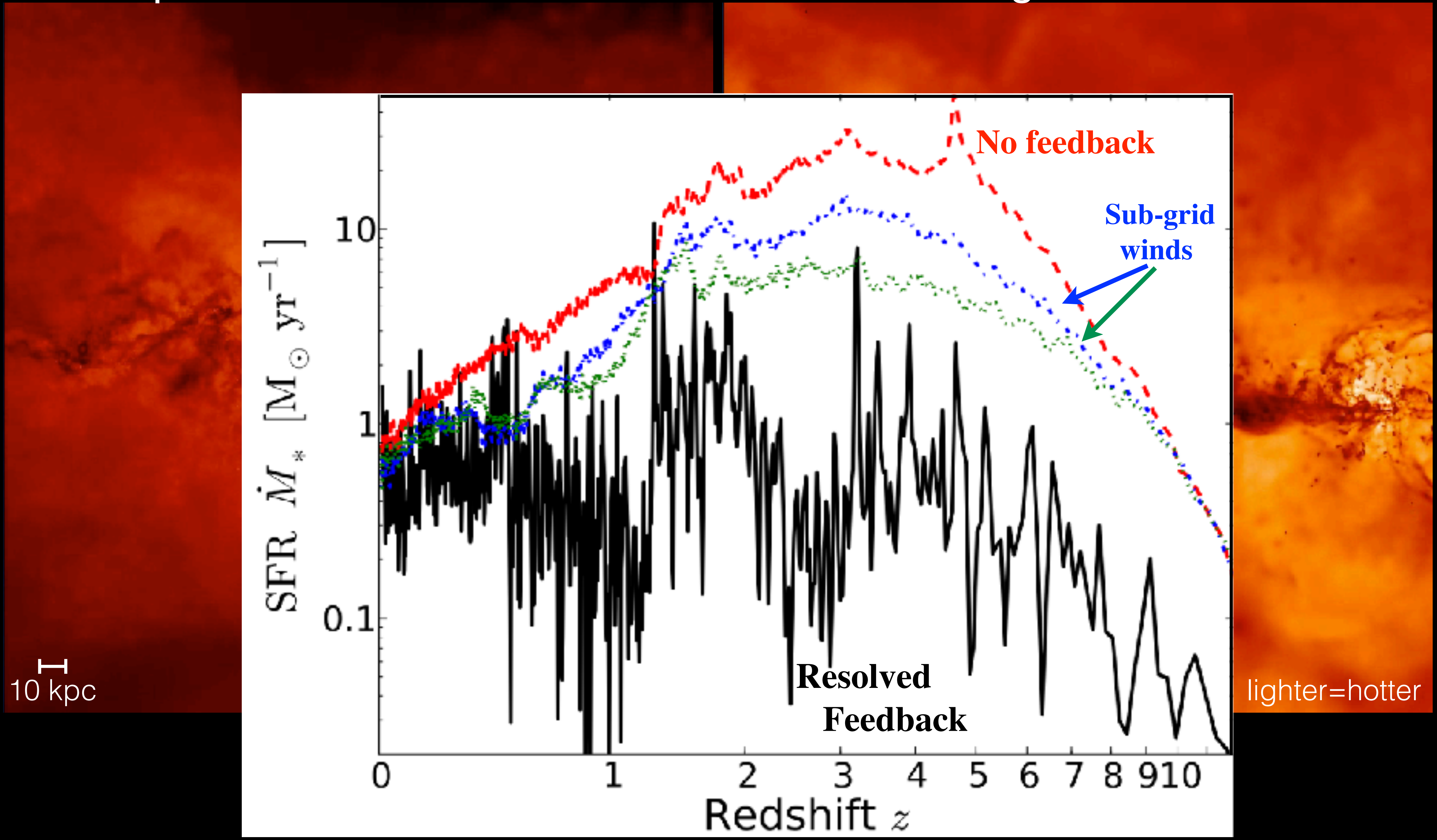
Doing the "sub-grid" right can matter

DANGERS OF ONLY FITTING MASSES

M. Sparre
arxiv:1510.03869

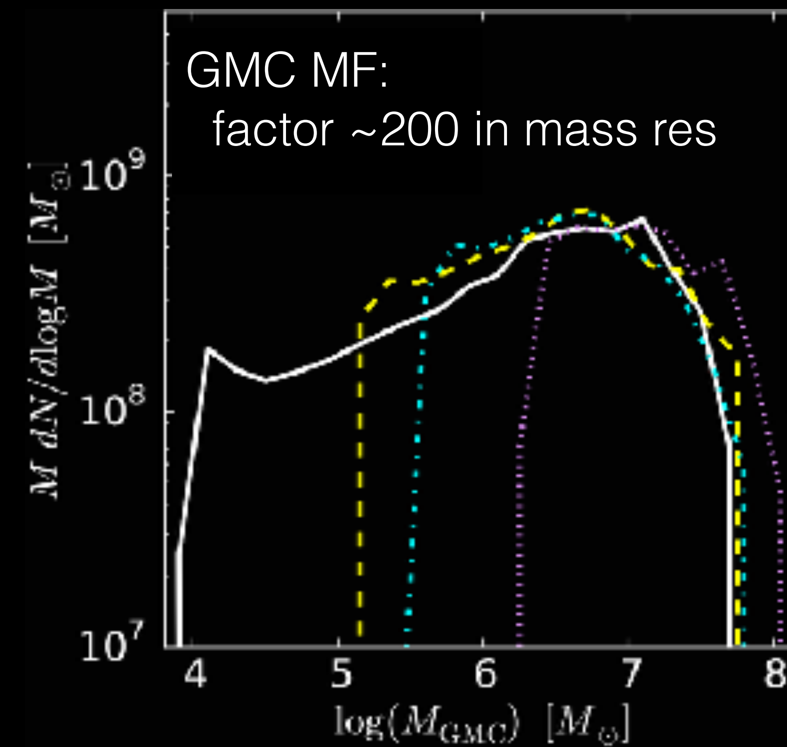
Proto-Milky Way: Gas Temperature:

Simple Sub-Grid $(\dot{M}_{\rm wind} = \eta\, \dot{M}_*)$ | Following Full Feedback



Resolution: Needs to Match Your Physics!

DIFFERENT PREDICTIONS REQUIRE DIFFERENT RESOLUTION



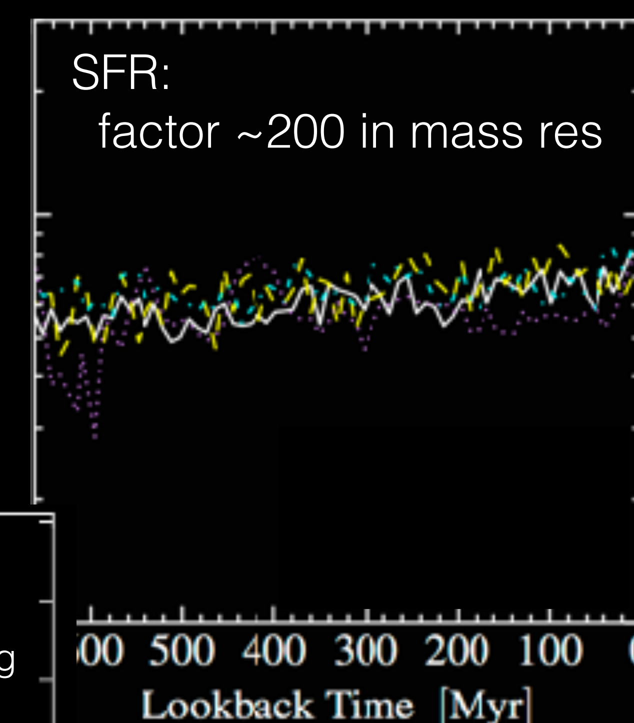
Fragmentation / GMCs / Dense Gas:

$$m_i \lesssim 10^5\,M_\odot \ll M_{\rm Toomre}$$

$$\epsilon_{\rm grav}^{\rm min} \ll 100\,{\rm pc}$$ [guaranteed if adaptive]

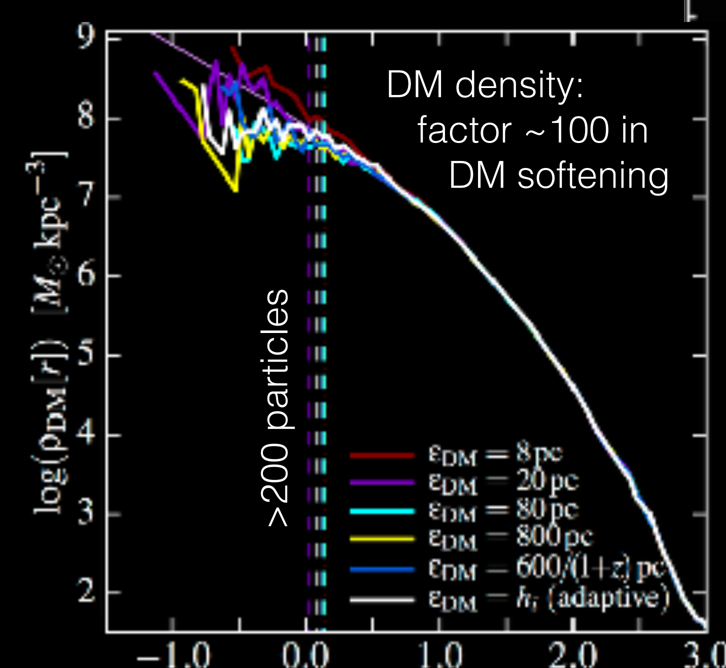
Super-bubbles / overlaps / chimneys:

$$m_i \lesssim 10^5\,M_\odot \ll M_{\rm Bubble}$$



Individual SNe (no sub-grid SNe momentum):

$$m_i \lesssim 10^3\,M_\odot \ll M_{\rm Cooling}$$



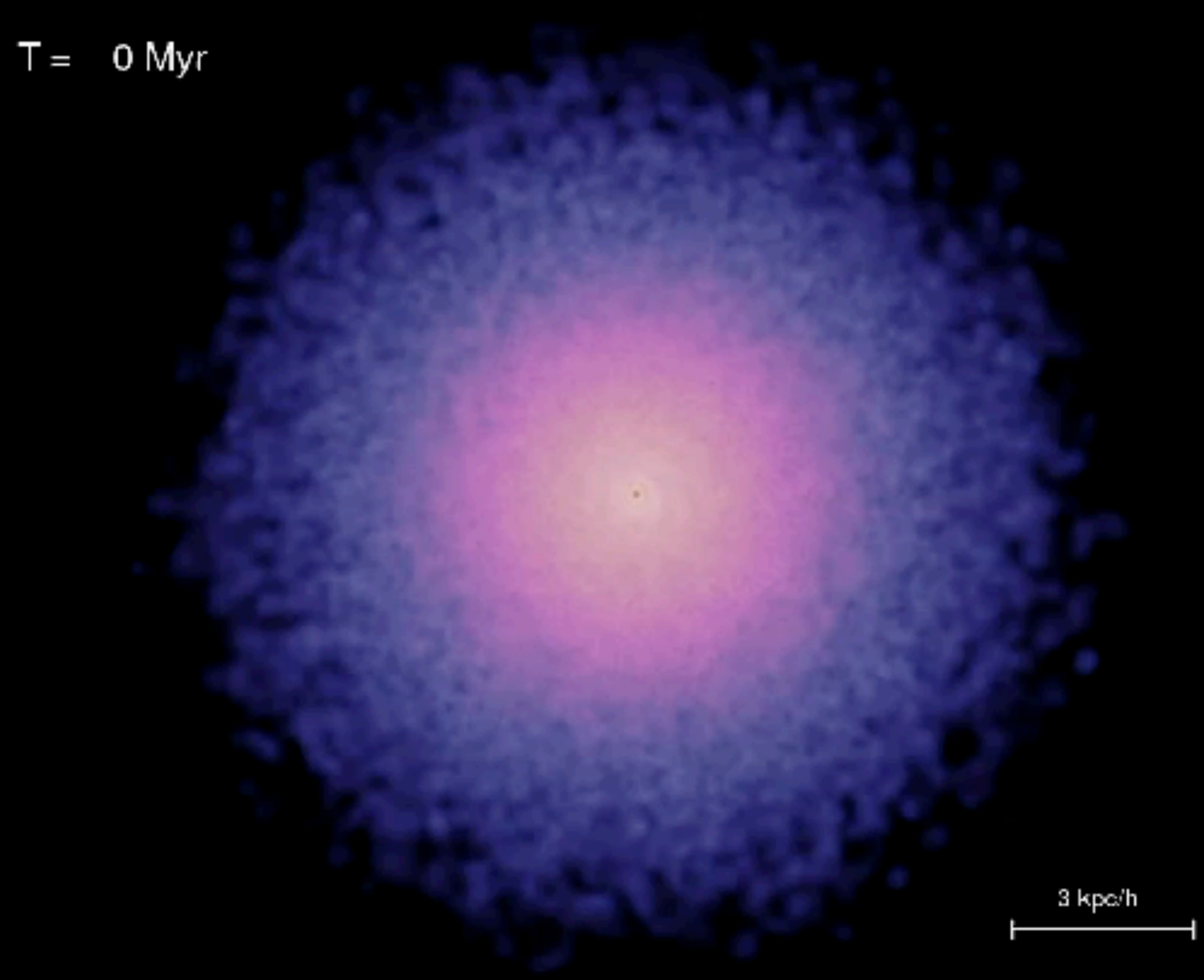
Dwarf galaxy "bursty-ness":

$$m_i \lesssim 10^{-6}\,M_{\rm halo}$$

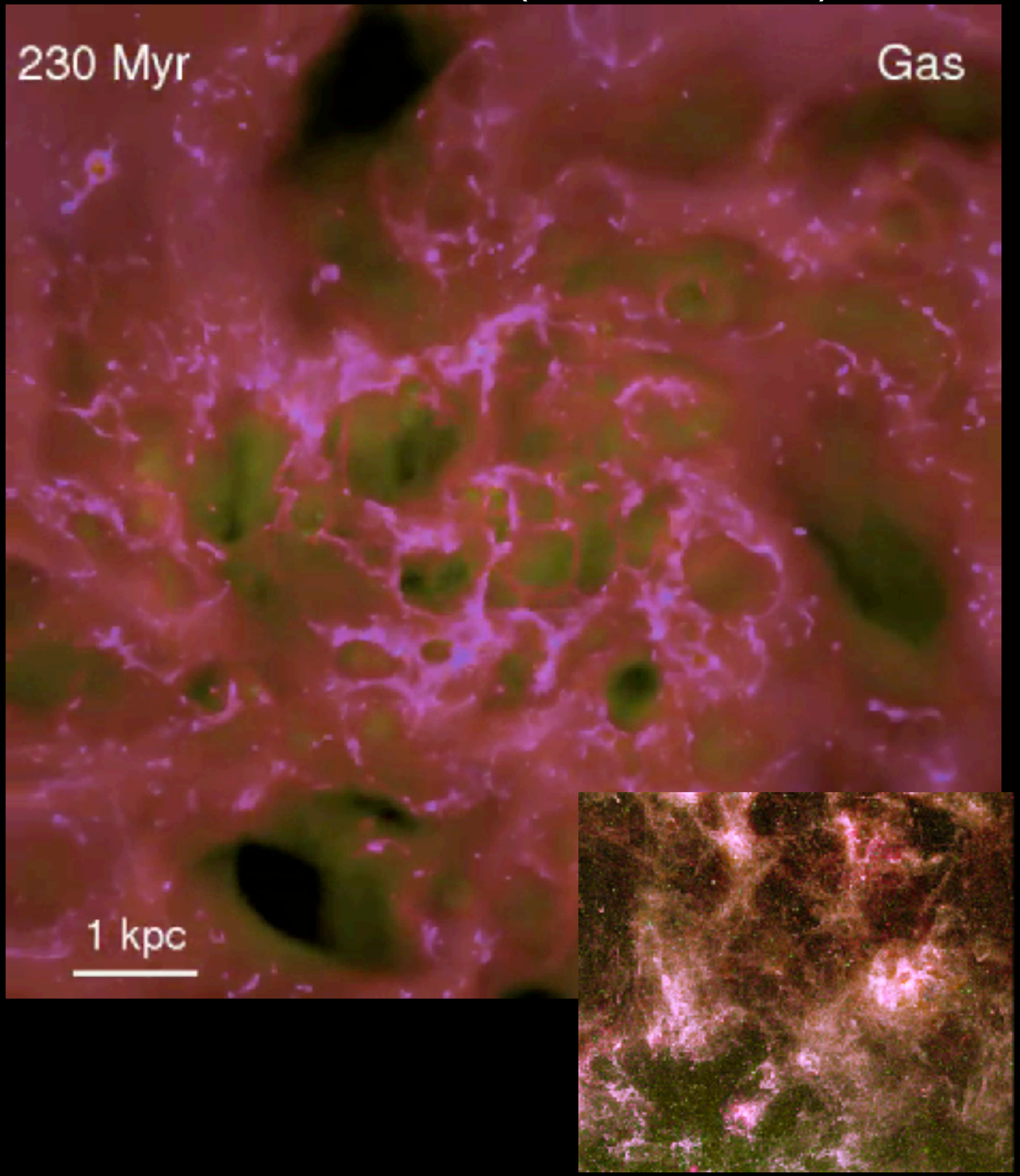
Doing the "sub-grid" right can matter

NEED PHYSICS TO PUSH BEYOND YOUR SUB-GRID SCALE

Sub-Grid ISM (Illustris, Eagle)



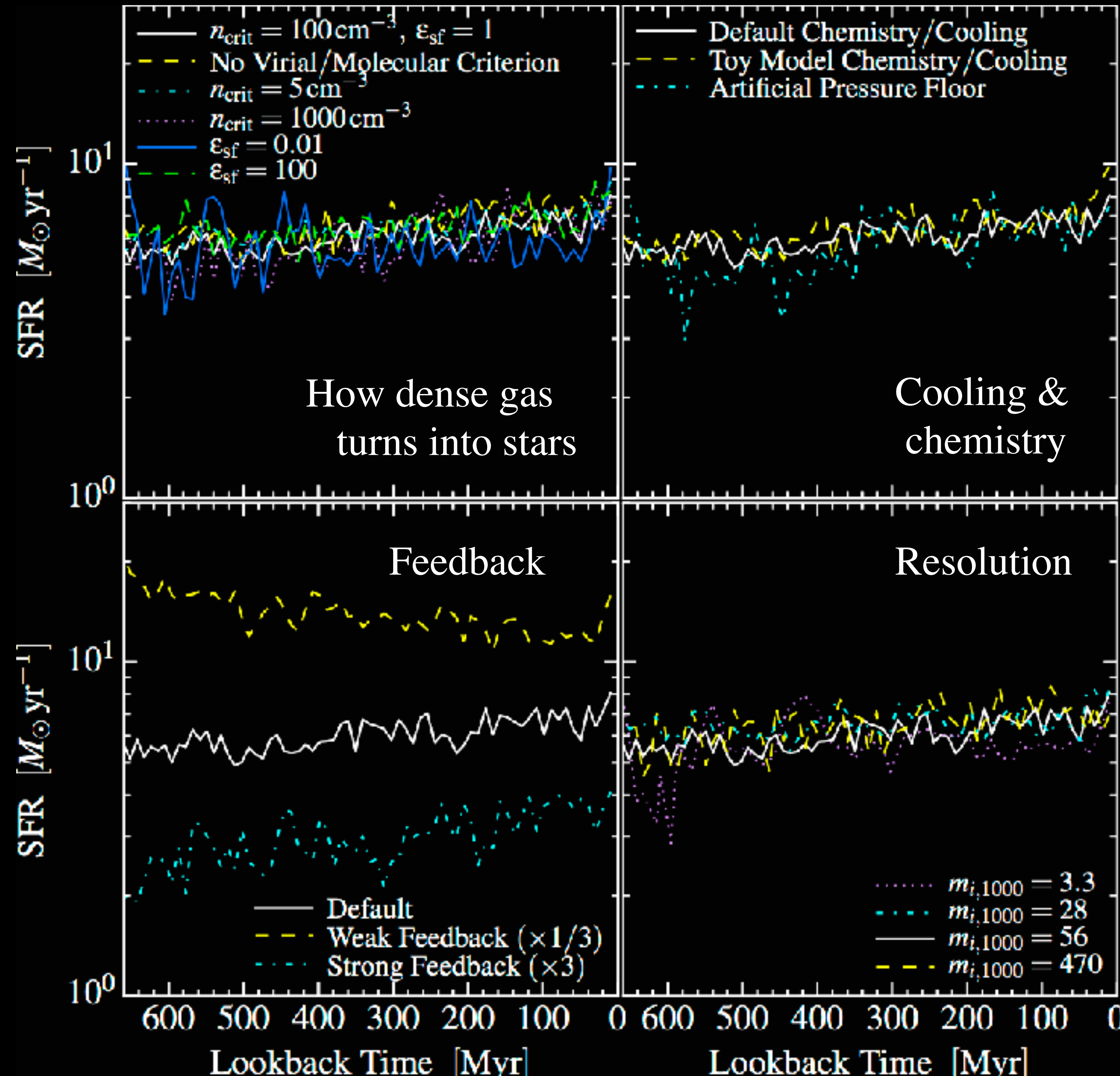
Resolved ISM (FIRE, SILCC)



What Doesn’t Matter?

(depends 100% on *what you care about predicting*)

(Galactic) Star Formation Rates are *INDEPENDENT* of how stars form!

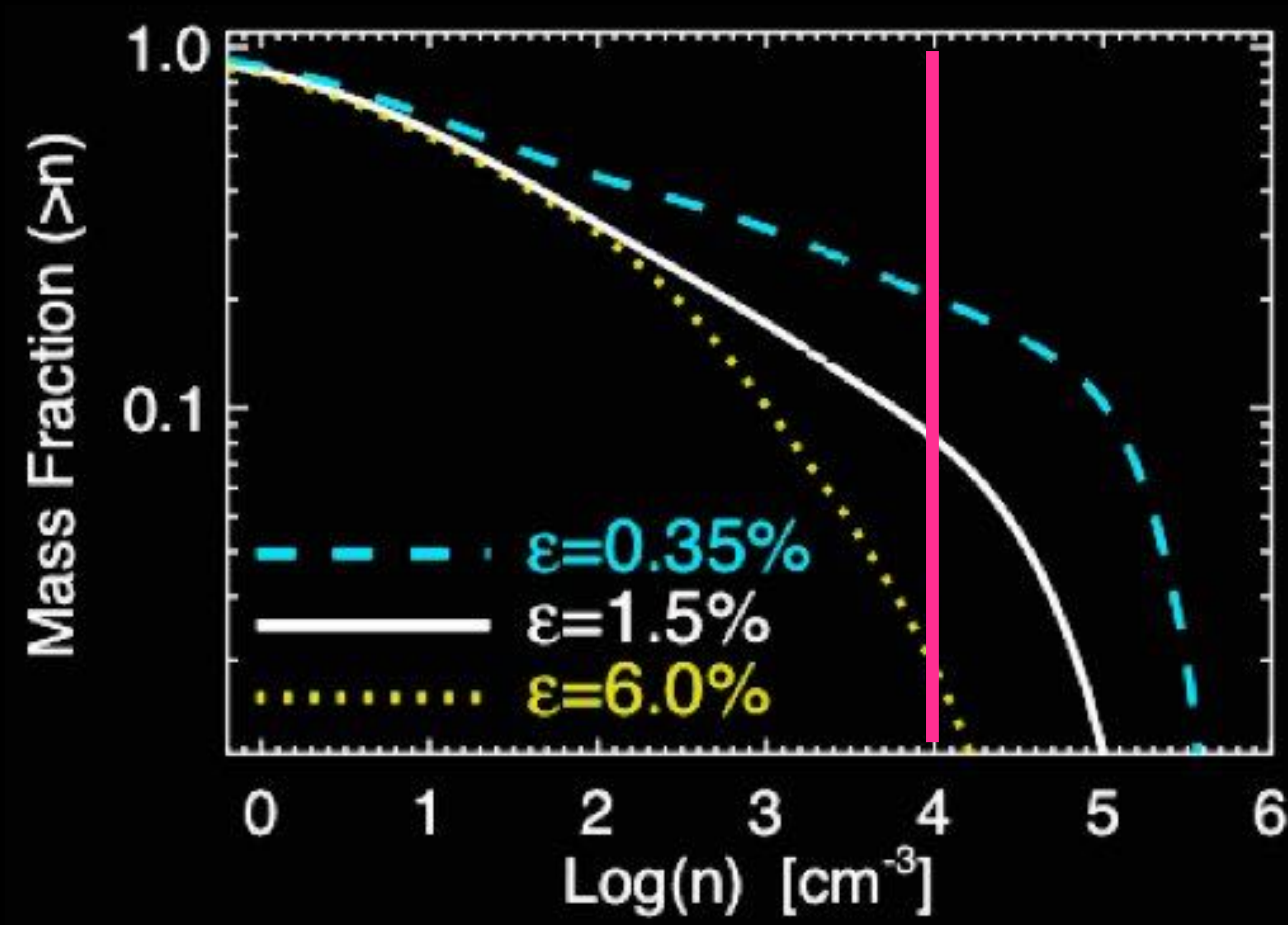


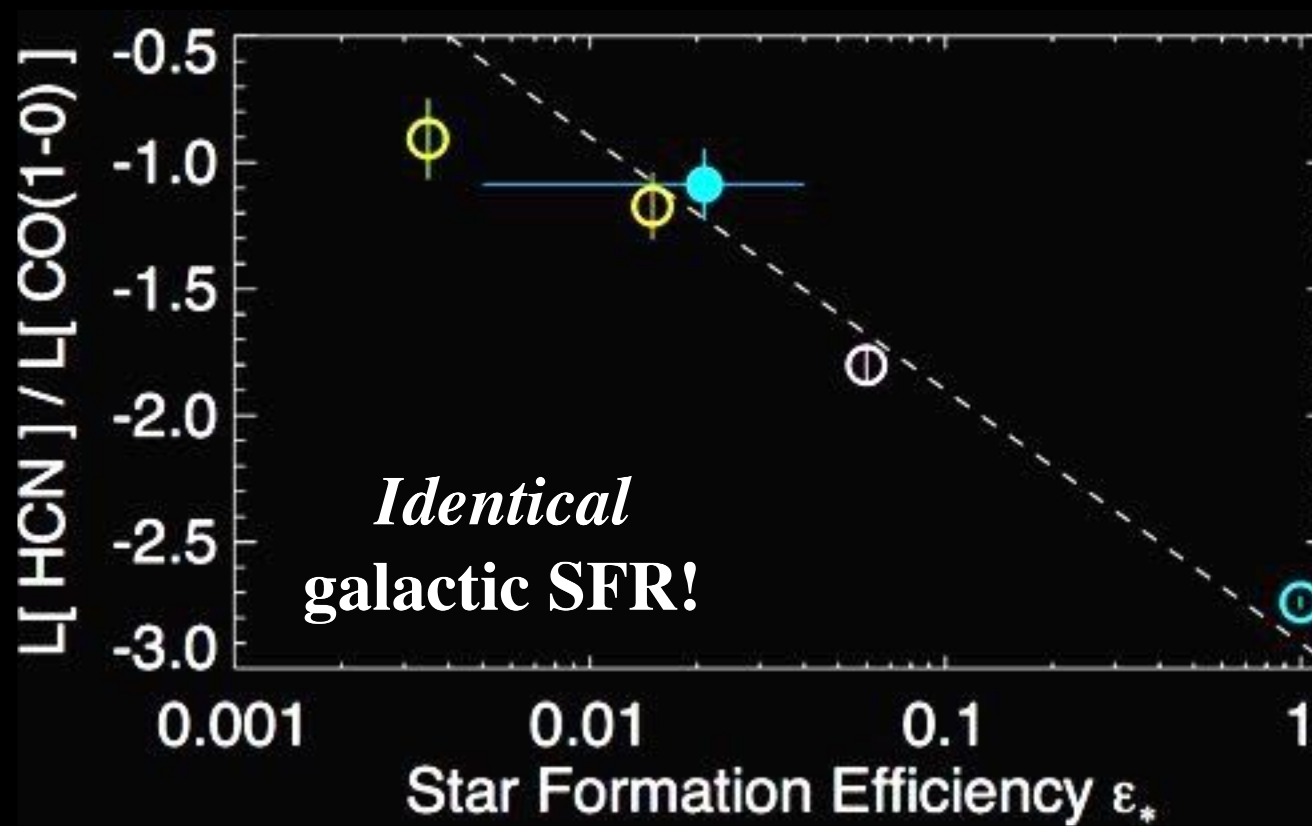
Matt Orr (in prep)
Saitoh+ 11
Hopkins+ 11,12,14
Agertz+14

Dense Gas *Does* Change

SELF-REGULATES TO "NEEDED" SFR LEVEL

Efficiency (SF per t_{dyn}) in *dense* gas





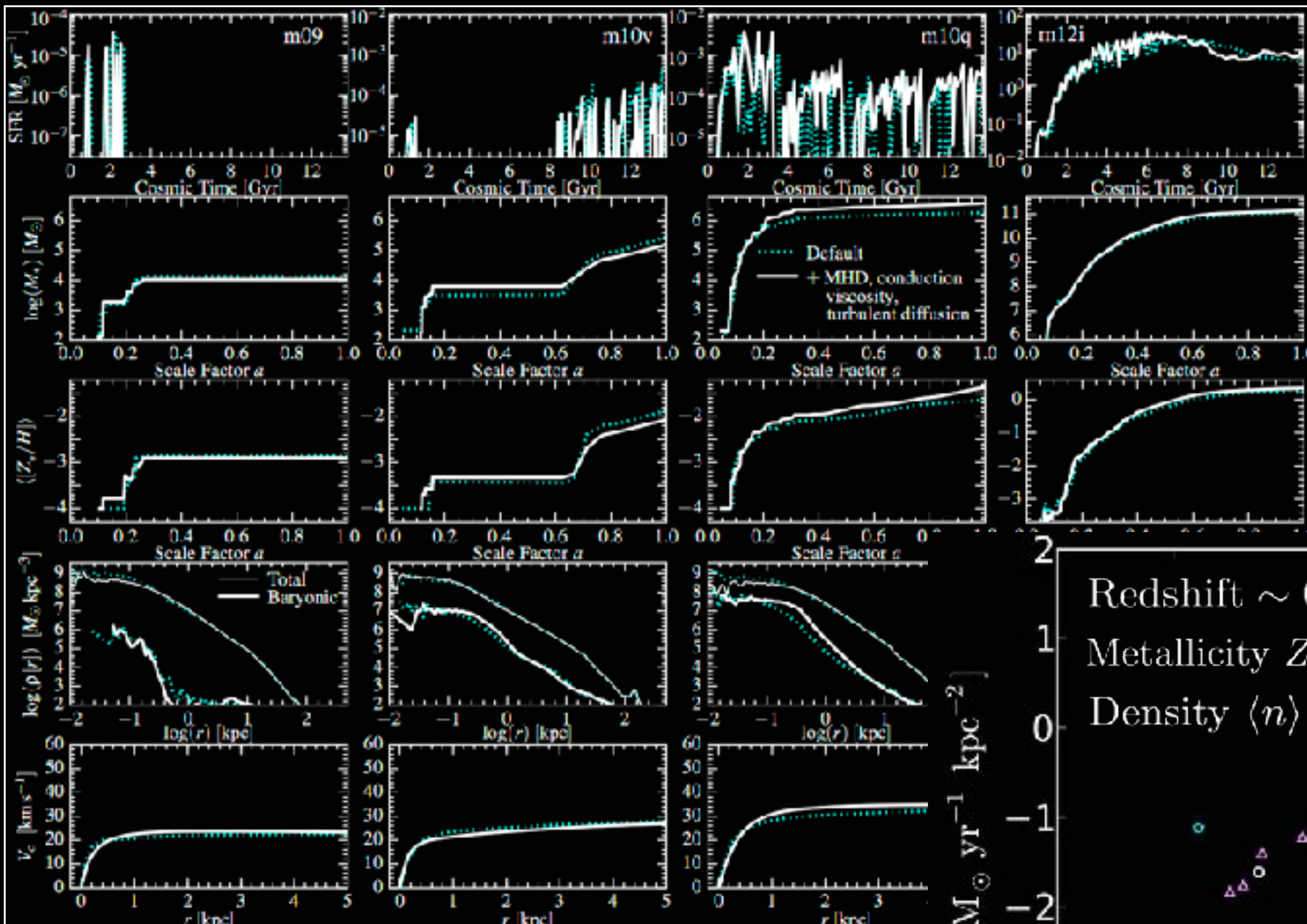
(molecular gas)

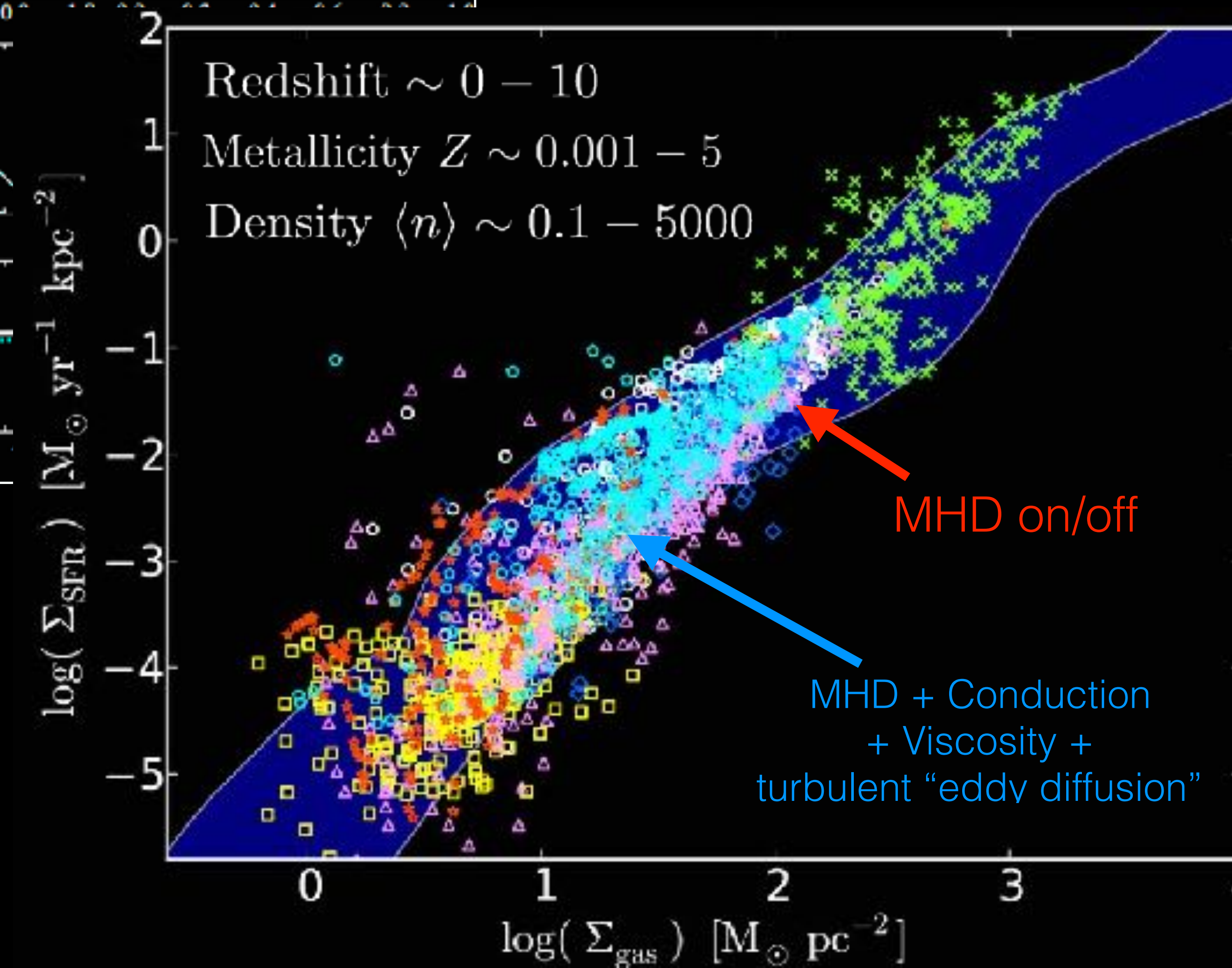
Matt Orr (in prep)
Hopkins+ 11,12,14
Shetty+ 14
Narayanan+ 13

Galaxy SFRs (sub-L*) independent of MHD+diffusion

MAY NOT APPLY TO COOLING IN HOT HALOS!

Kung-Yi Su
(in prep.)





Numerical Methods

(aka: why did we switch from SPH?)

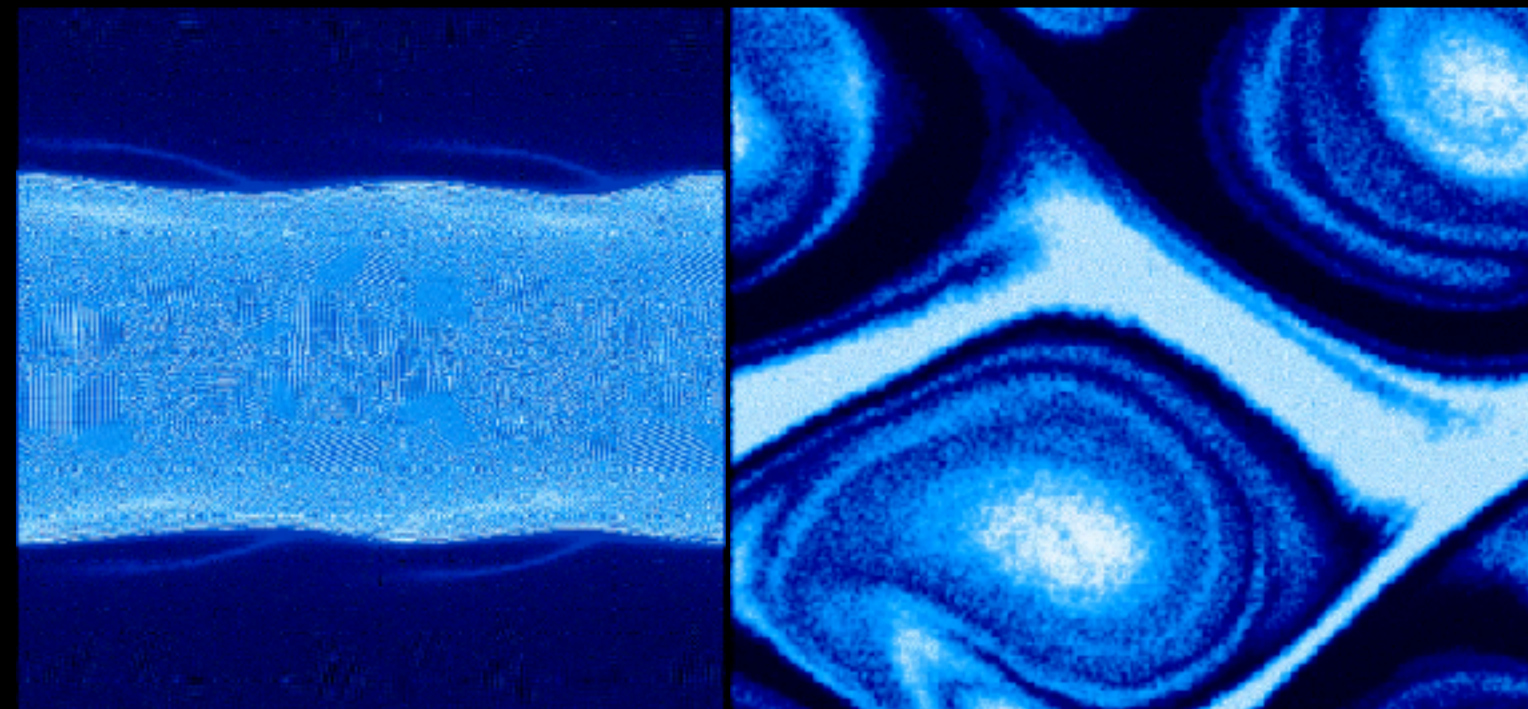
Smoothed-Particle Hydrodynamics (SPH)

CHALLENGE: POPULAR METHODS HAVE PROBLEMS

- Lagrangian, adaptive, simple, conservative
- Artificial diffusion terms:
    - excess diffusion, viscosity

Ritchie & Thomas 01, Agertz 07, Price 12, Read 12

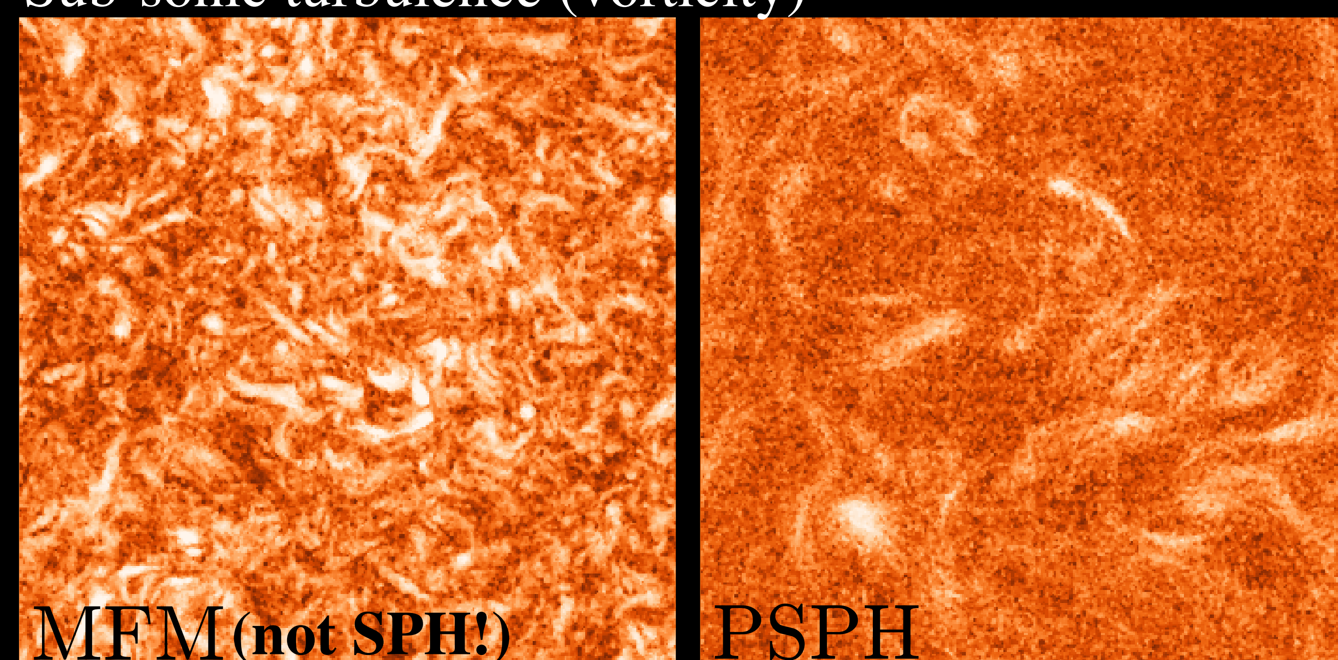
Kelvin-Helmholtz Instabilities



"old" SPH (Springel 02)

"new" SPH (PSPH) (Hopkins '13): >>100 neighbors

Sub-sonic turbulence (vorticity)

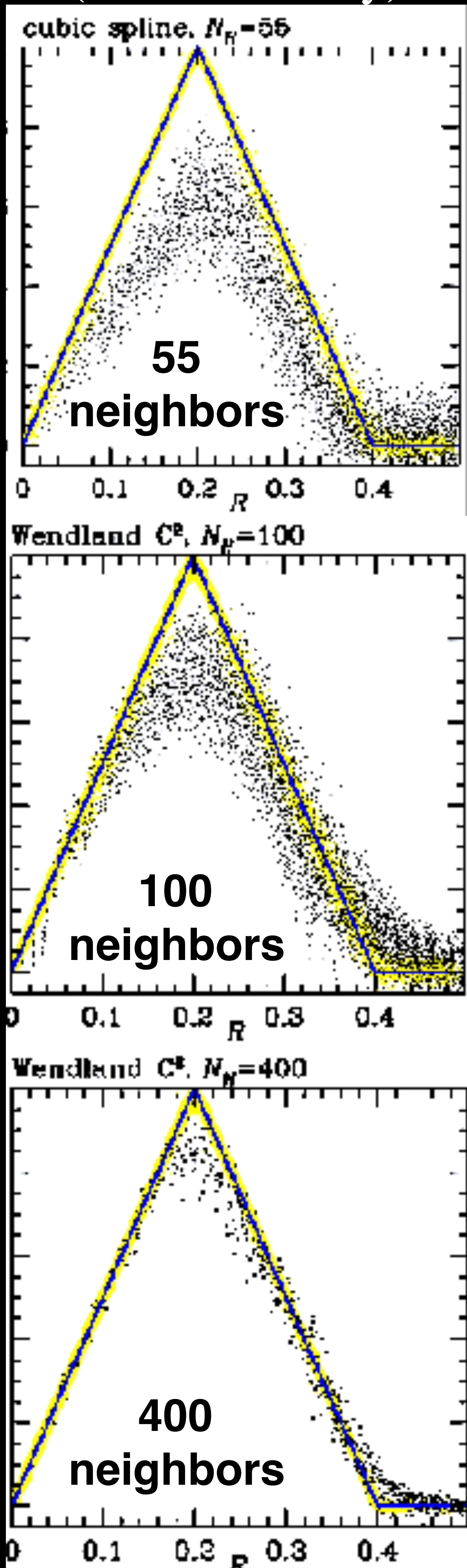


Smoothed-Particle Hydrodynamics (SPH)

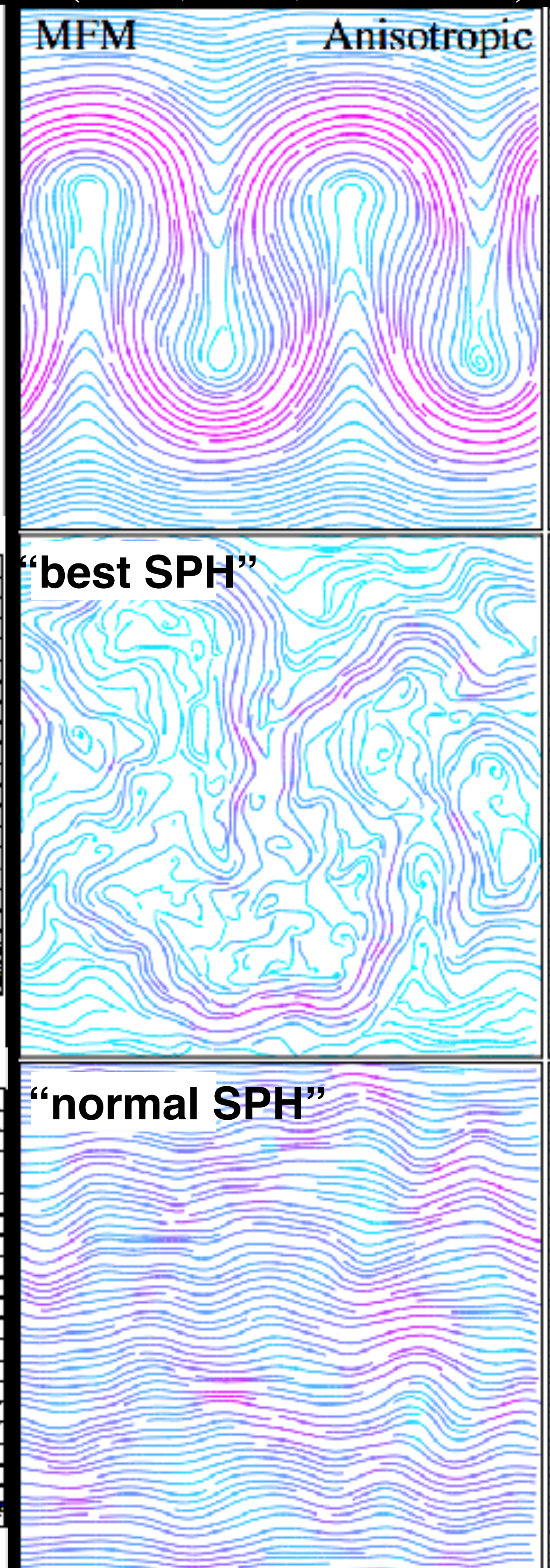
CHALLENGE: POPULAR METHODS HAVE PROBLEMS

- Fundamental low-order errors:
  - converges slowly:
    "beat down" by
    increasing kernel size,
    but this is *not efficient!*

Gresho vortex
(Dehnen & Aly)



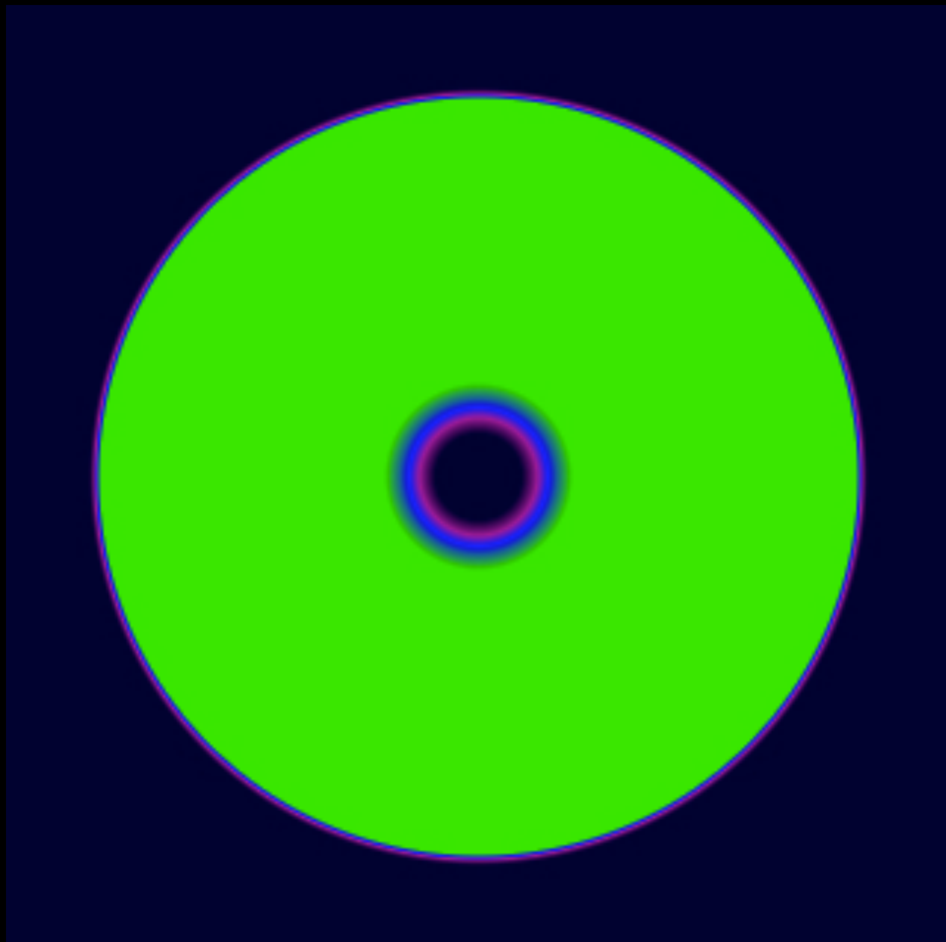
Anisotropic Conduction
(MTI, HBI, Hall MRI)

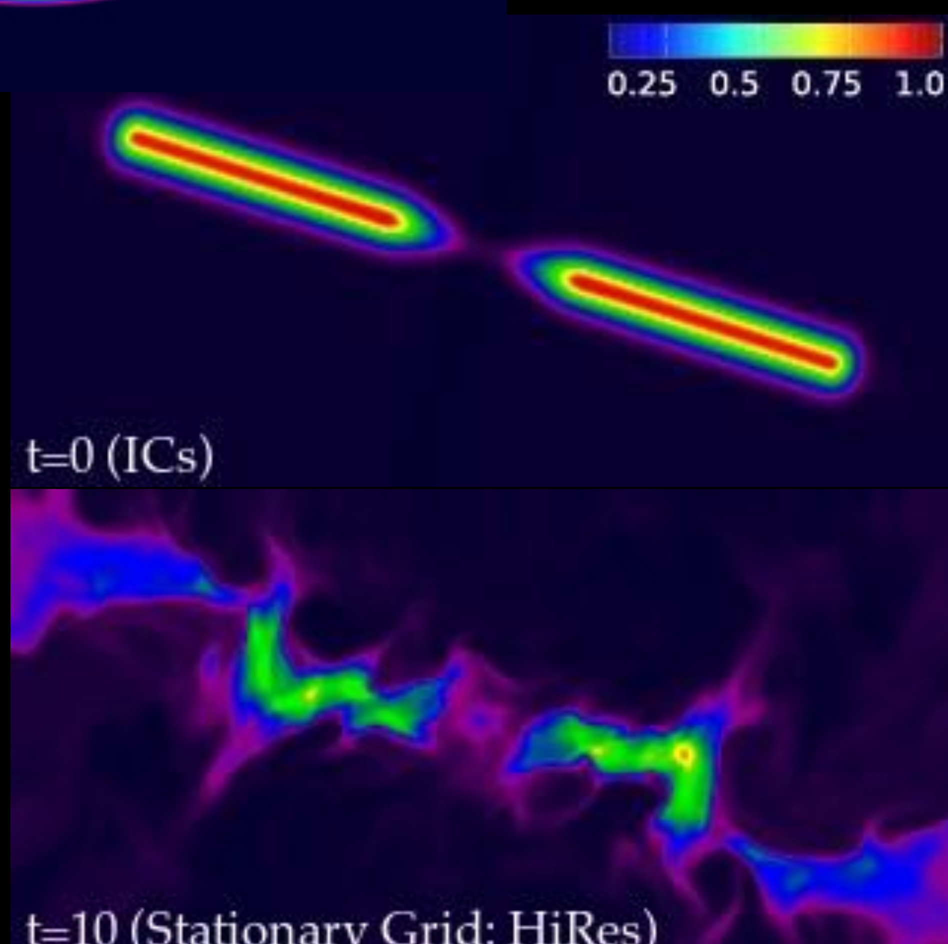


Dehnen & Aly, Rosswog, Hopkins, Tricco & Price, Read

Adaptive Mesh Refinement (AMR)

CHALLENGE: POPULAR METHODS HAVE PROBLEMS



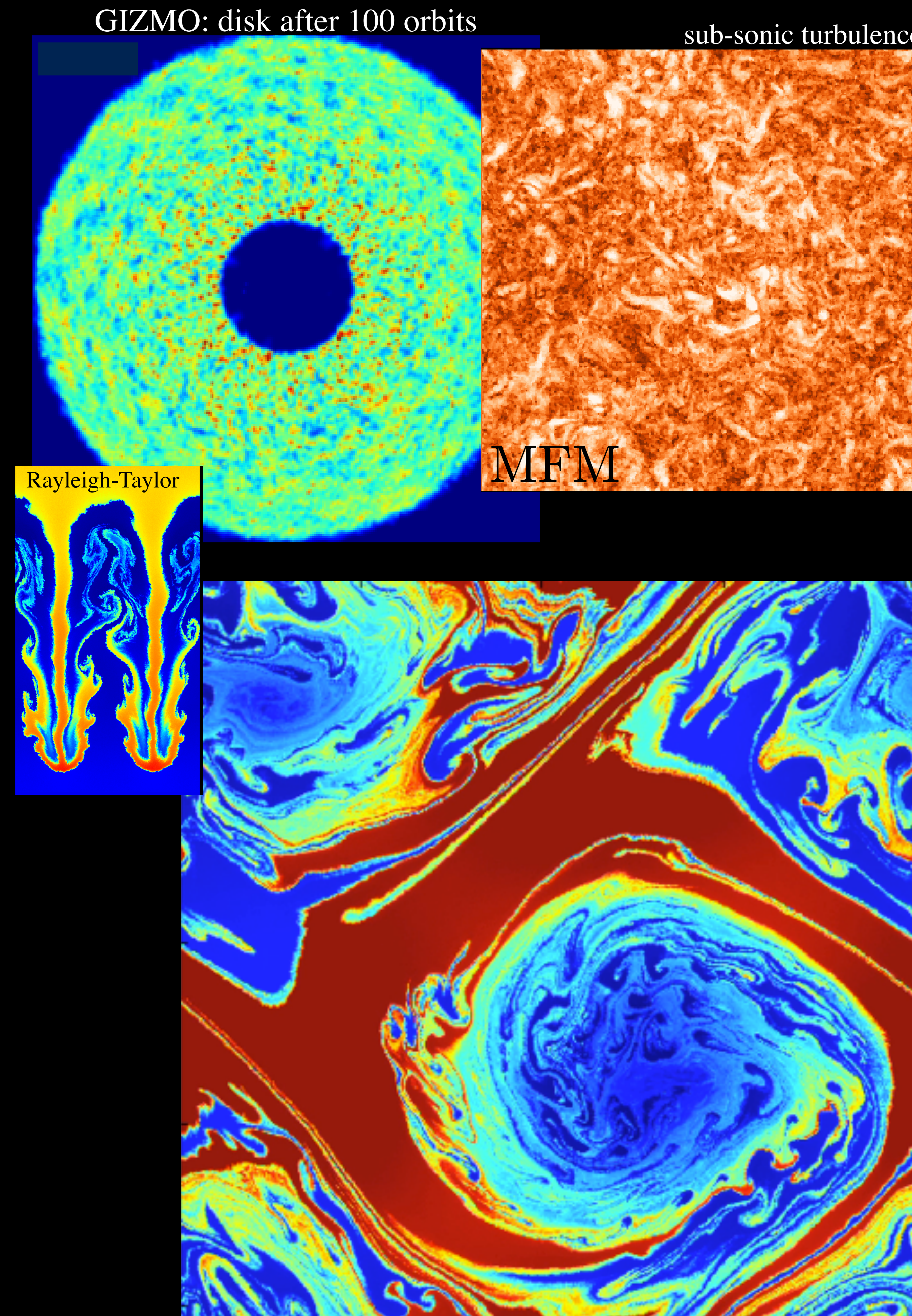


- Eulerian, well-studied, high-order
- Excessive mixing/diffusion
  when fluid moves over cells
- Geometric effects:
  - carbuncle instability (shocks)
  - loss of angular momentum
  - grid-alignment (disks)
- Also “beaten down” with resolution, but *expensive*
  - Hahn ’10: >>512^2 resolution to avoid grid-alignment

Peery & Imlay 88,
Mueller & Steinmetz 95, Hahn 10

New Methods Combine (some) Advantages of Both:

(BUT REMAIN LESS WELL-TESTED)



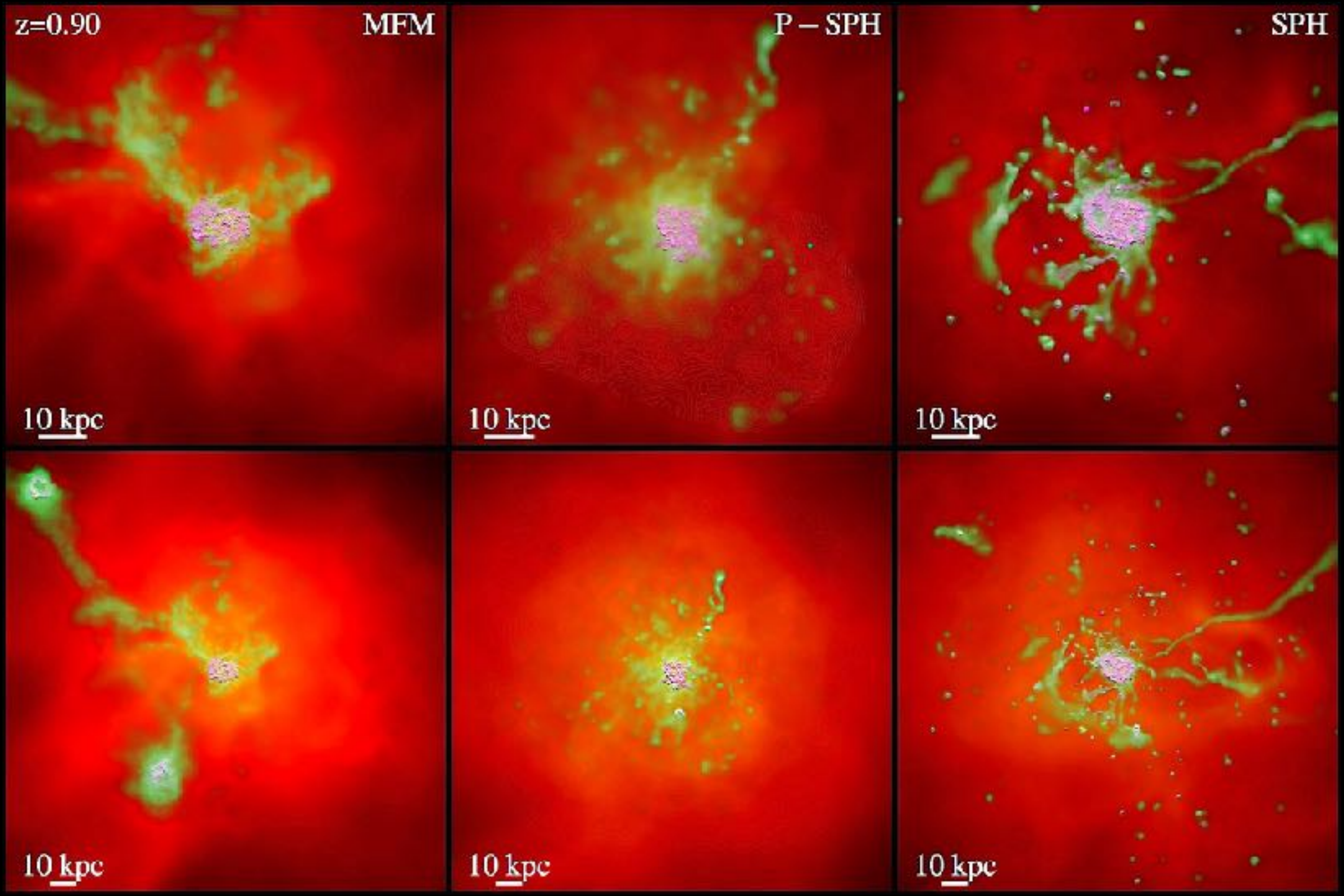
- Moving-meshes (AREPO), meshless finite-volume (GIZMO), high-order ALE methods
- Move with flow, no preferred geometry, but also accurate, high-order, and shock-capturing
- Grid noise is more severe

AREPO: Springel 2010
TESS/DISCO: Duffel 2011
FVMHD3D: Gaburov 2012
GIZMO: Hopkins 2015

Getting the Hydro Right Can Matter

BUT IT DEPENDS ON WHAT YOU CARE ABOUT

New Methods (GIZMO) | "New" SPH | "Old" SPH



Agertz 07 & many others

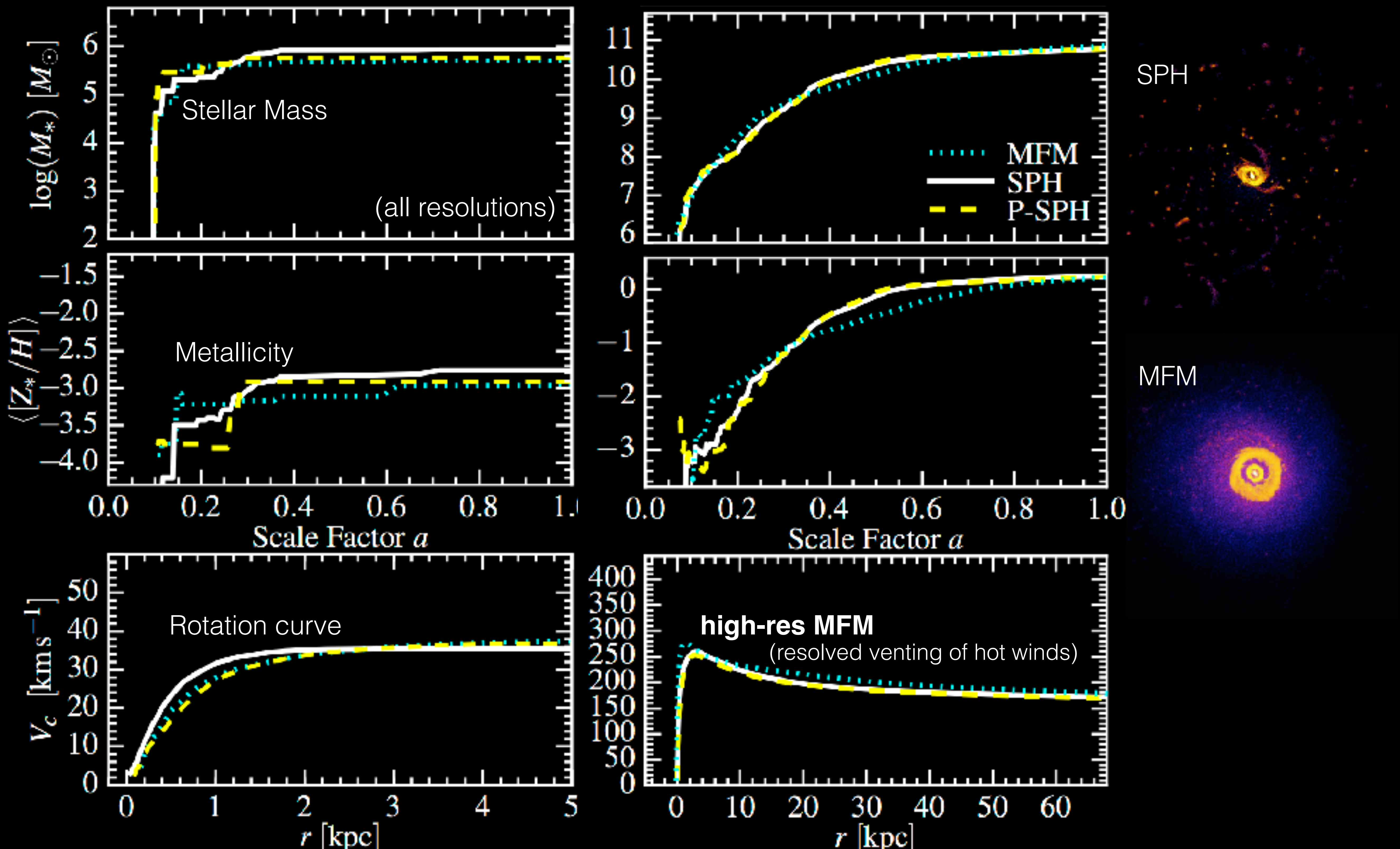
BUT only factor ~1.5 difference in mass!

Getting the Hydro Right Can Matter

DEPENDS ON WHAT YOU CARE ABOUT

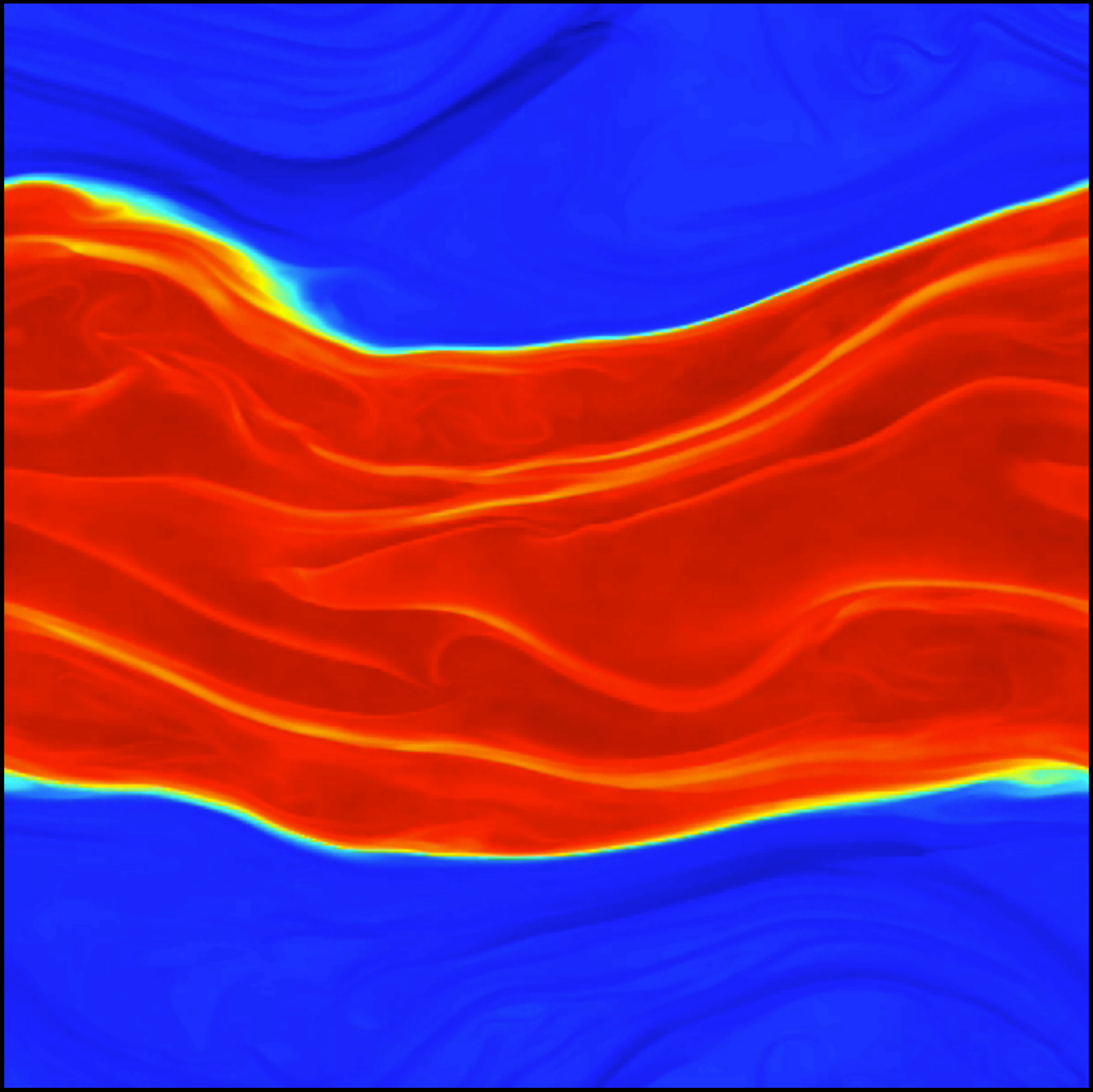
Dwarfs ("cold mode"): no effects

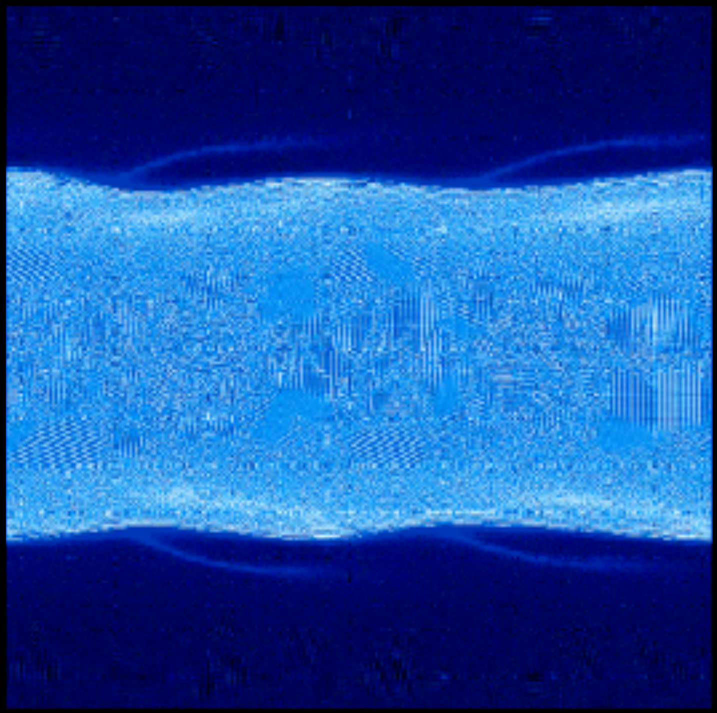
Massive Galaxies ("hot mode"): cooling & wind "venting"



A Caution: You can get the “right” answer for the wrong reasons

DON’T MISTAKE NUMERICAL PRECISION FOR PHYSICAL ACCURACY



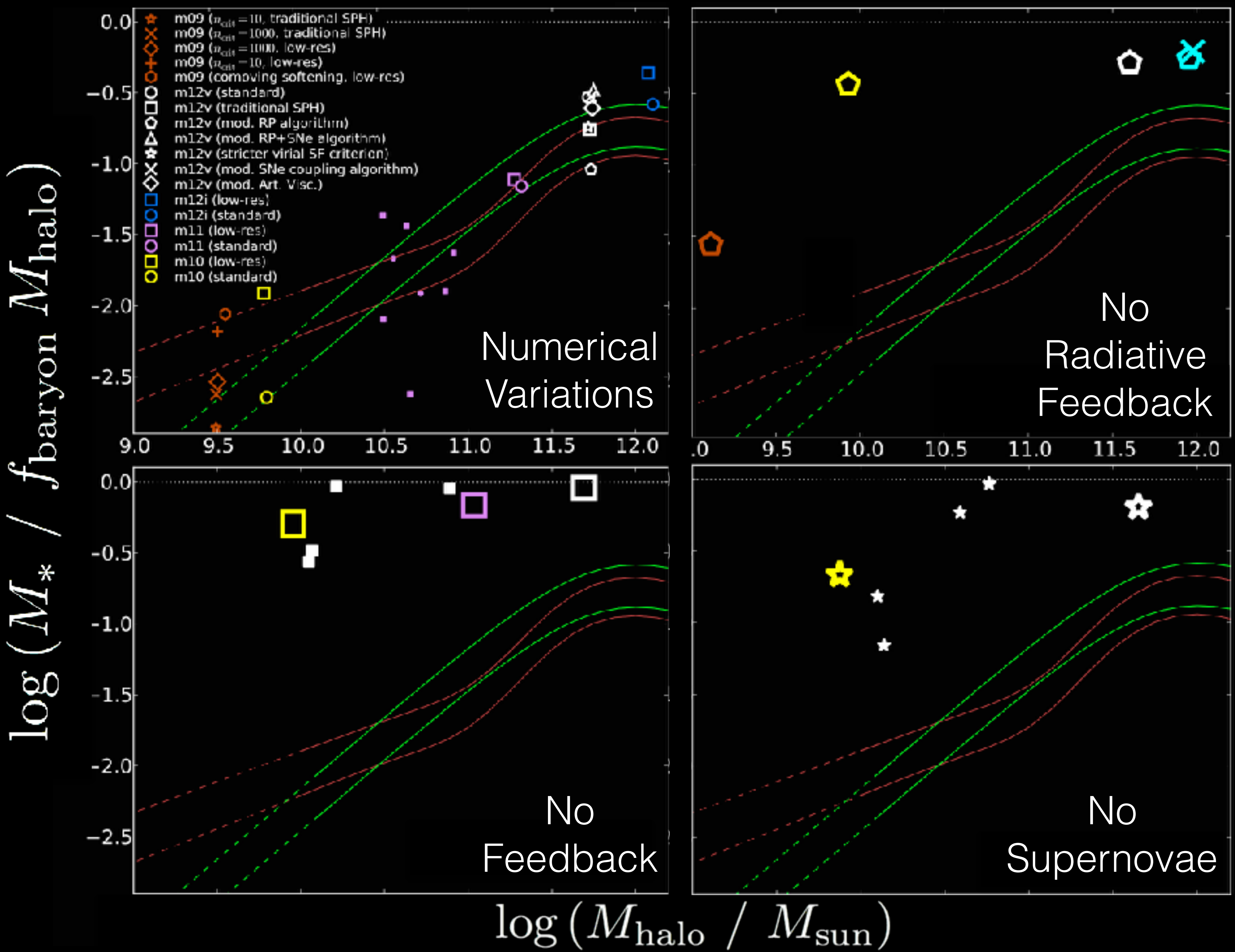


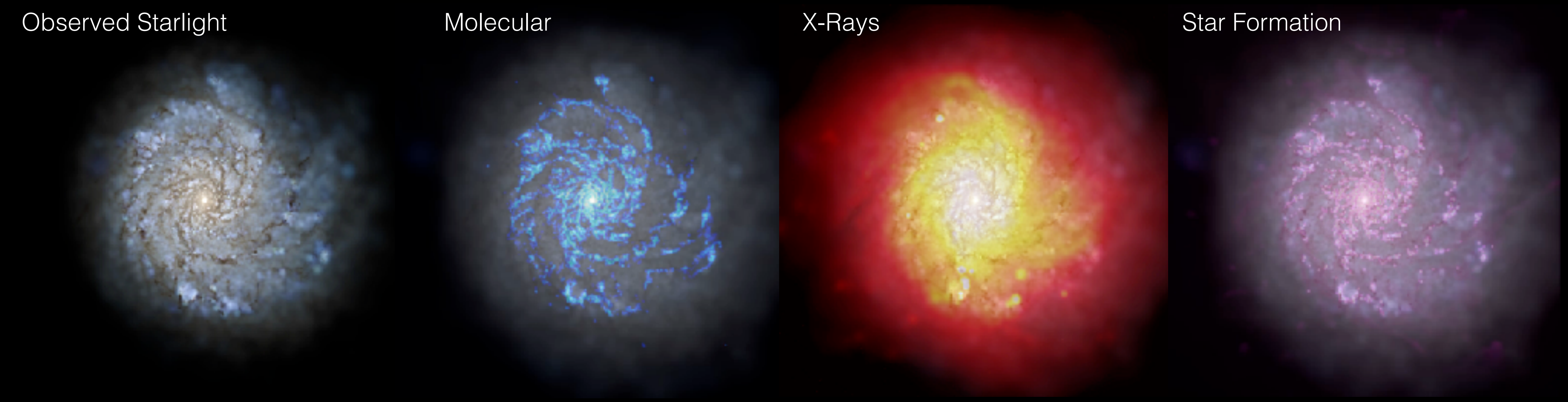
“old” SPH
(Springel 02)

Magnetic KH
(Equipartition field)
with a “good” code

Getting the Hydro Right Can Matter

DEPENDS ON WHAT YOU CARE ABOUT





- **Numerics can be important**
  - SPH: is high N_{NGB} worth it? MHD, conduction, RT, issues: significant differences in “hot halos”
  - Quasi-Lagrangian schemes: “grid noise” at very low Mach numbers (<0.01)
  - ***Physics*** usually dominates

- **Everything is sub-grid: but there are “good” and “bad” models, and different philosophies**
  - FIRE: trying to “build up” from small scales: works surprisingly well!
  - Need resolution to match your physics, but also need ***physics*** to match your resolution (no meaning in resolving scales you don’t have the physics for)

- **What is needed? Depends 100% on what you want to predict**
  - Resolve dense gas: resolve fragmentation (Toomre), ***physics*** for GMC destruction (radiative FB)
  - Resolve SNe overlaps/bubbles: need to treat them explicitly, account for unresolved cooling
  - SFR surprisingly insensitive to small-scale SF physics, MHD, diffusion: ***feedback*** dominates